TORUS2018 on December 14, 2018 in Puerto Varas, Chile

BASS Survey XI: The covering factor of dust and gas in Swift/BAT AGN

Ichikawa et al. '17, ApJ, 835, 74
Ichikawa et al. '18, ApJ in press.
arXiv:1811.02568

happy peaceful

Kohei Ichikawa (市川幸平)

FRIS fellow, Tohoku Univ. (from Oct 2018-)

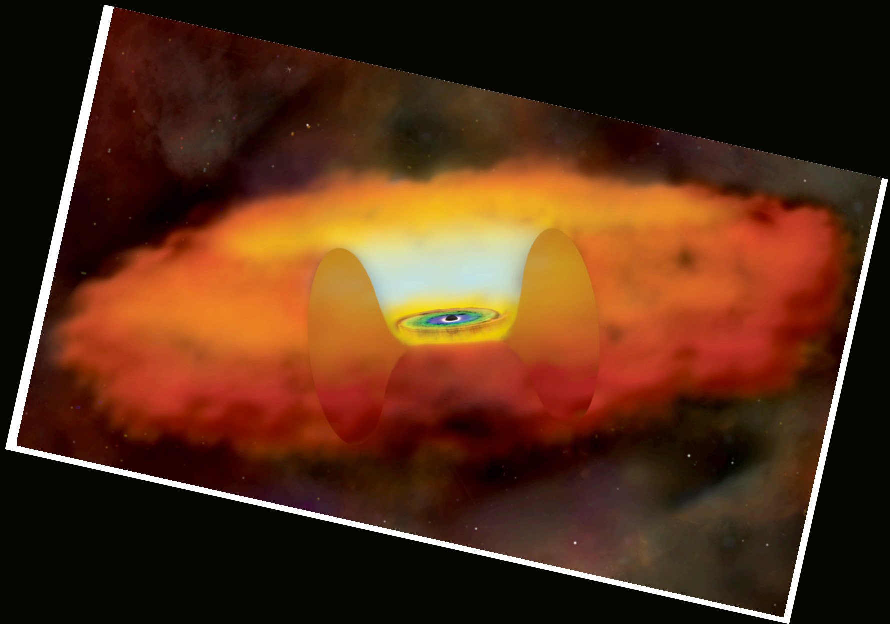
TOHOKU UNIVERSITY

In collaboration with

C. Ricci, Y. Ueda, **F. Bauer, T. Kawamuro**, M. J. Koss, K. Oh, D. J. Rosario, T. T. Shimizu, **M. Stalevski, L. Fuller**, **C. Packham**, B. Trakhtenbrot, and the **BASS team**

(Mid-)IR emission of AGN= nuclear dust

Nuclear (MIR) dust emitting region is compact w/ < 10pc



Urry & Padovani '95

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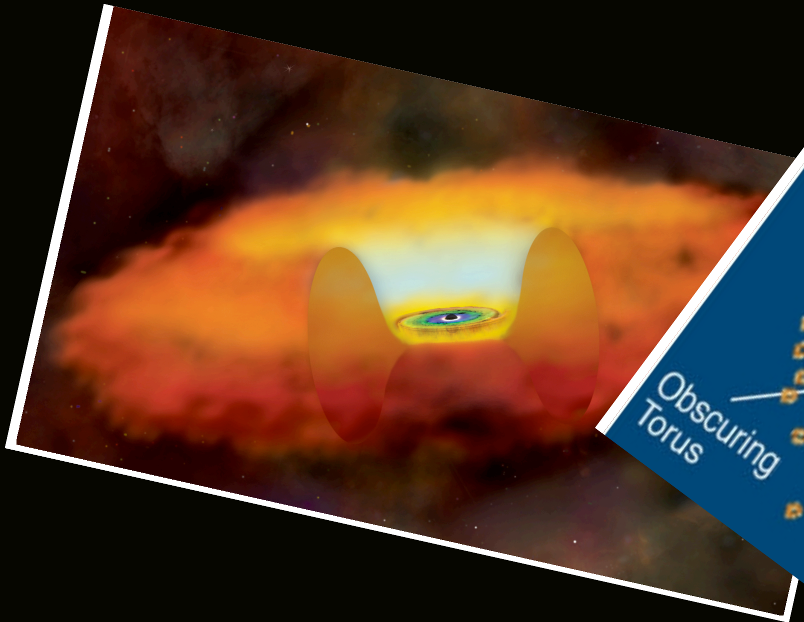

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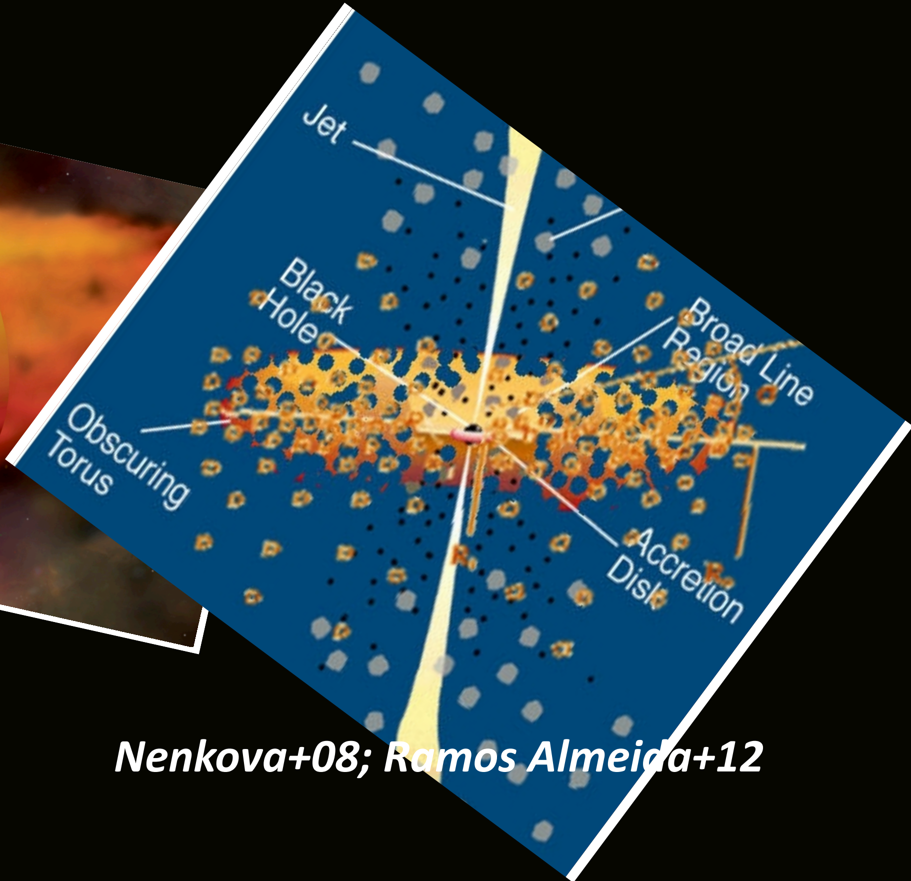

Nenkova+08; Ramos Almeida+12

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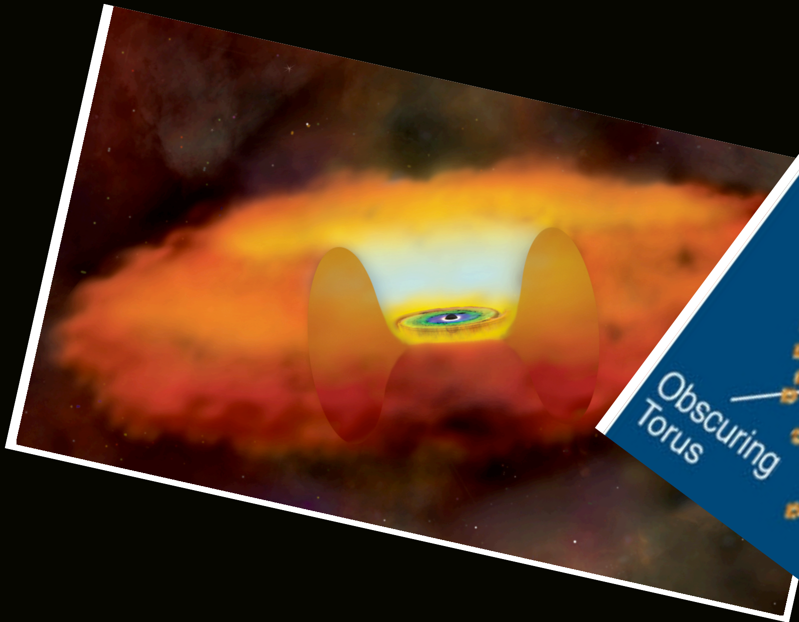

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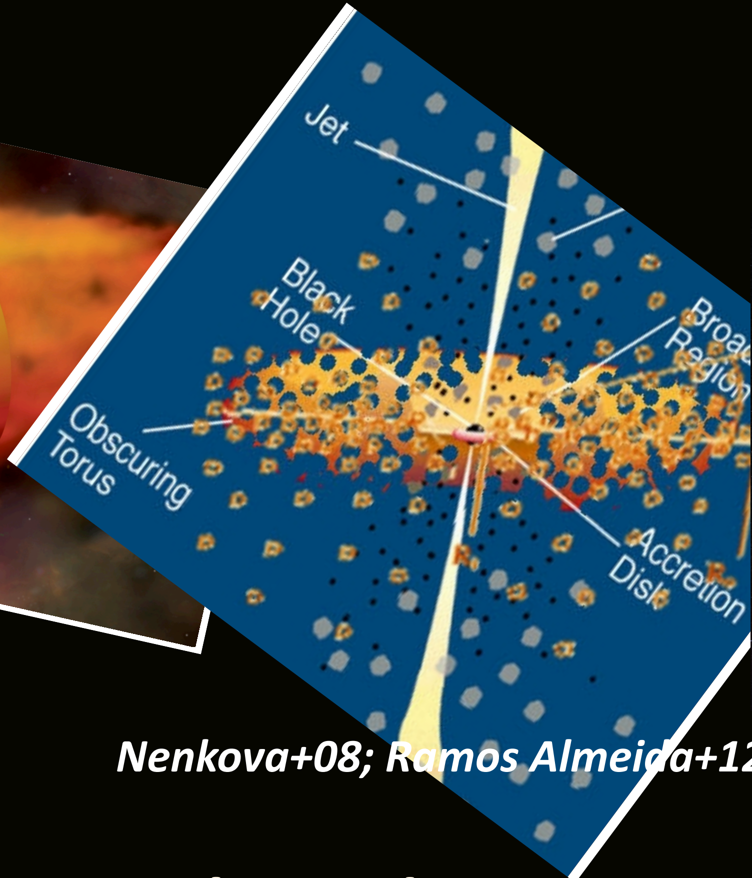


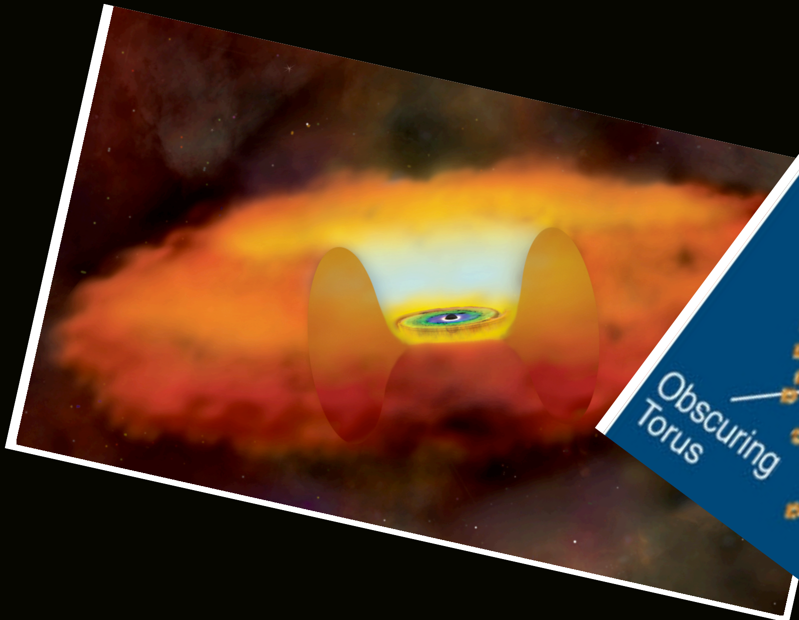
Nenkova+08; Ramos Almeida+12

e.g., Hoenig+12

∴ **Resolving the regions is limited to < 10 sources**

Geometry of (nuclear) dust emission

Nuclear (MIR) dust emission is compact w/ < 10pc



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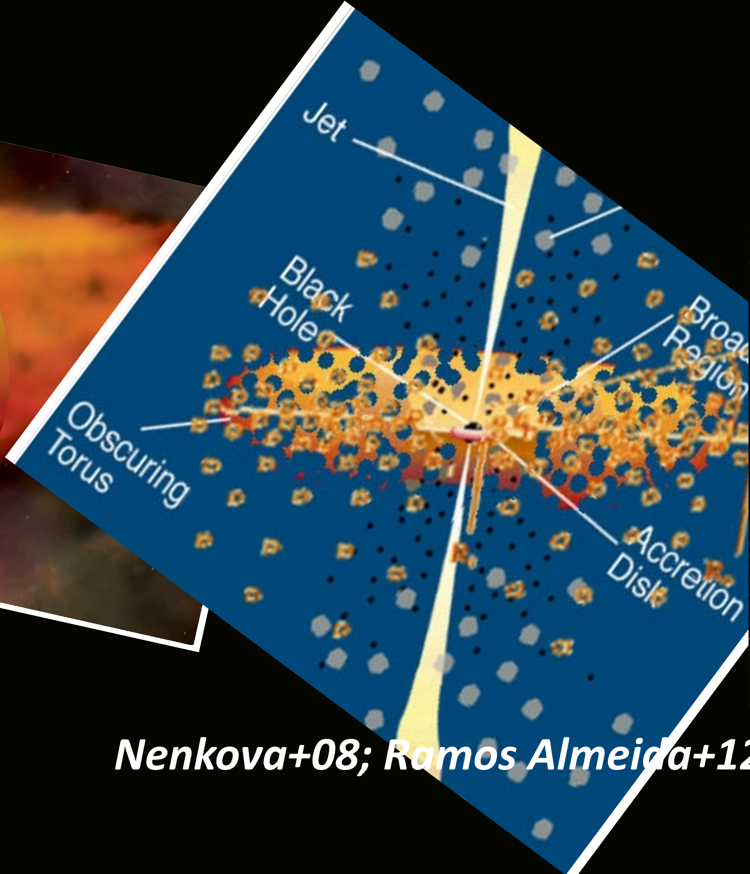


Nenkova+08; Ramos Almeida+12

e.g., Hoenig+12

∴ Resolving the regions is limited to < 10 sources

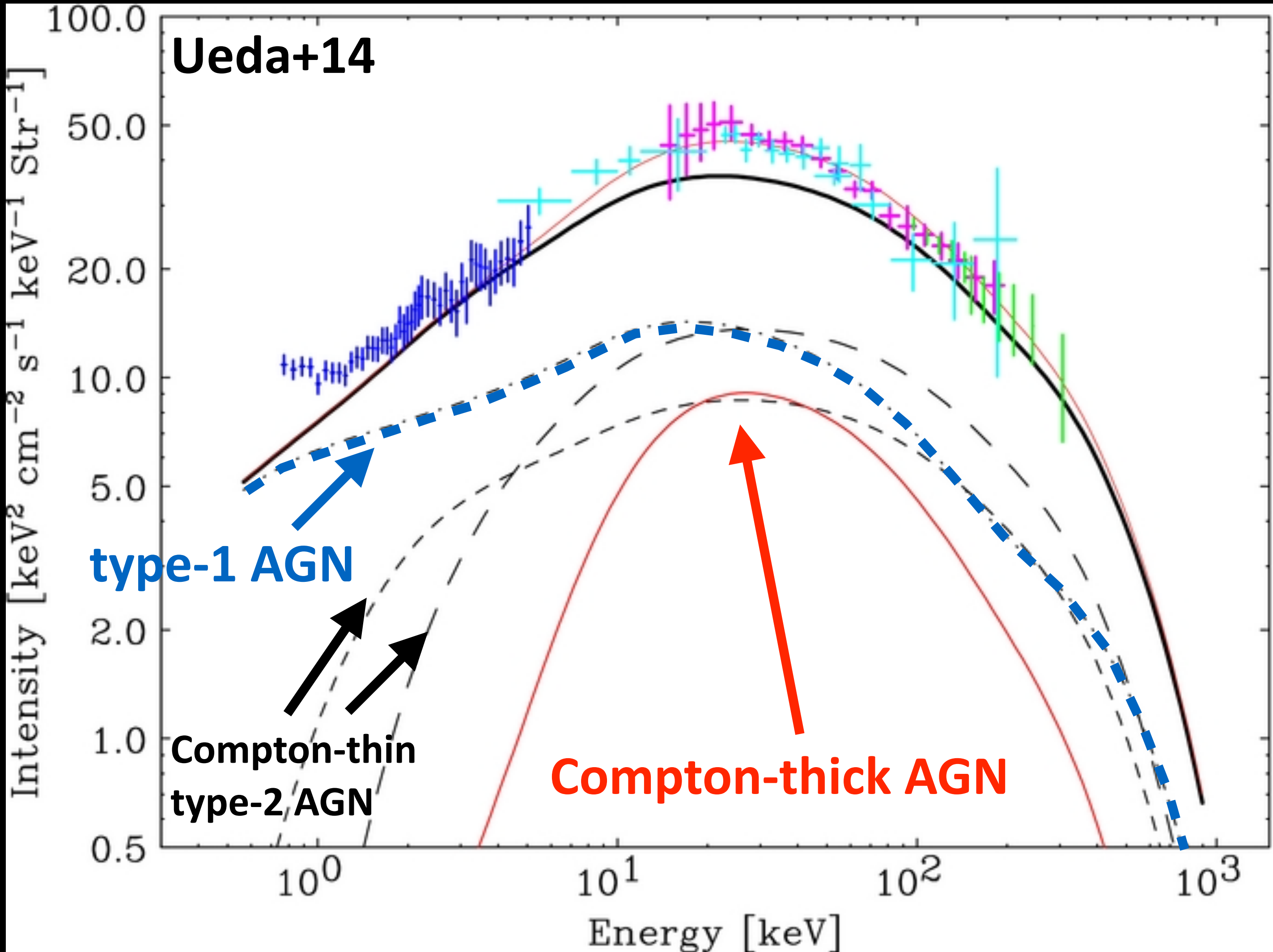
Q. How much do we constrain the (averaged) dust geometry?

$$C_{\mathrm{T}}(\mathrm{dust}) \propto L_{\mathrm{IR}}(\mathrm{AGN})/L_{\mathrm{bol}}(\mathrm{AGN})$$

Our Goal: Obtaining C_{T}(dust) using the complete AGN sample

Most of AGN are obscured

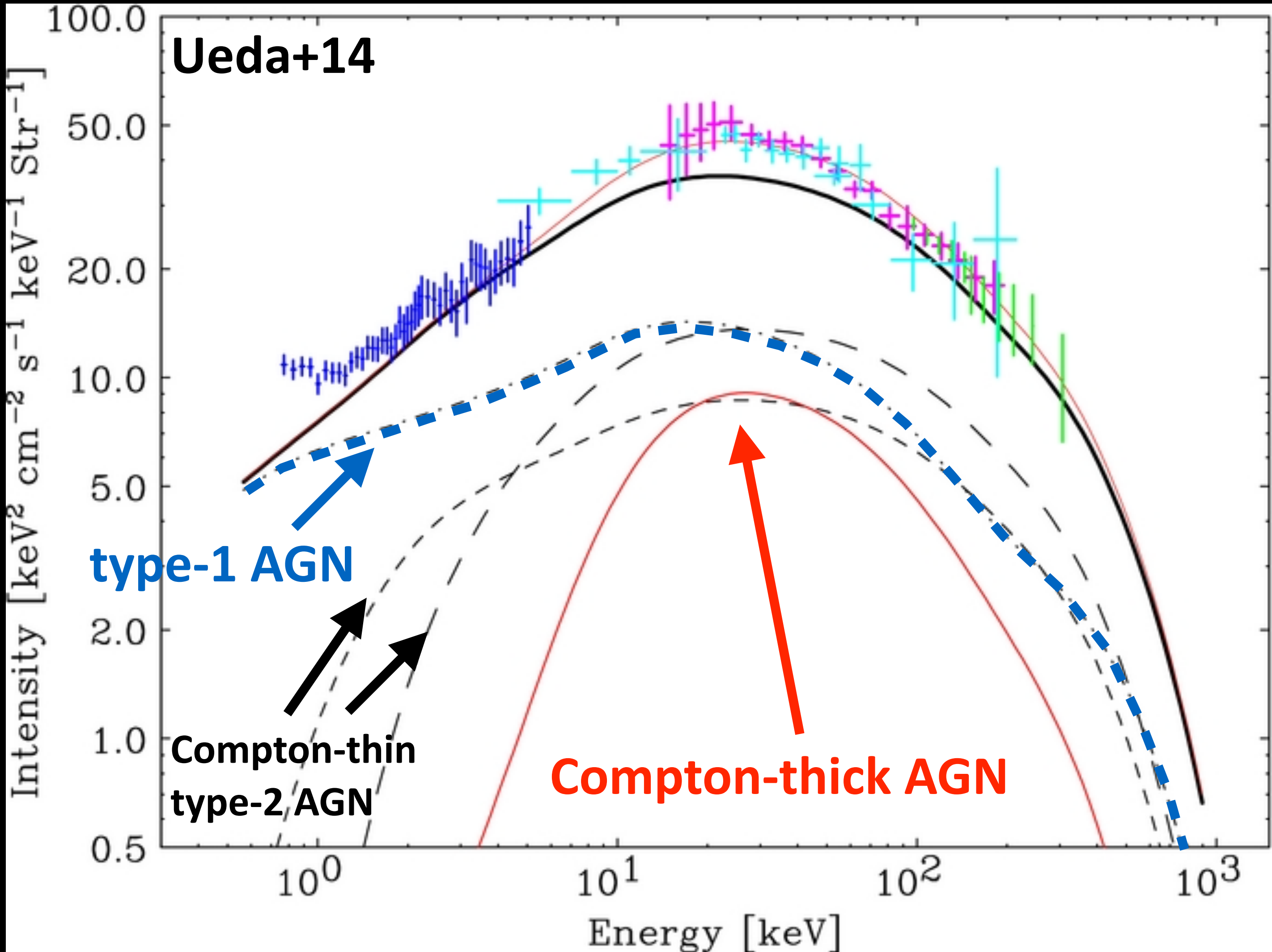
XRB indicates that most of AGN are obscured



☑ energy density peaks at **~30 keV**

Most of AGN are obscured

XRB indicates that most of AGN are obscured



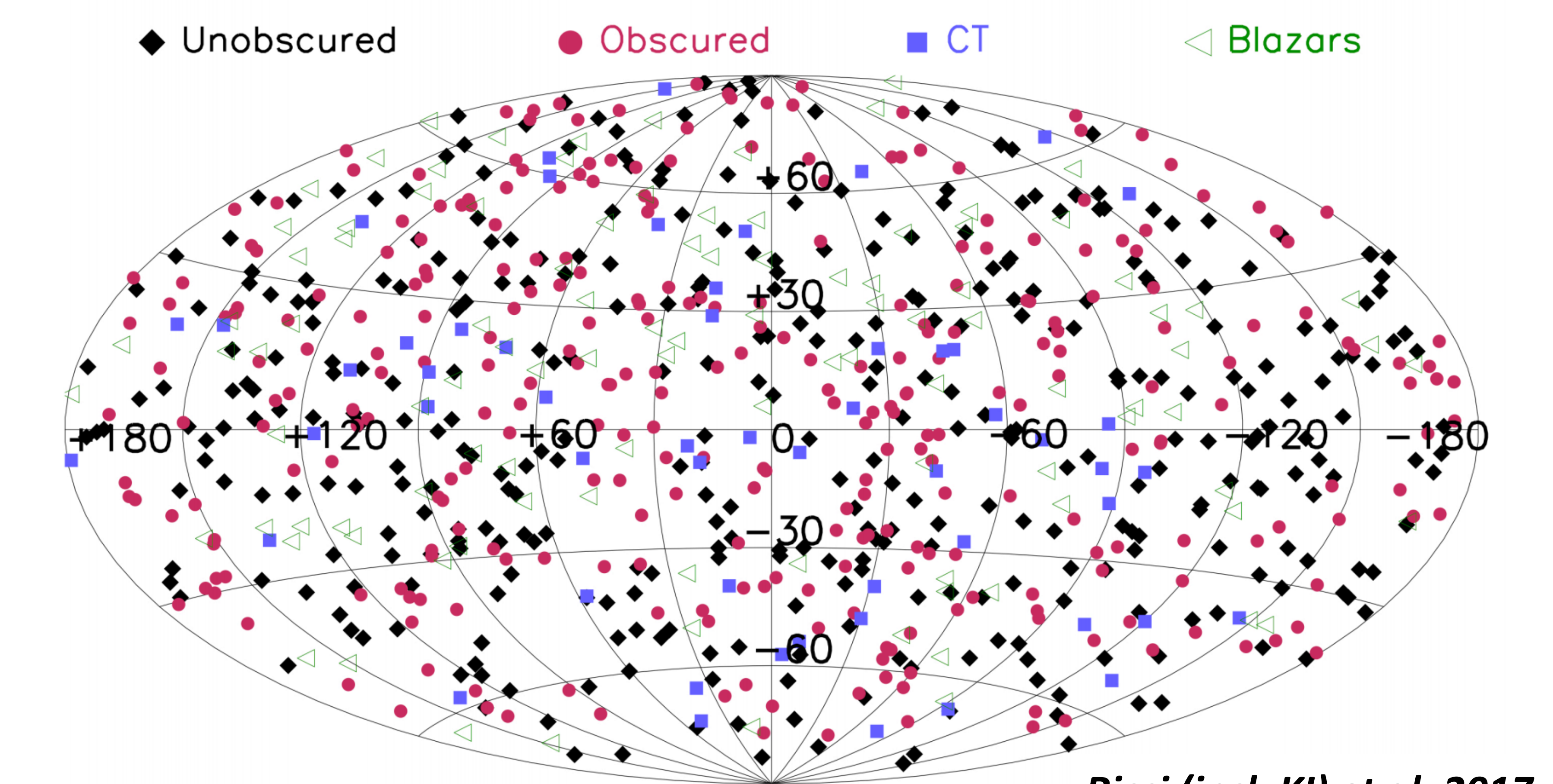
☑ energy density peaks at **~30 keV**

☑ **E>10 keV:** best energy band to detect obscured (log N_H>22) AGN

Swift/BAT AGN (14-195 keV)

70 month catalog: 836 AGN (728 non-blazars)

FYI, 105 month catalog is public (Oh et al., '18)



Ricci (incl. KI) et al. 2017

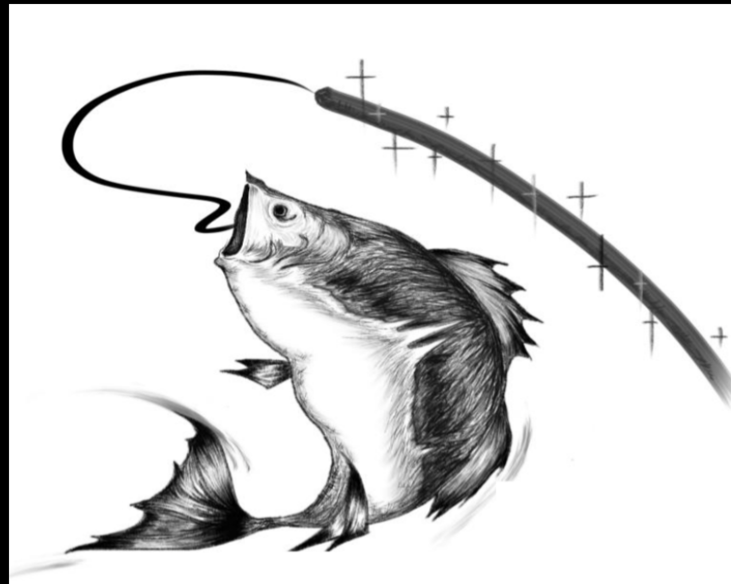
☑ most complete up to logN_H=24.5 in the local Universe *(Koss+16; Ricci+15)*

☑ **606 out of 728 have z info and are located at |b|>10°**

BASS=BAT AGN Spec Survey

Multi-wavelength Follow-up of BAT-AGN

co-lead by M. Koss, ***C. Ricci***, B. Trakhtenbrot, K. Oh



by Courtesy of K. Oh

- ☑ X-ray (L_X, N_{H}, Γ) Ricci et al. (2017)
- ☑ Optical Spec (M_{BH}, λ_{Edd}) Koss et al. (2017)
- ☑ NIR Spec (σ, M_{BH}) Lamperti et al. (2017)

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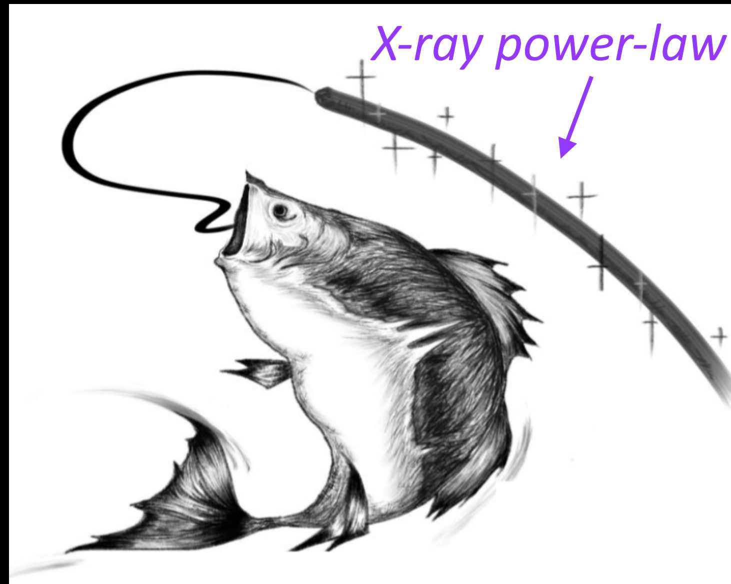

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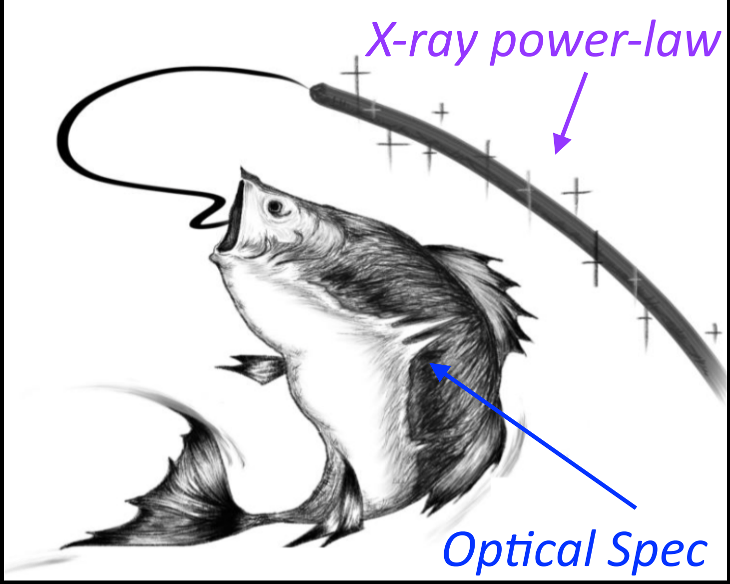

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- ☑ Optical Spec (M_{BH}, λ_{Edd}) Koss et al. (2017)
- ☑ NIR Spec (σ, M_{BH}) Lamperti et al. (2017)

More studies and Data, see BASS website!

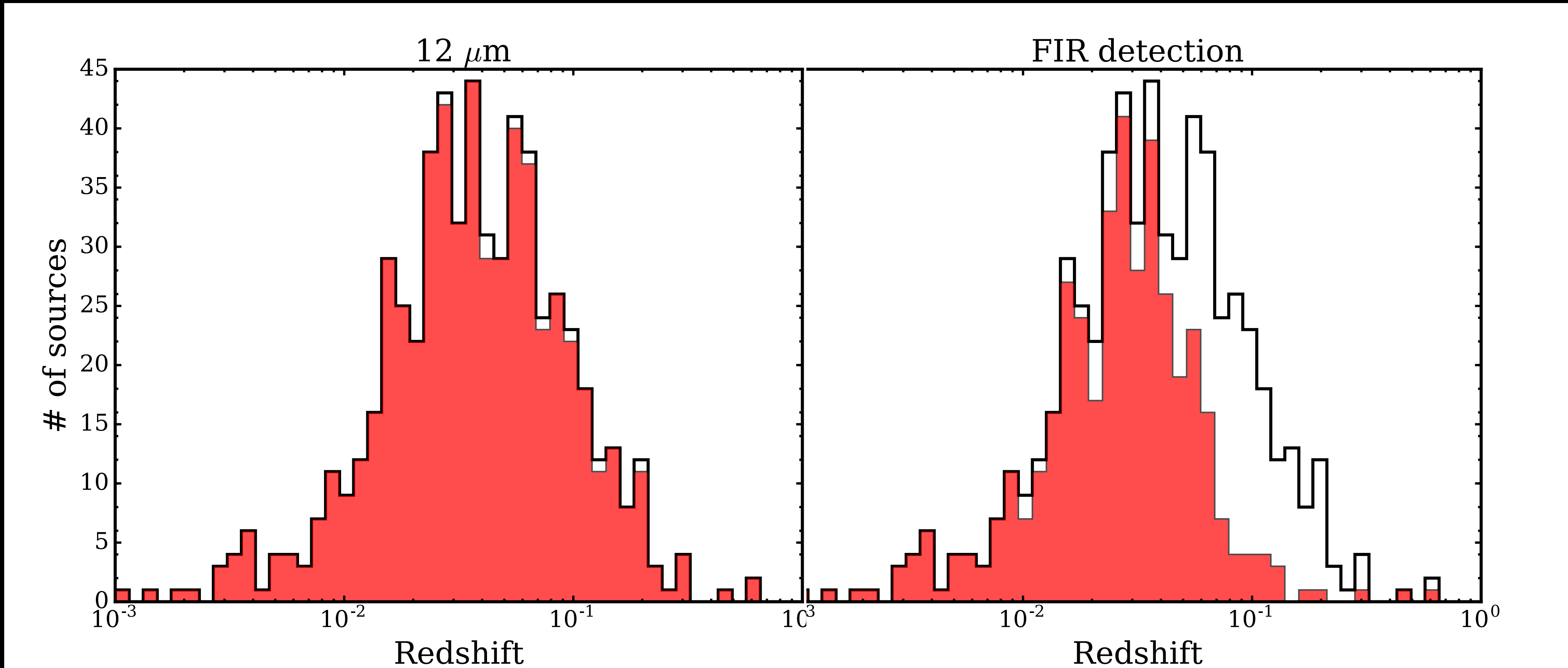
Today's topic

- ☑ IR catalog (3-500 um) ***Ichikawa et al. (2017a)***
- ☑ IR SED Decomposition; ***Ichikawa et al. (2018), arXiv:1811.02568***

IR counterparts of BAT AGN

☑ **3-500 um IR data from WISE, AKARI, IRAS, and Herschel**

(see Ichikawa+17 for more details)



☑ **601/606 MIR (, NIR) and 402/606 FIR counterparts**

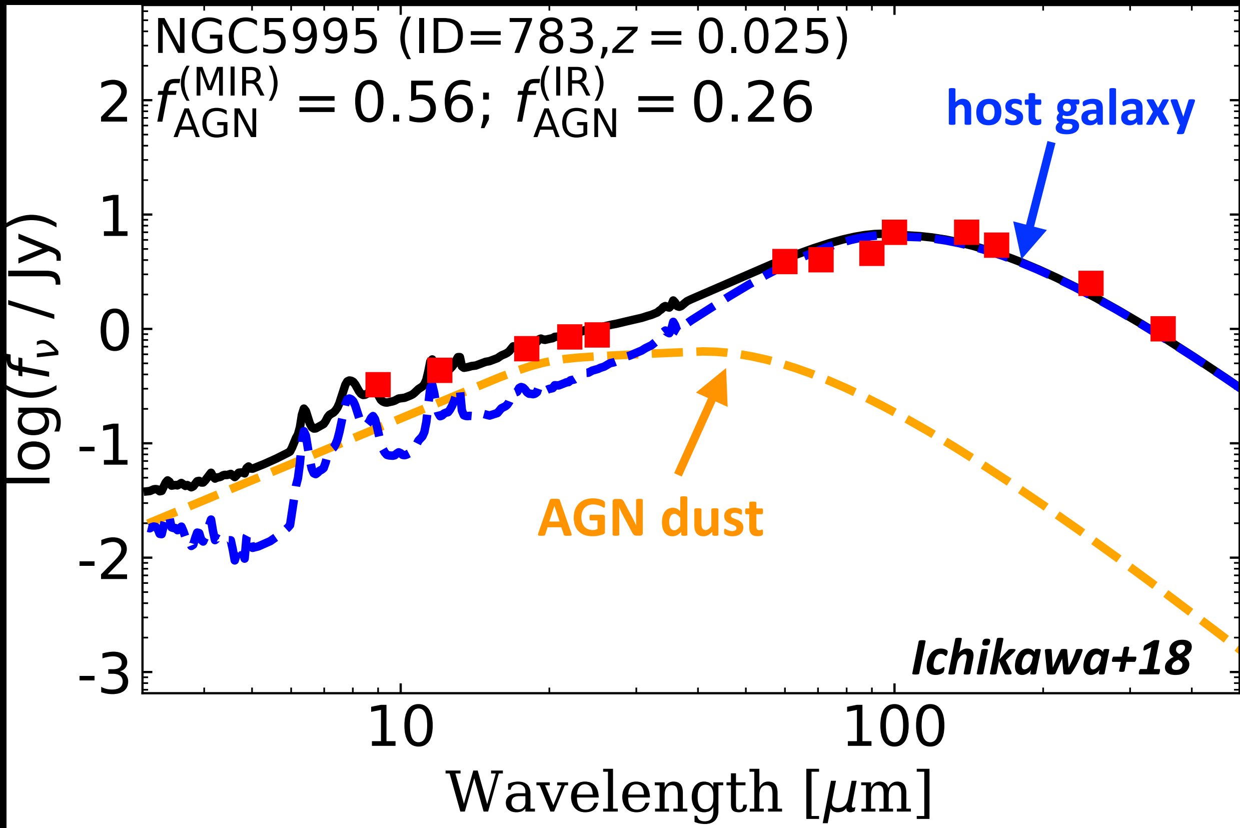
☑ **suitable for the AGN dust/host galaxy studies**

☑ **IR Data is already public.** http://iopscience.iop.org/0004-637X/835/1/74/suppdata/apjaa5154t1_mrt.txt

SED Decomposition in IR bands

☑ **SED Decomposition is done using simple AGN/(SB+stellar) templates**

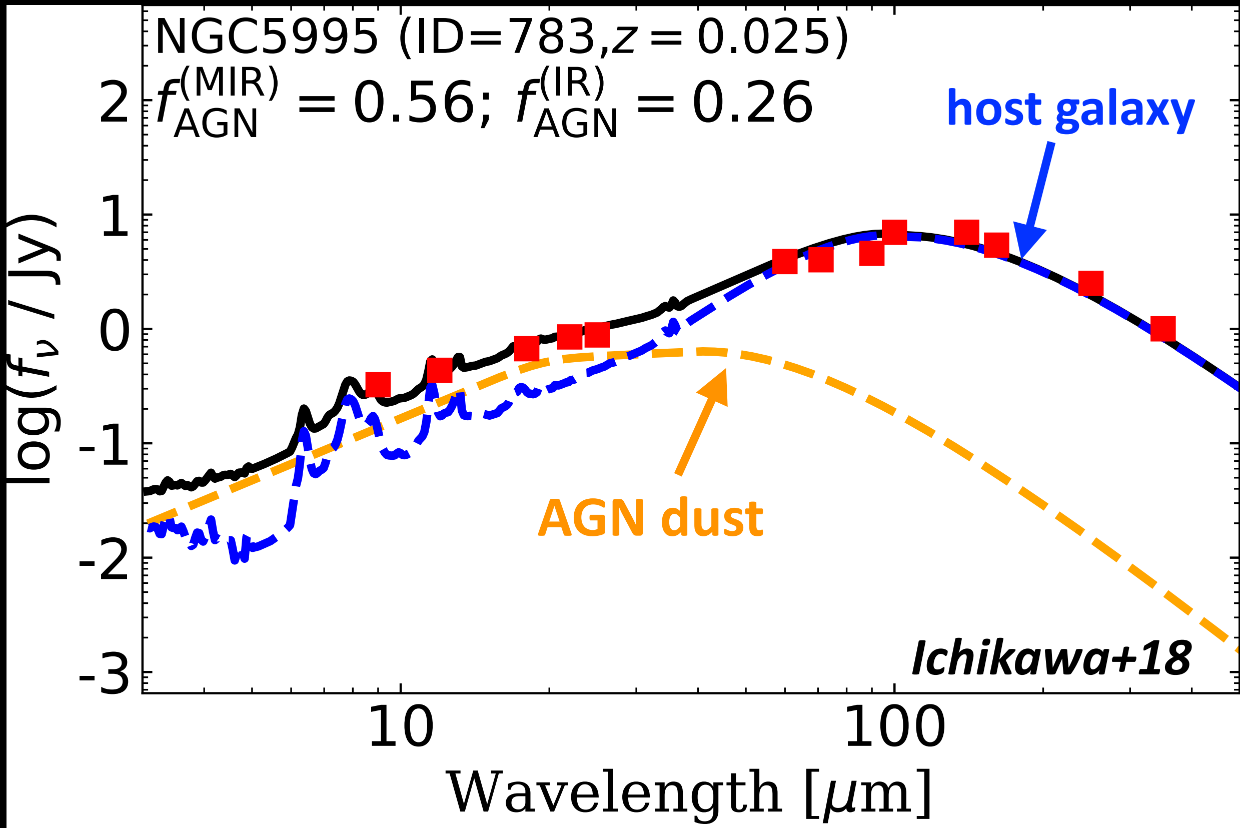
(see Ichikawa+18 for more details)



SED Decomposition in IR bands

☑ **SED Decomposition is done using simple AGN/(SB+stellar) templates**

(see Ichikawa+18 for more details)



☑ **SED decomposition: 587/606 sources**

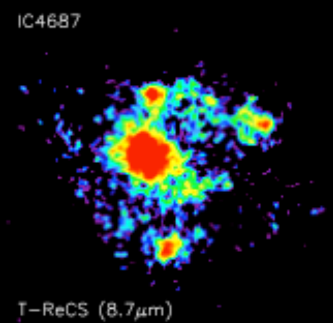
☑ **Disentangling AGN/host galaxy (SB+stellar) component**

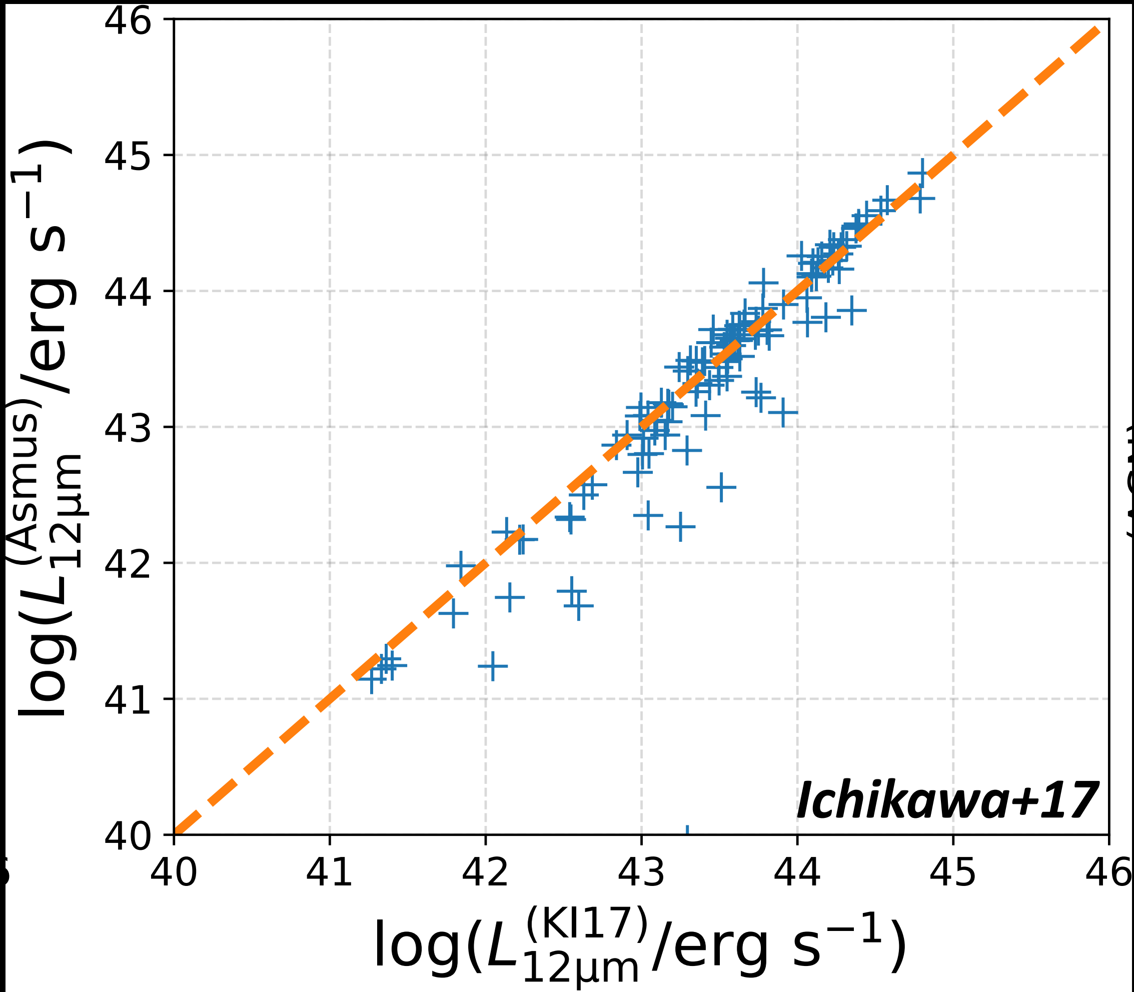
=> AGN IR emission w/o host galaxy contamination

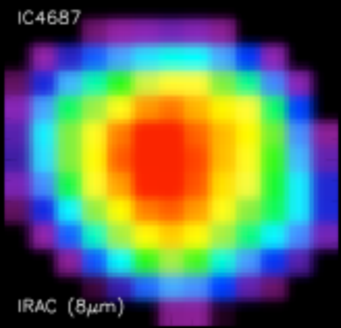
FYI, All info incl. IR SEDs, decomposed SEDs, M_{BH}, $L_{x, bol}$ will be public

Comparison with high-spatial resolution observations

High spatial. resol. obs.
(Asmus+14,+15)







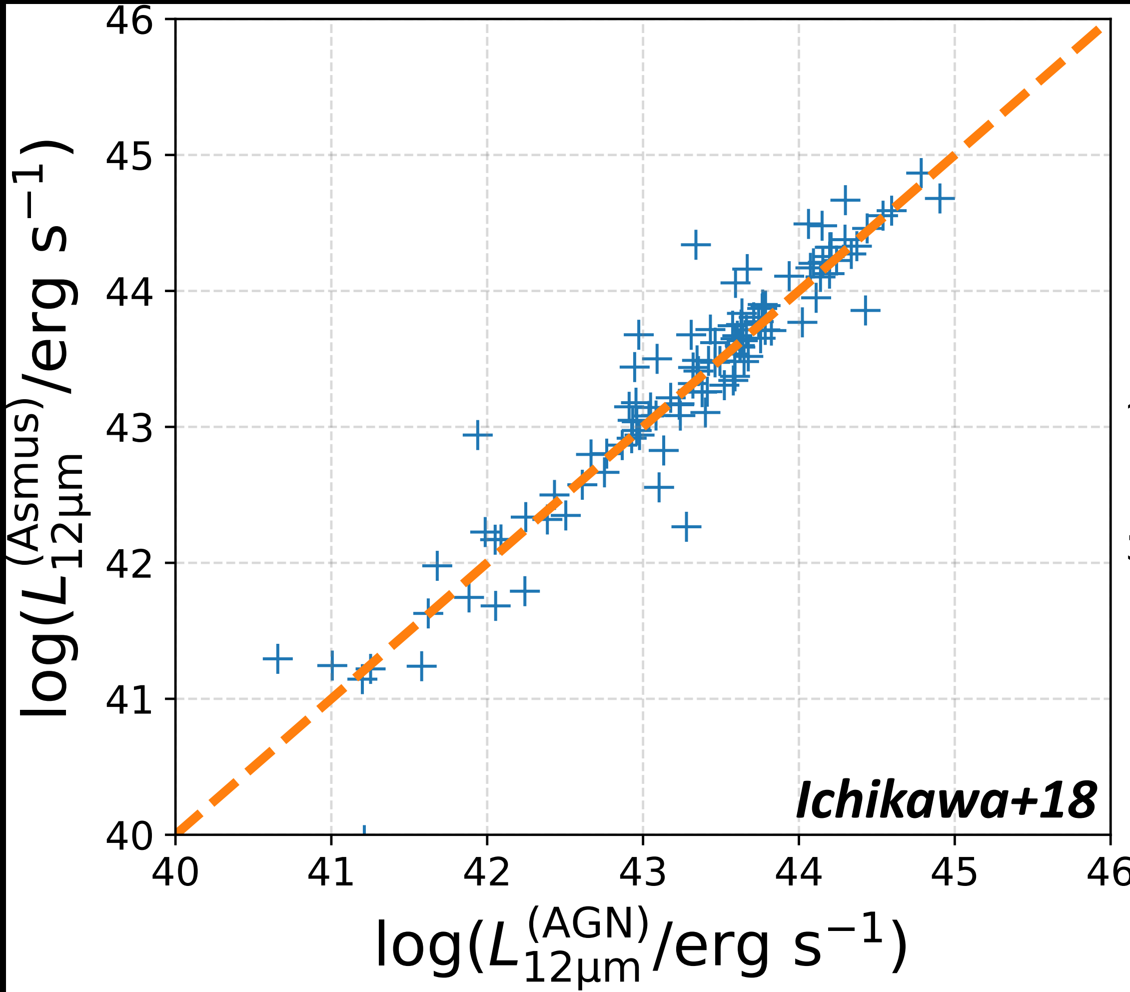
L_{12um} "Before" SED decomposition

☑ $L_{12um}^{(KI17)} >= L_{12um}^{(Asmus)}$

Comparison with high-spatial resolution observations

☑ SED Decomposition works well!

High spatial. resol. obs.



L_{12um} *"after"* SED decomposition

☑ SED decomposition reproduces L_{12um} of 0."3-0."7 scale high spatial resolution observations (Asmus+14;15)

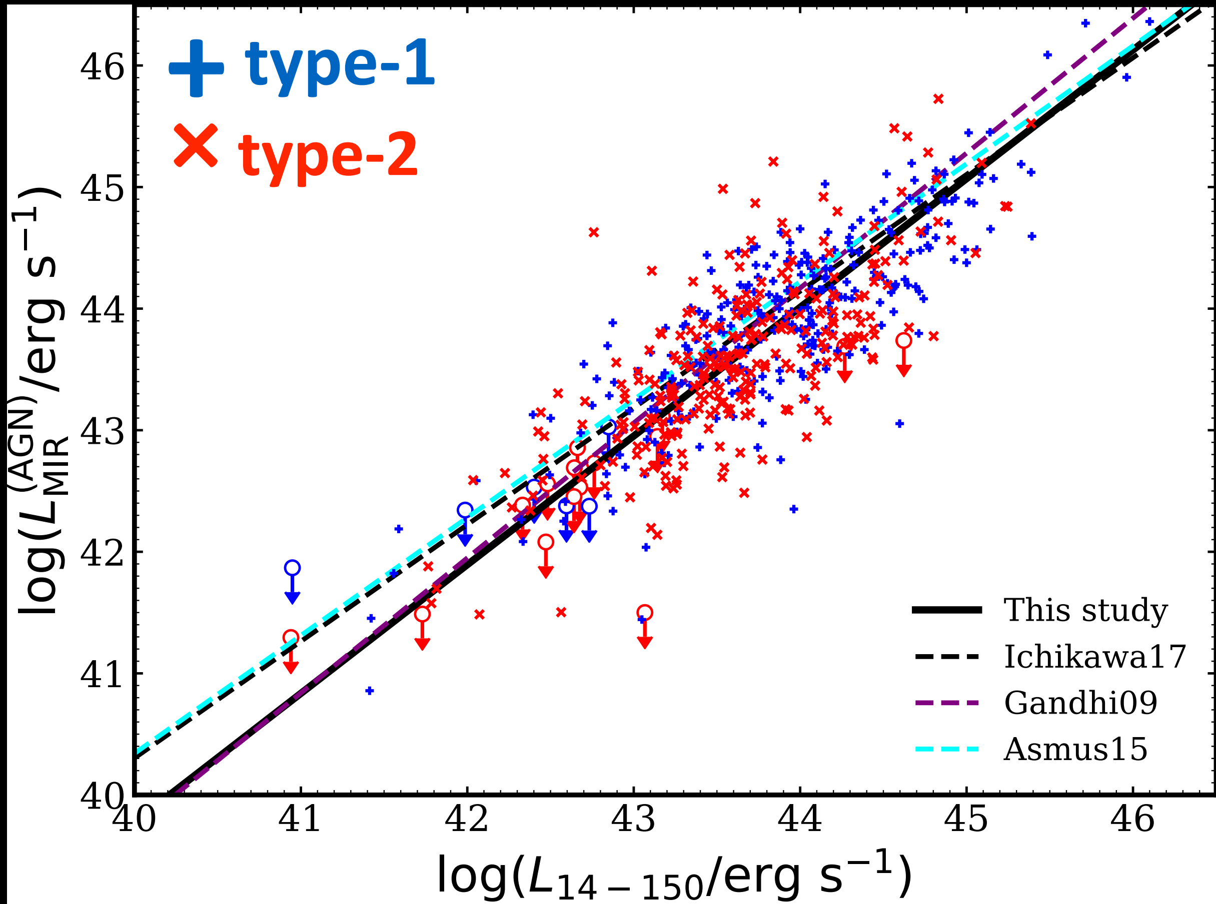
L_{IR}(AGN) vs. $L_{14\text{-}150keV}$

Ichikawa+18

Our study

L_{MIR}/L_X (type-1) ~ L_{MIR}/L_X (type-2)

MIR emission: isotropic



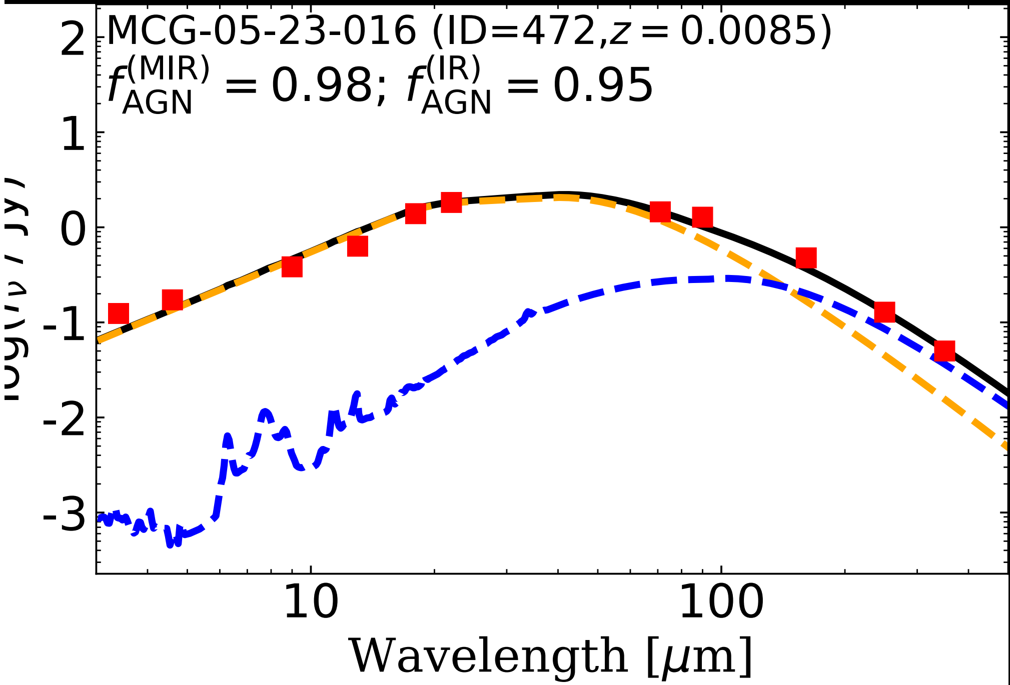
log L_{MIR} ∝ 1.06 log L_X ∴ slope b=1.06 (+/-0.03)

☑ **b=0.9-1.1** from local/X-ray selected AGN

(e.g., Gandhi+09; **Ichikawa+12,+17**; Asmus+15; Mateos+15)

IR-Pure AGN candidates

We found 9 "IR-pure AGN" candidates



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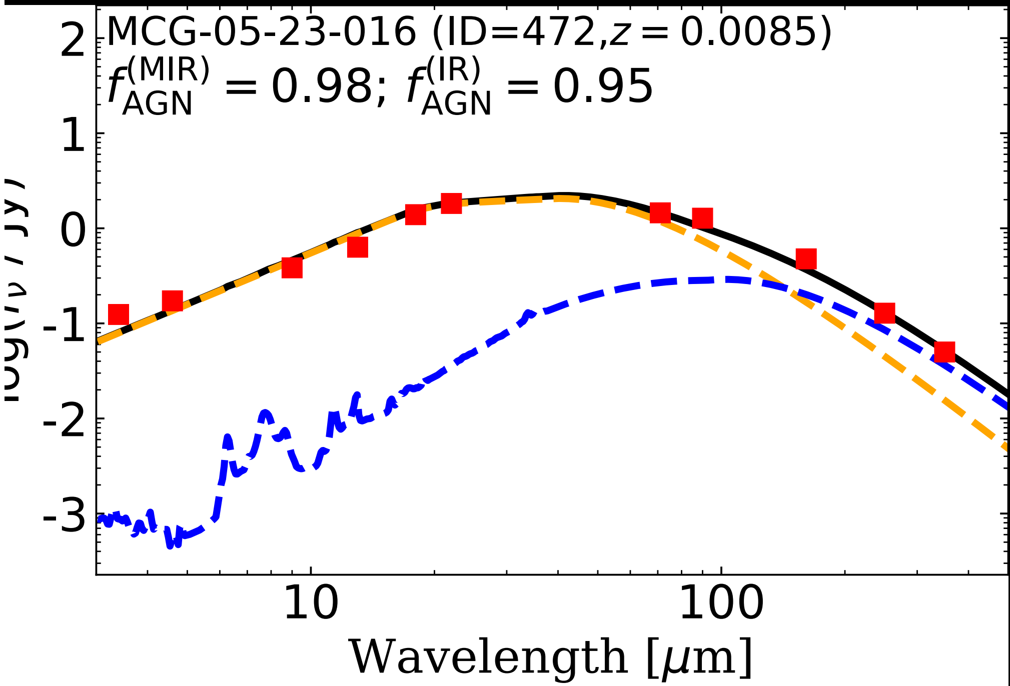
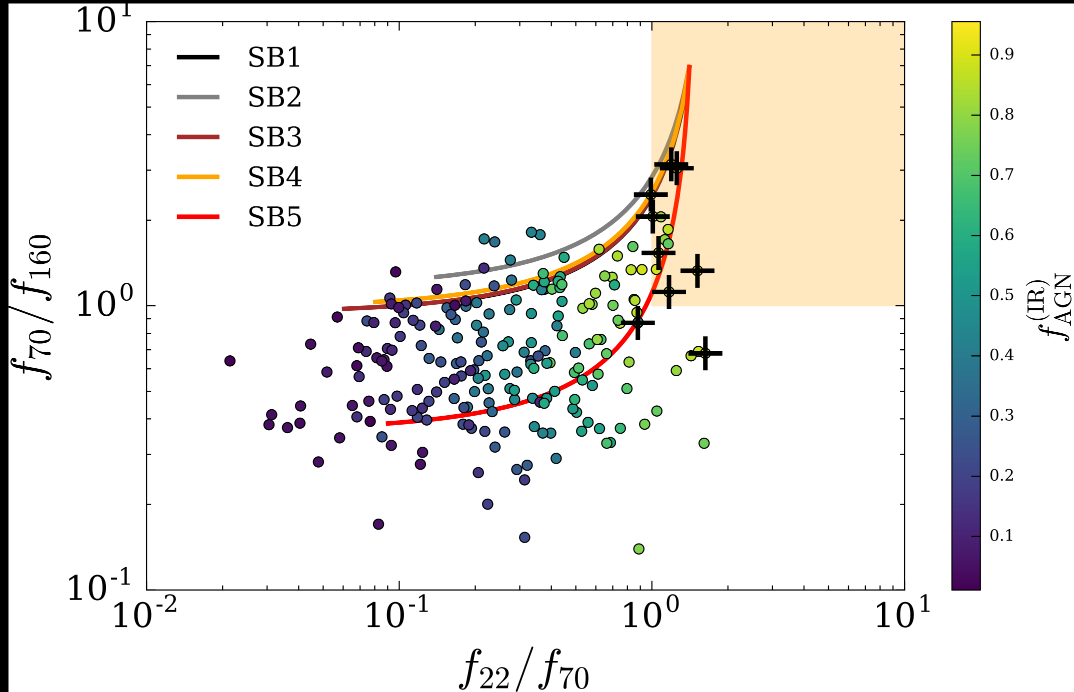




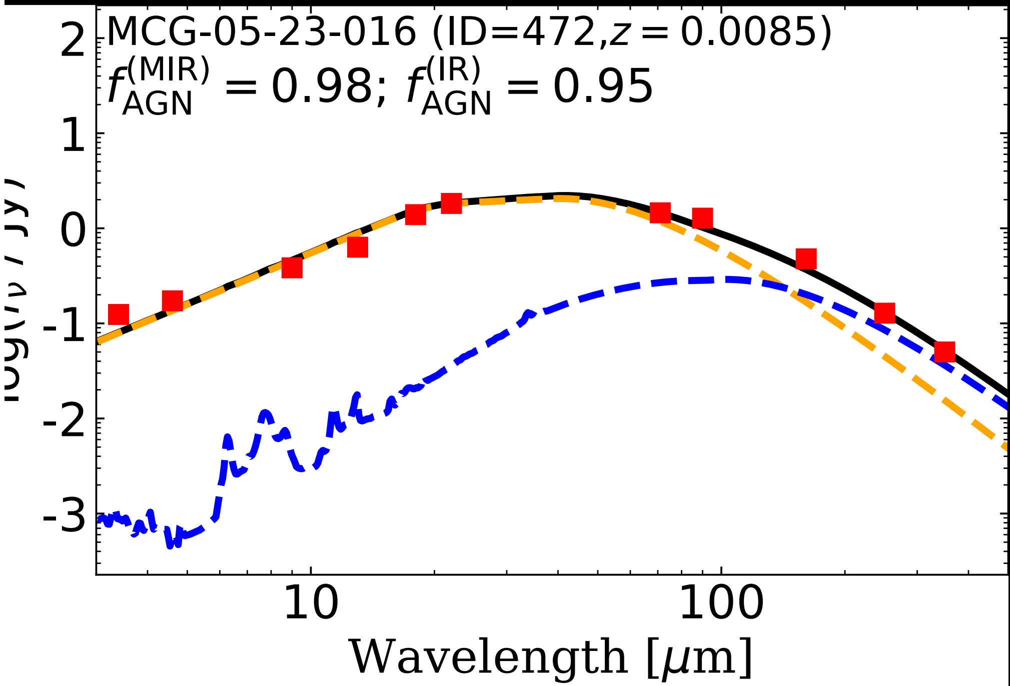
☑ **FIR (up to ~100um) is dominated by AGN torus emission**

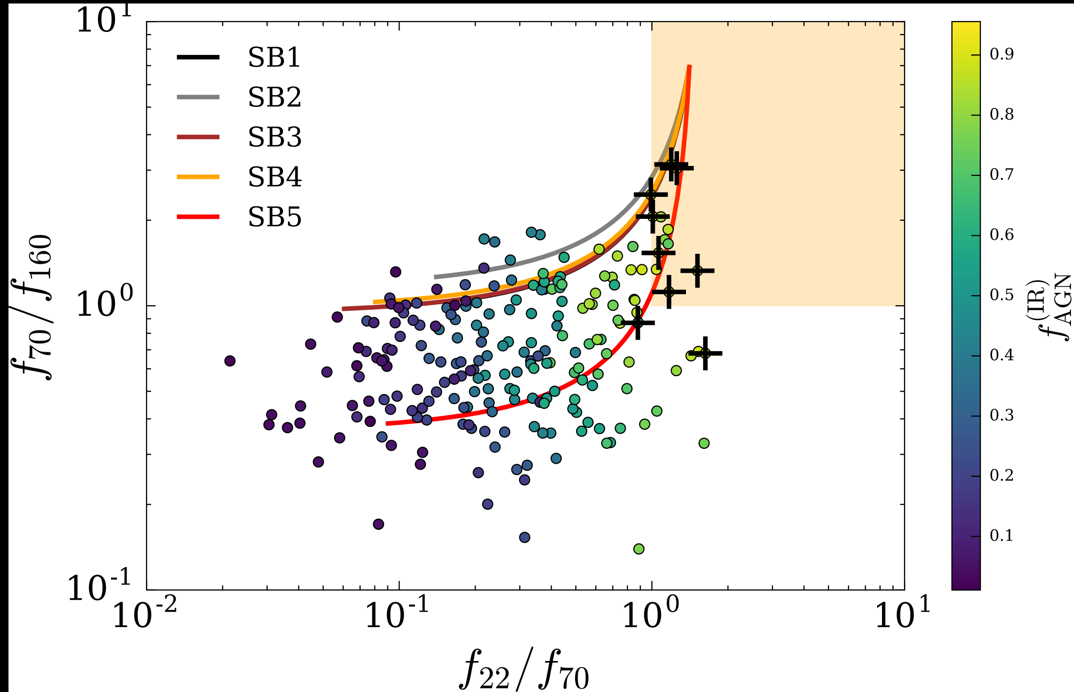
☑ **IR-pure AGN shows the SED w/ $f_{22um} > f_{70um} > f_{160um}$**

IR-Pure AGN candidates

We found 9 "IR-pure AGN" candidates

Ichikawa+18





☑ **FIR (up to ~100um) is dominated by AGN torus emission**

☑ **M_{BH}, $L_{14\text{-}150keV}$ distribution is same as the parent sample (<log M_{BH}>=7.8, <log $L_{14\text{-}150}$>=43.7)**

→ Suggesting **weaker SF activities in the host**

→ **good candidates of final stage AGN?**

Dust Covering factor (C_T) vs. L_{bol}

L_X => L_{bol} (const) and **L_{IR}(AGN) /L_{bol} => C_T** (see Stalevski+16)

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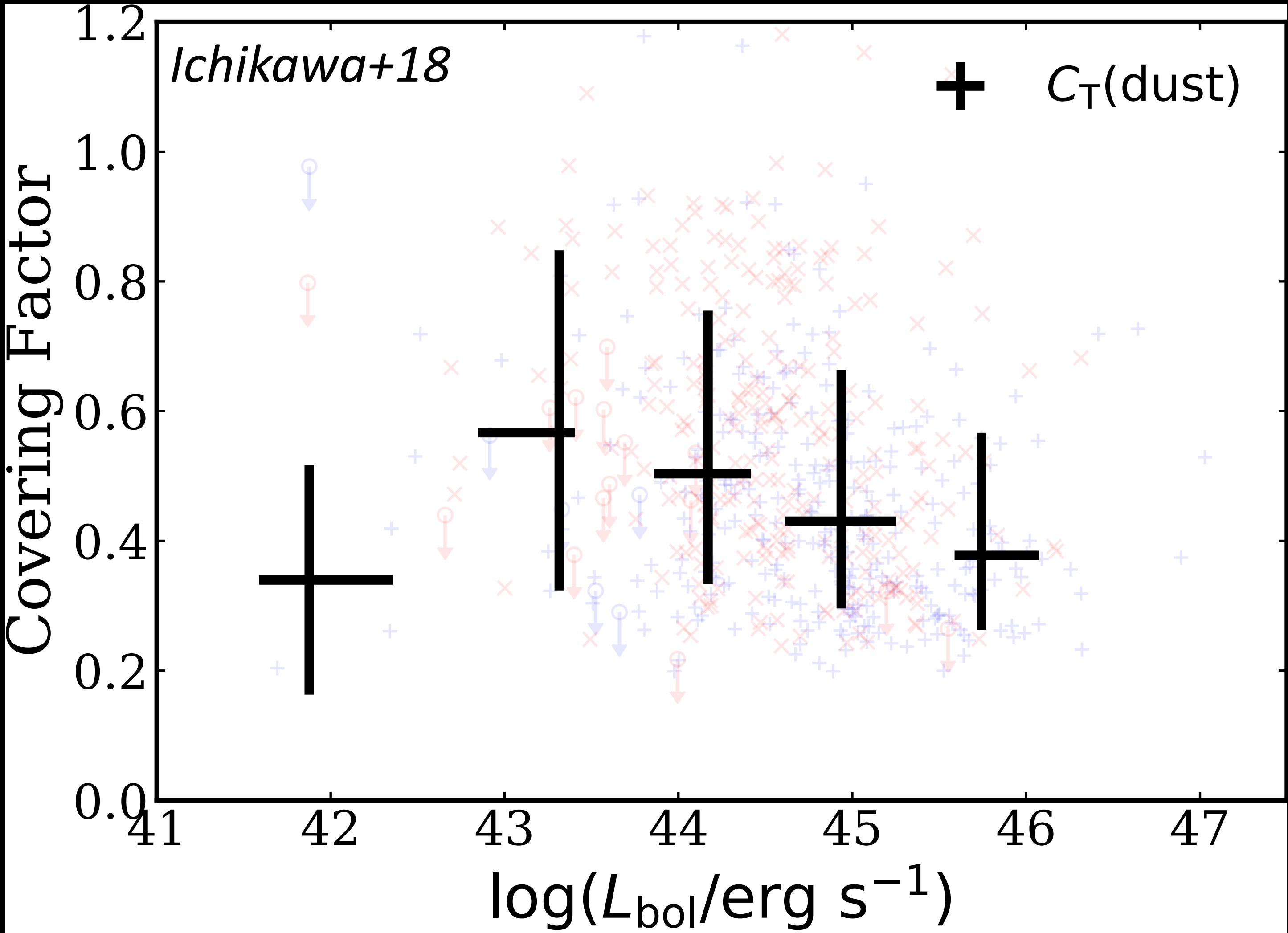
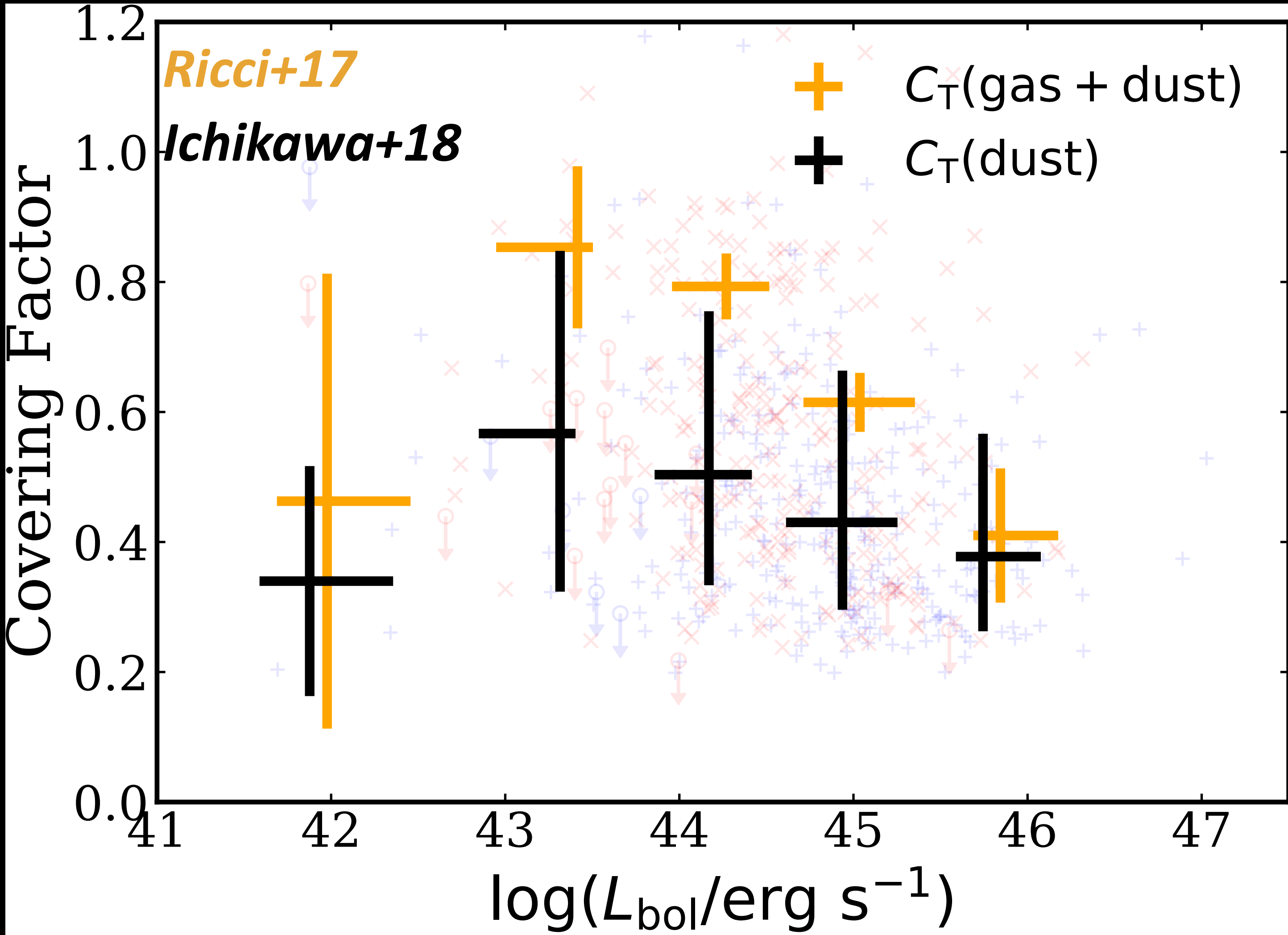


☑ **C_T (dust) : 0.4-0.6,** very weak or almost independent of L_{bol}

(see also Merloni+14, Netzer+16, Stalevski+16, Mateos+17)

Dust Covering factor (C_T) vs. L_{bol}

Lx => L_{bol} (const) and **L_{IR}(AGN) /L_{bol} => C_T** (see Stalevski+16)

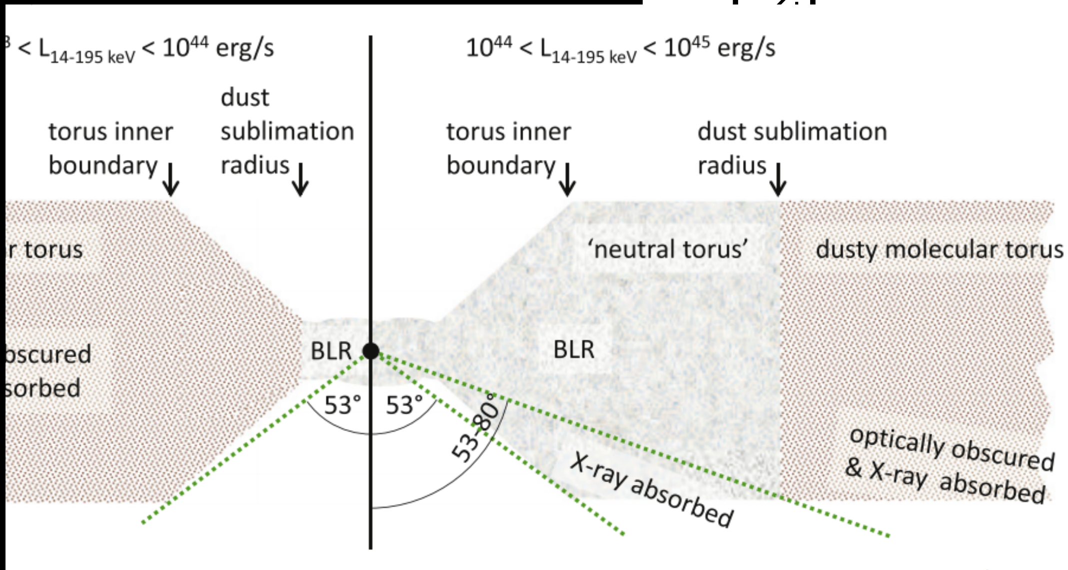


☑ **C_T(dust) < C_T (dust+gas) <= obtained from X-ray obs.**

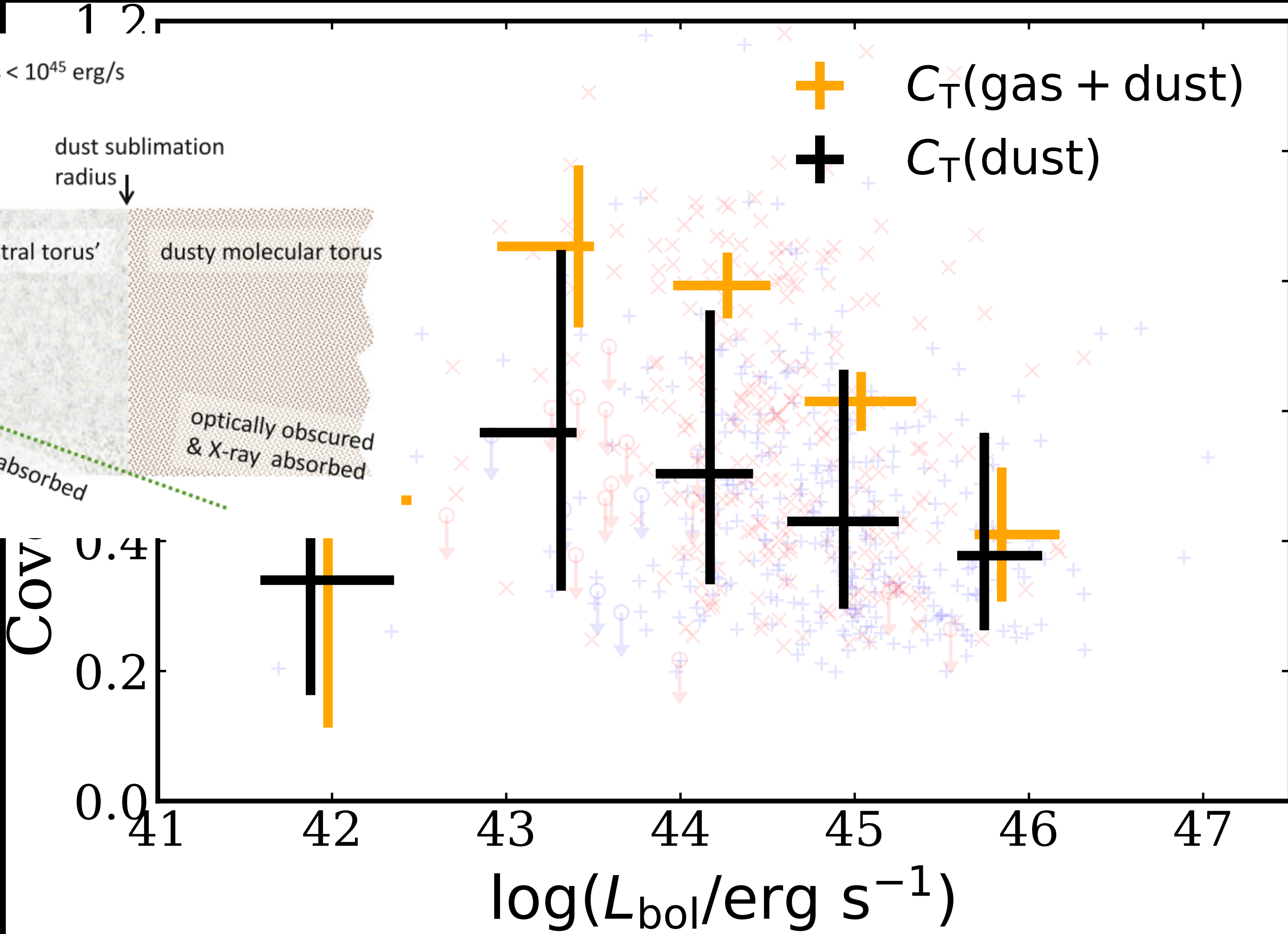
☑ **There is a dust-free (X-ray) obscuring region**

Dust Covering factor (C_T) vs. L_{bol}

Lx => L_{bol} (const) and L_{torus}/L_{bol} => C_T (dust) (see Stalevski+16)



Davies+15

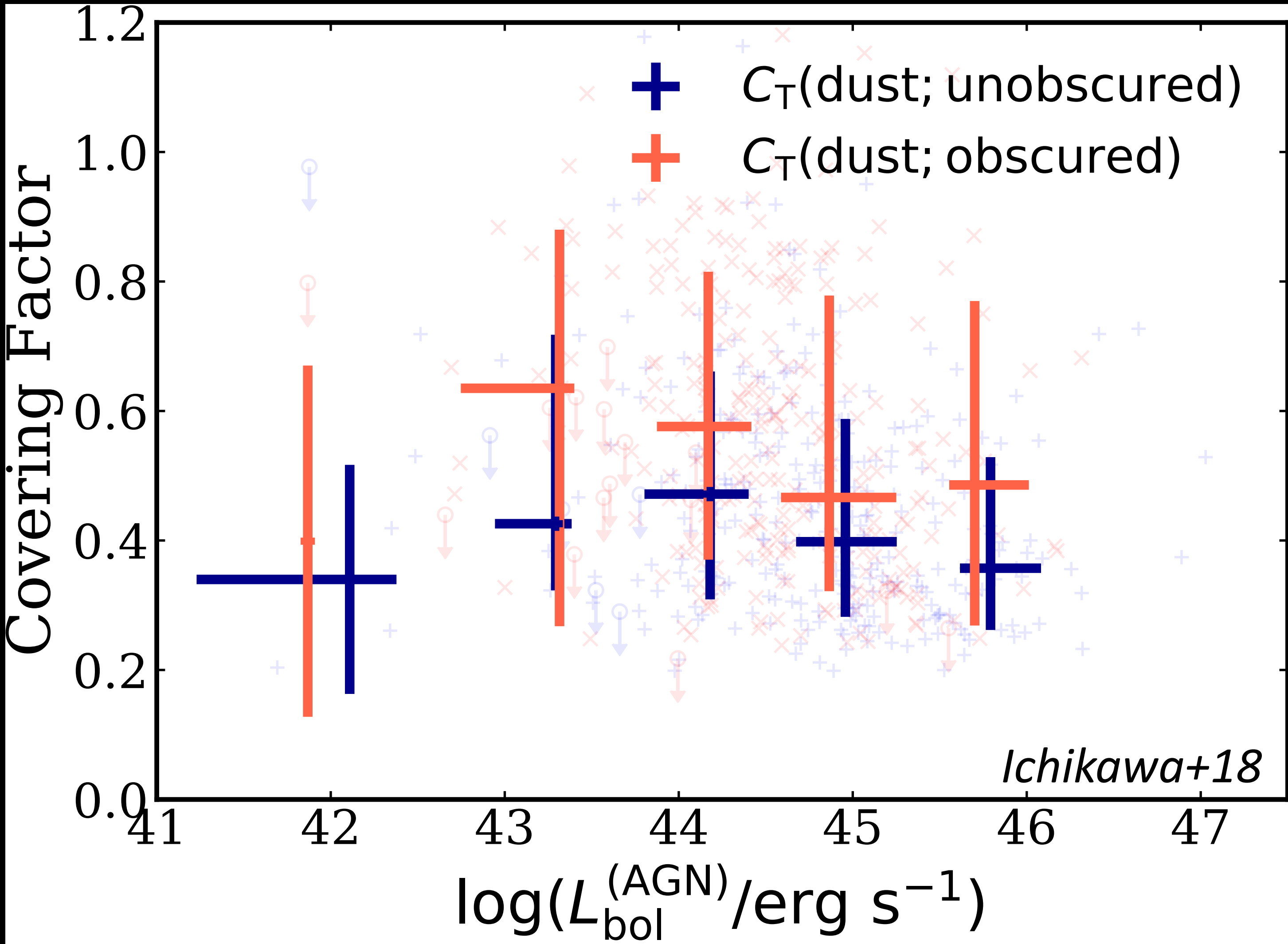


☑ C_T(dust) < C_T (dust+gas) <= obtained from X-ray obs.

☑ There is a dust-free (X-ray) obscuring region

(see also Markowitz+14; Davies+15; Liu+18)

Dust Covering factor (C_T) for un-/obscured AGN



☑ **C_T (obscured) is (on average) always larger than C_T (unobscured)**

=> larger (line of sight) N_H sources tend to have larger (geometrical) C_T

Summary

Swift/BAT (14-195 keV) AGN catalog

☑ suitable sample of an unbiased census of AGN

☑ **BASS provides L_X, N_H, M_{BH}, and λ_E**

☑ **almost complete 3-500 um IR catalog (601/606 at MIR, 402 at FIR, see Ichikawa+17)**

IR and X-ray properties of BAT AGN

☑ 9 IR-pure AGN are found

☑ C_T(dust) < C_T (dust+gas) => dust-free obscuring region

☑ **C_T (obscured) is (on average) always larger than C_T (unobscured)**

see Ichikawa et al. (2017, 2018) for more details

Appendix

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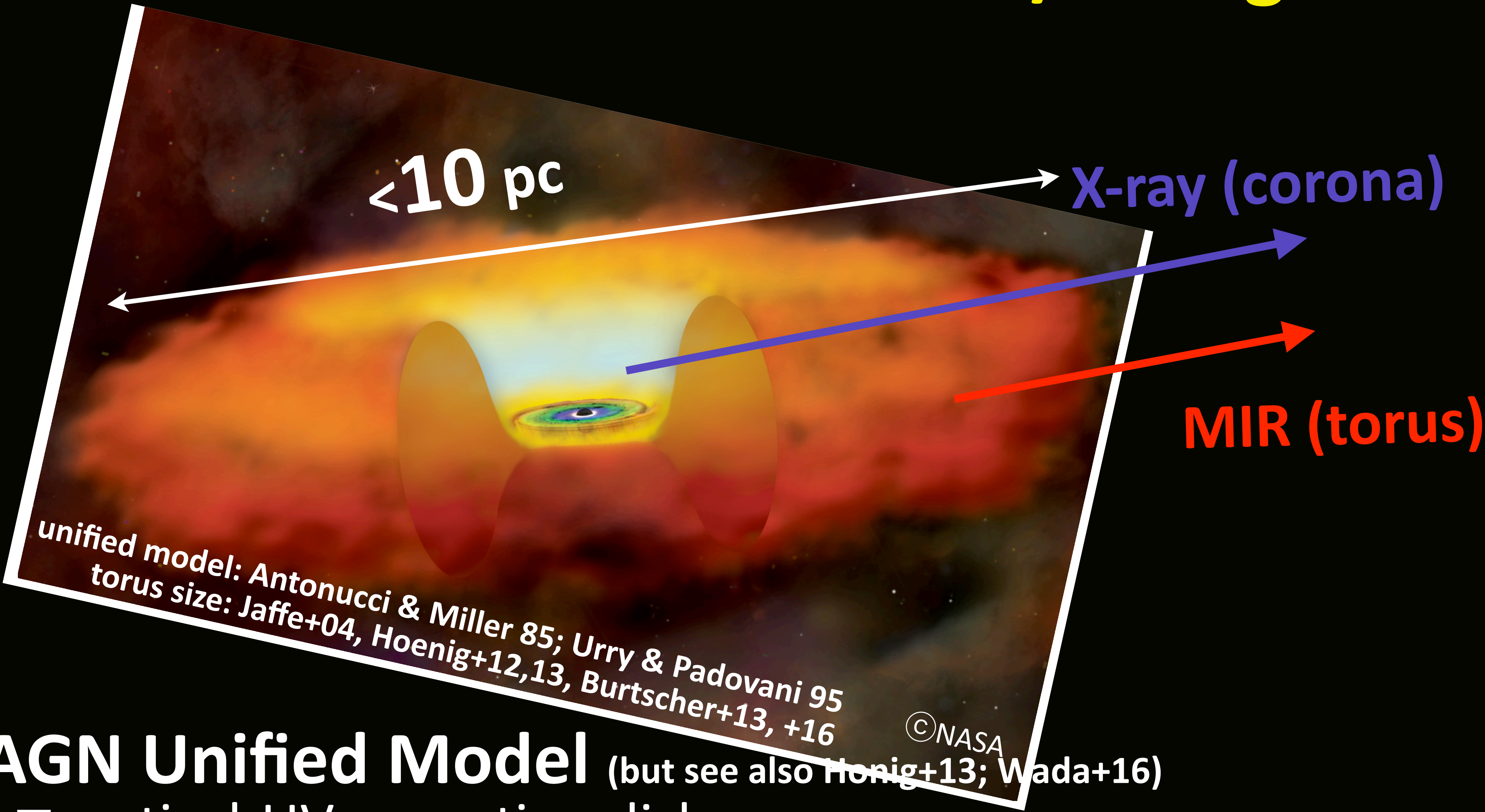
Credit: ESO/M. Kornmesser

IR observations of AGN= torus/host galaxies



AGN Unified Model (but see also Honig+13; Wada+16)

- ☑ optical-UV: accretion disk
- ☑ X-ray: accretion disk+hot electron corona
- ☑ **mid-IR (MIR): dusty torus (dust/gas provider to SMBH)**
- ☑ **far-IR (FIR): host galaxy**

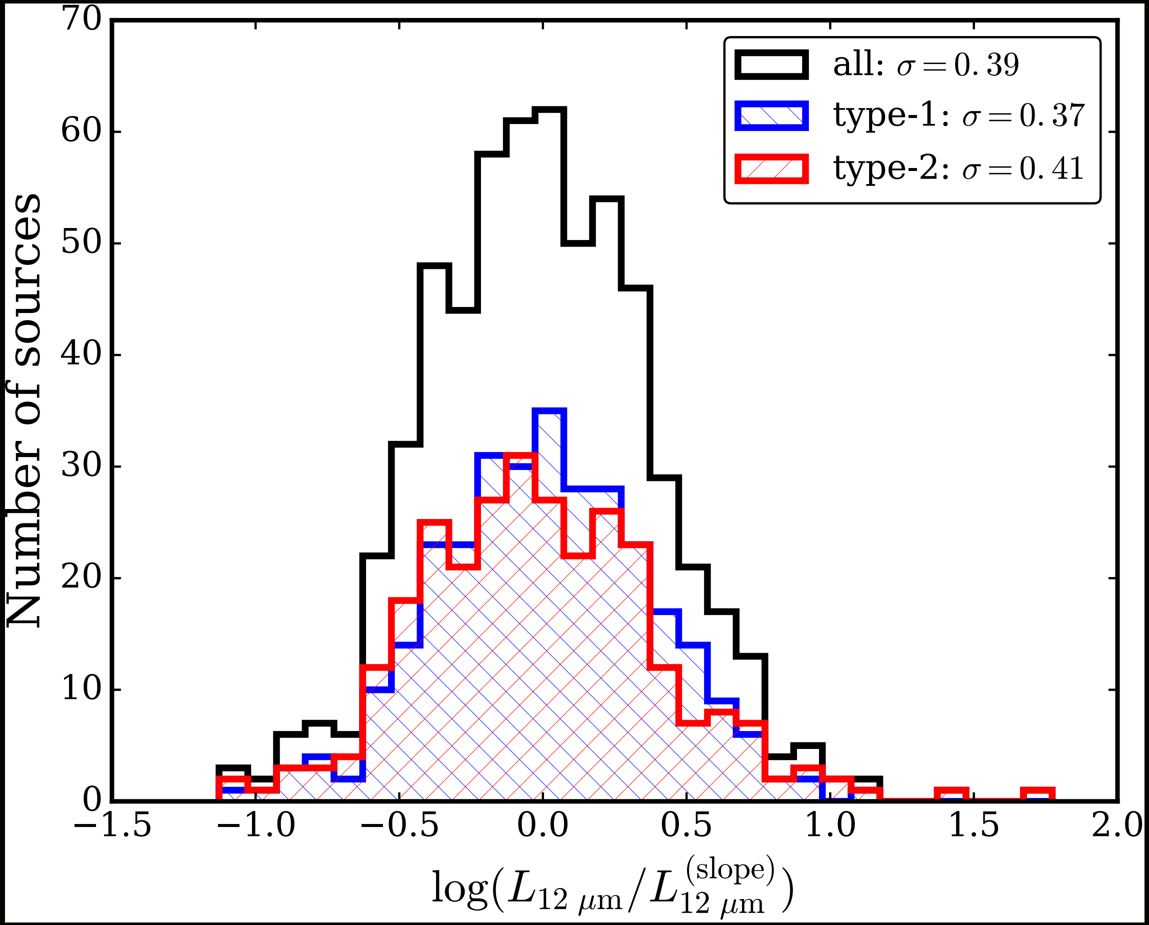
IR properties of BAT AGN

(see Ichikawa+17a for more details)

SED Decomposition

Ichikawa et al. (2018)

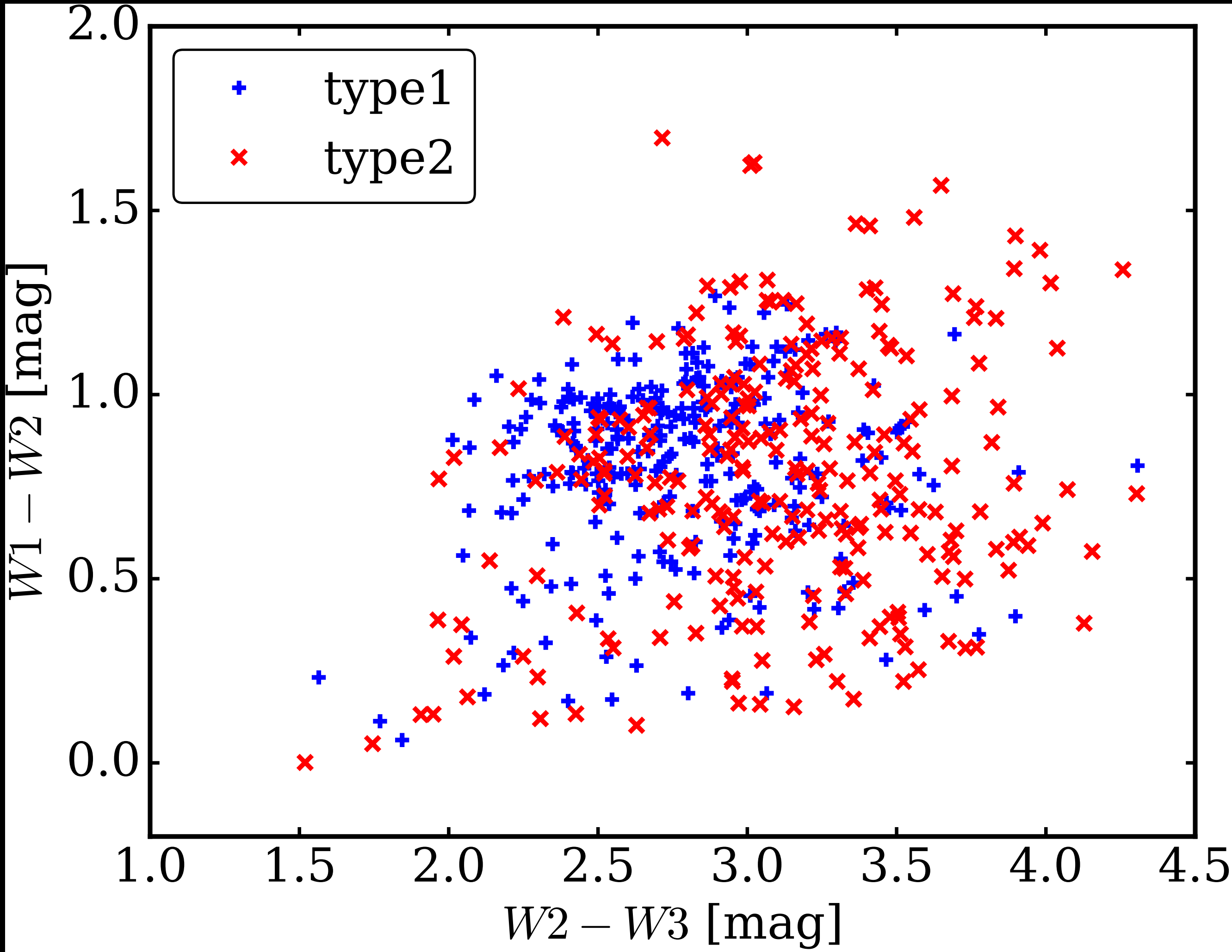
Consistency with dust polar emission



- ☑ type-1/-2 has same distribution => isotropic emission
- ☑ consistent with MIR polar emission or fountain model

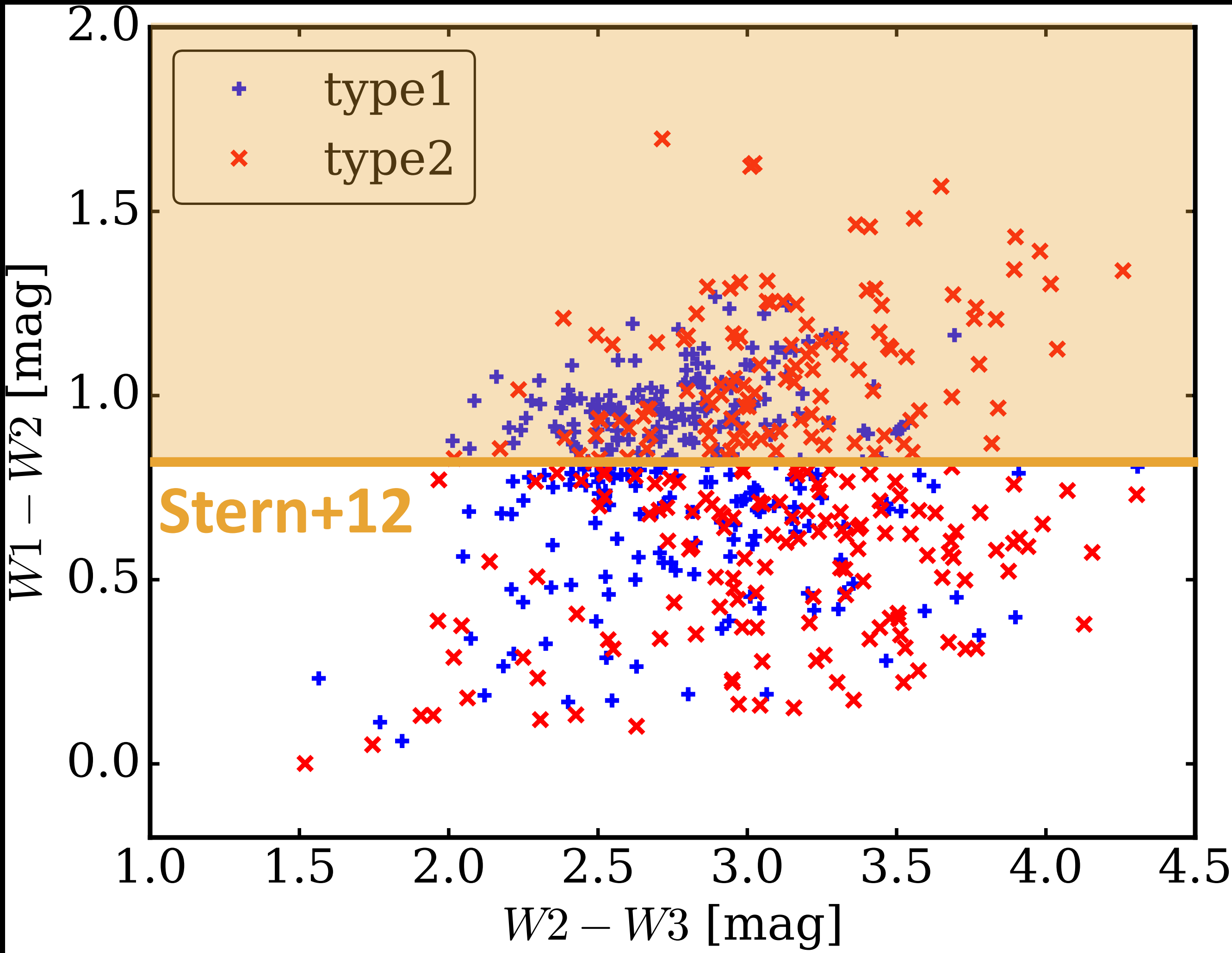
obs: Honig+13,+14, see also Asmus+16
model: Wada 12, Wada+16

WISE IR color-color selection of AGN



Ichikawa+17

WISE IR color-color selection of AGN



Ichikawa+17

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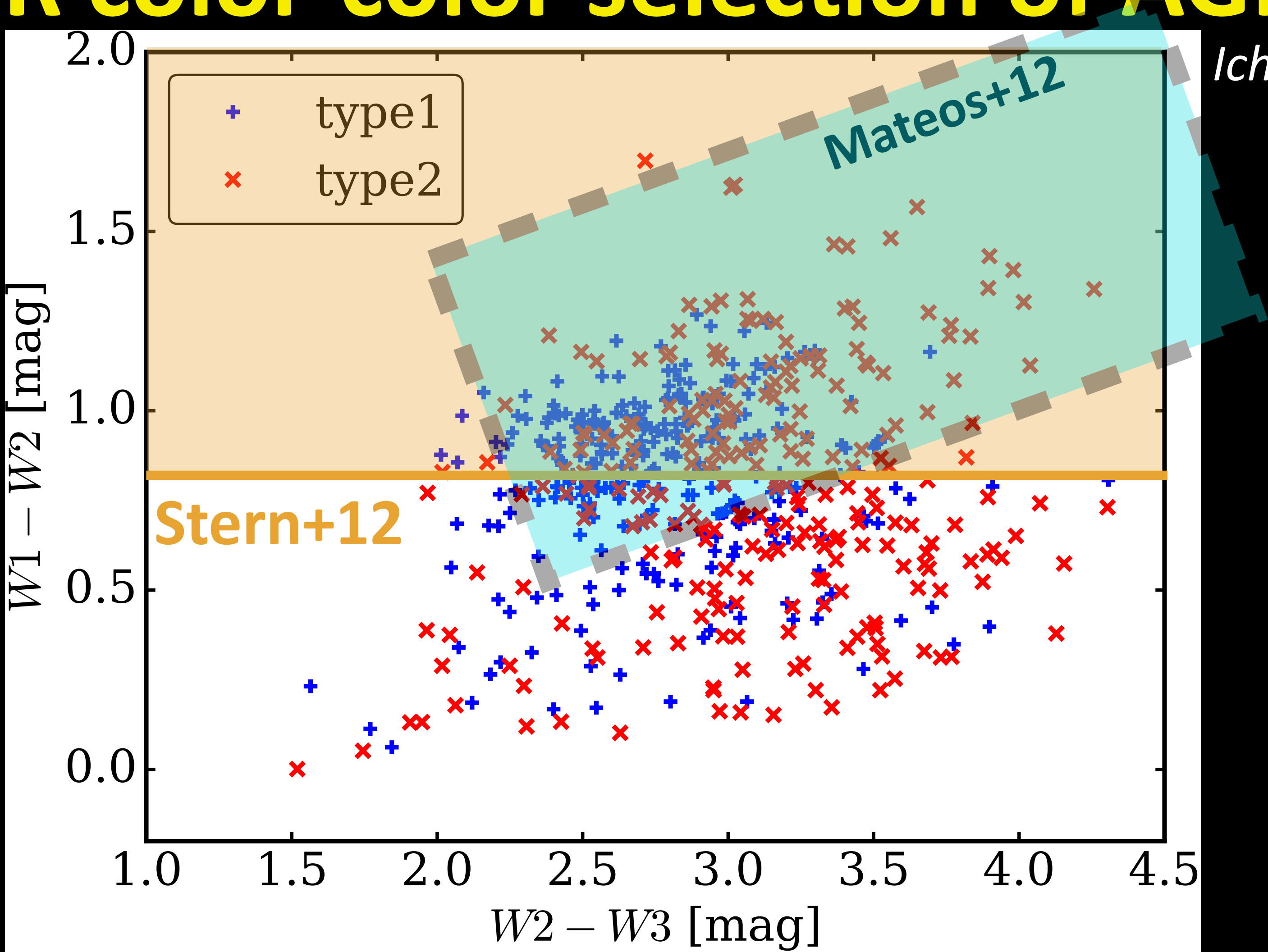


Ichikawa+17

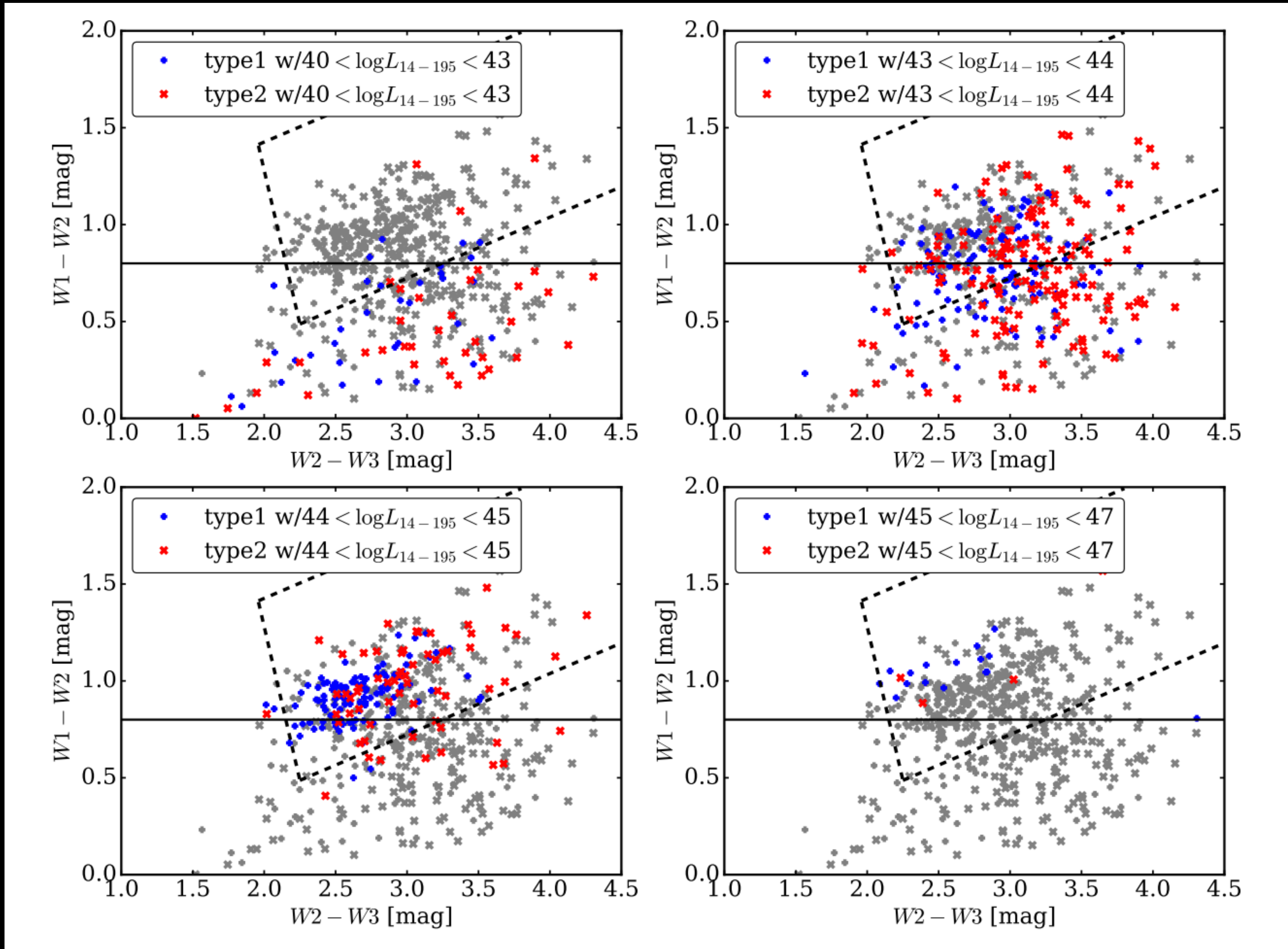
☑ **BAT-**AGN **do not always locate** at the IR selection areas of. Stern+12, Mateos+12

WISE IR color selections miss some AGN population

(see also Mateos+12, 13; Gandhi+16; Kawamuro+16; Tanimoto+16)

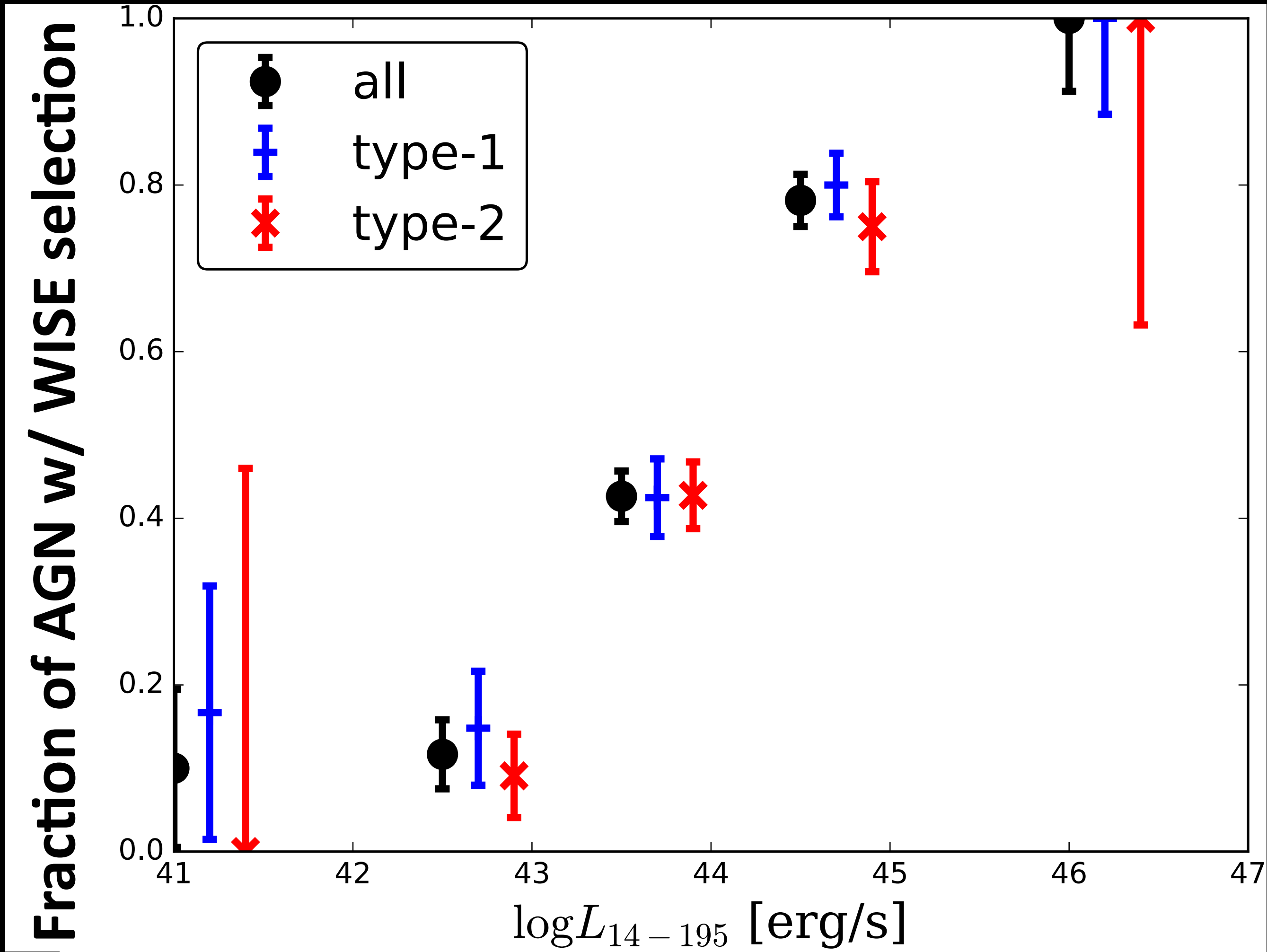
WISE IR color-color selection of AGN

Ichikawa+17



☑ **WISE IR color: insensitive to low-luminosity AGN**

Success rate of WISE color selection



Ichikawa+17

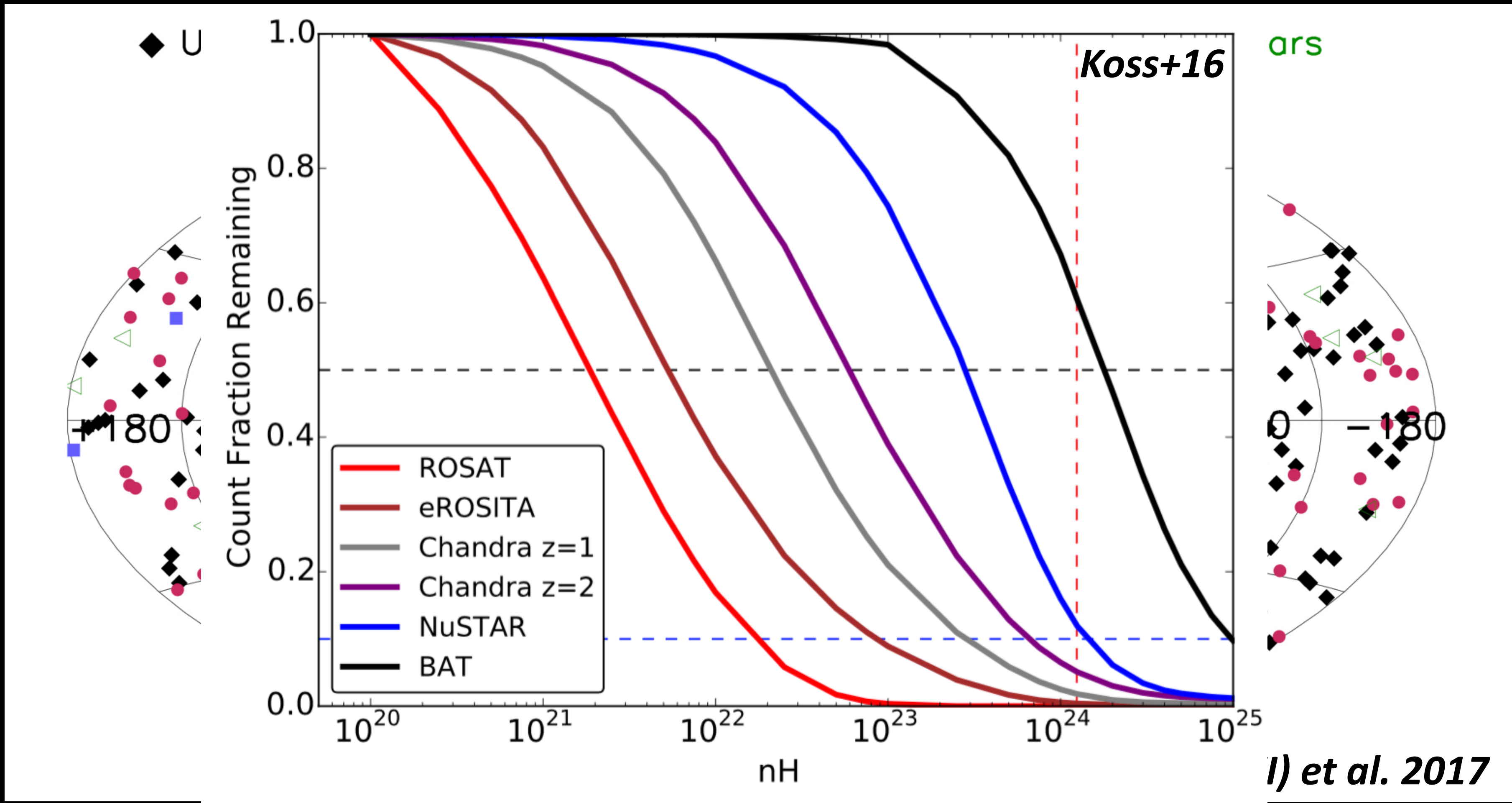
☑ **WISE IR color: insensitive to low-luminosity AGN**

☑ **<20% success rate** for low-luminosity AGN of **log Lx < 43**

Swift/BAT AGN (14-195 keV)

70 month catalog: 836 AGN (728 non-blazars)

FYI, 105 month catalog is public (Oh et al., '18)



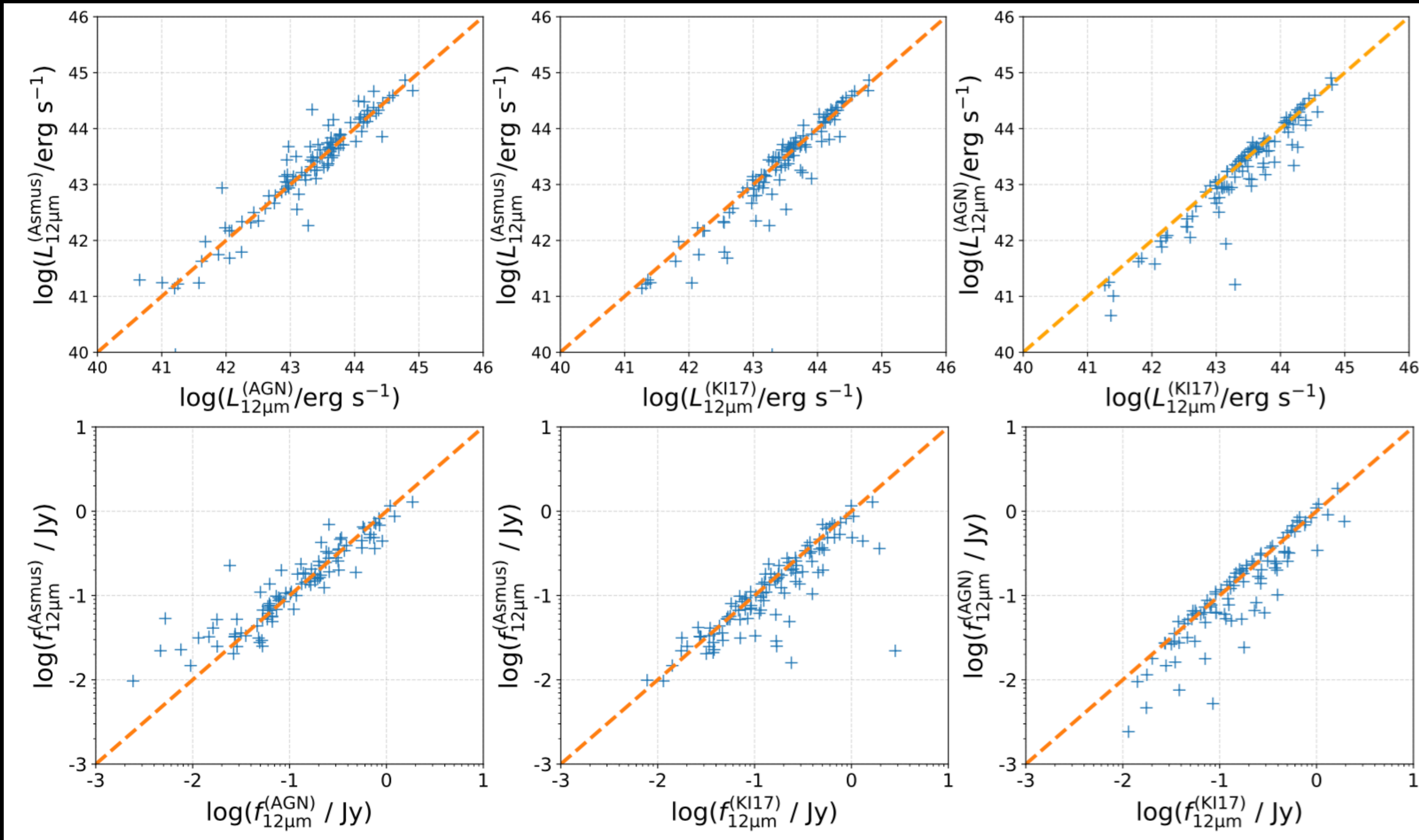
☑ most complete up to $\log N_H$=24.5 in the local Universe

(Koss+16; Ricci+15)

☑ **606 out of 728 have z info and are located at |b|>10°**

Comparison with high-spatial resolution observations

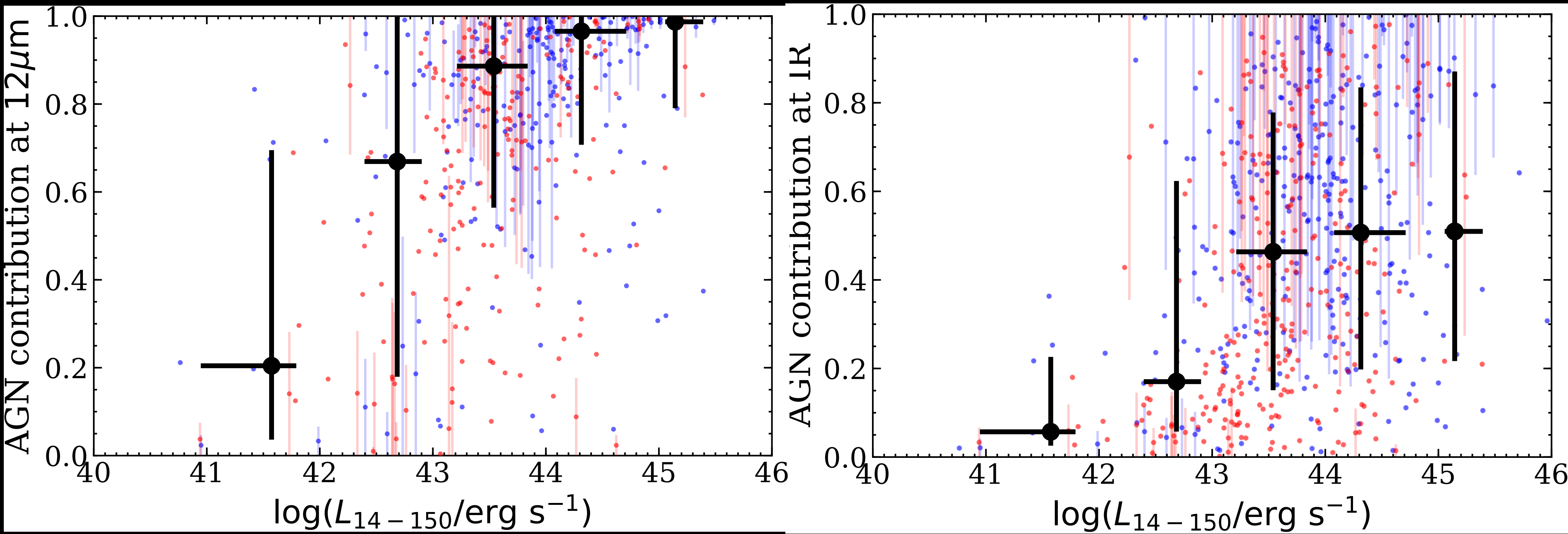
☑ **Decomposition works really well!**



☑ **Disentangling AGN/(SB+stellar) component**

☑ **suitable for the AGN torus/host galaxy studies**

AGN contribution as a function of L_{BAT}



Ichikawa+18

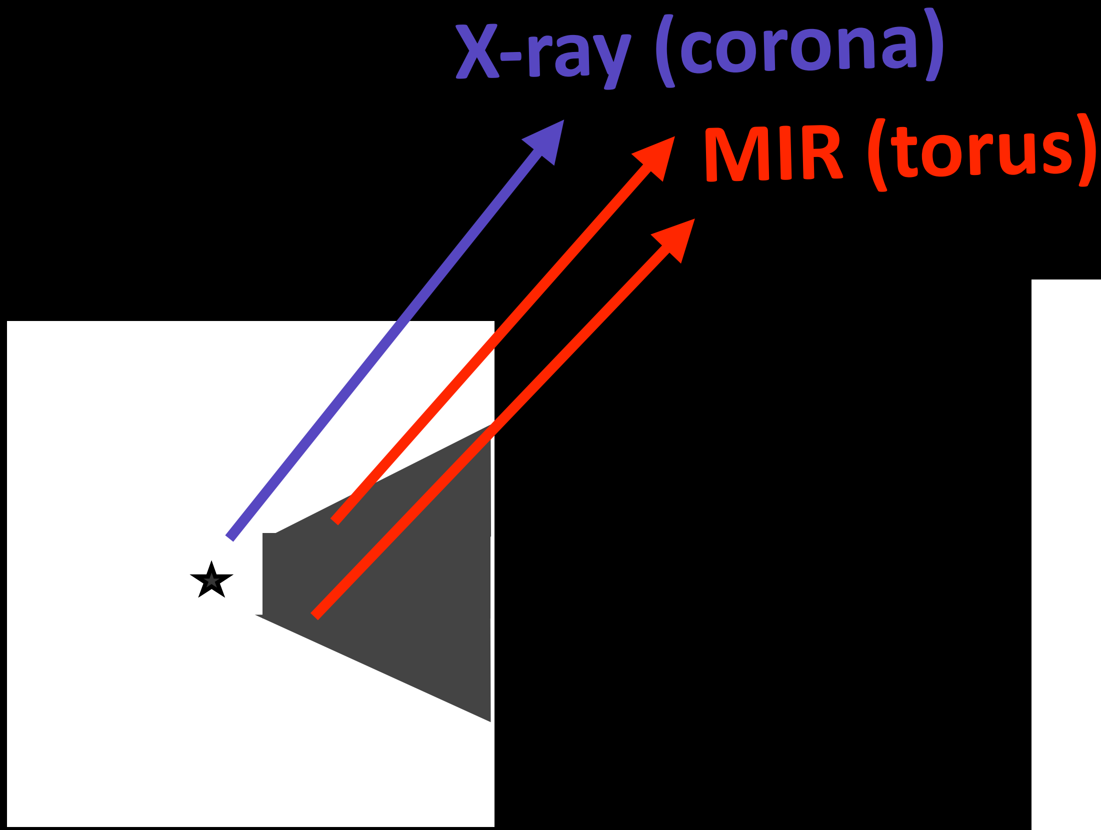
☑ **At high L_{BAT} end,** contribution reaches **~100% at 12um, 80% at MIR (5-40um), and 50% at total IR**

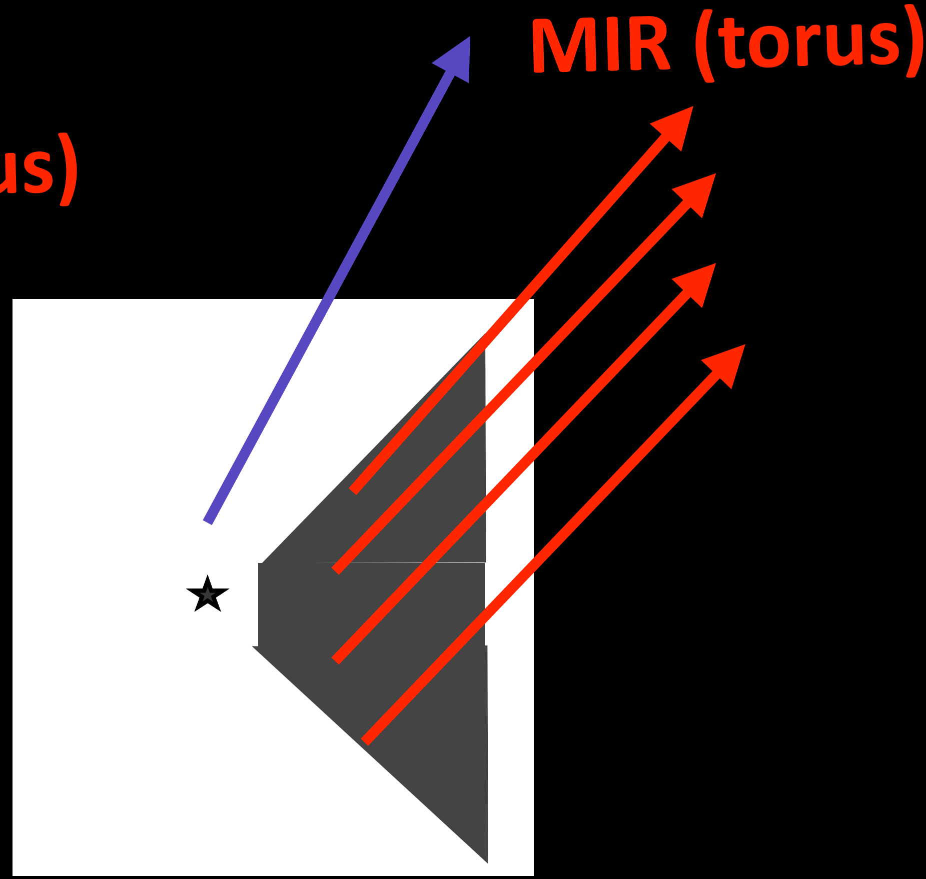
☑ **At low L_{BAT} end,** contribution goes down to **~20% at 12um, 20% at MIR (5-40um), and <10% at total IR**

→ **SED decomposition is crucial for low-luminosity AGN**

Dust Covering factor (C_T) vs. L_{bol}

$L_X \Rightarrow L_{bol}$ and $C_T \propto L_{MIR}/L_{bol}$ (see Stalevski+16)



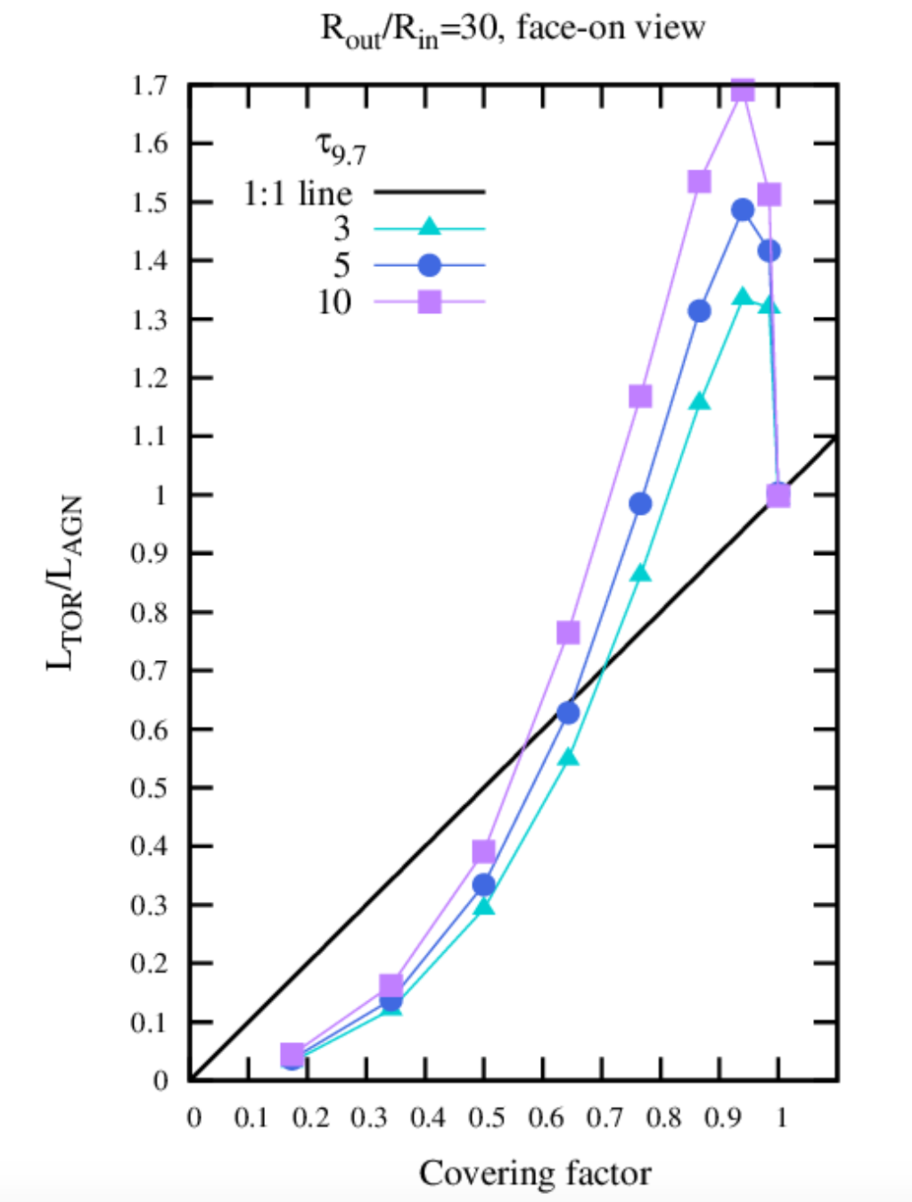


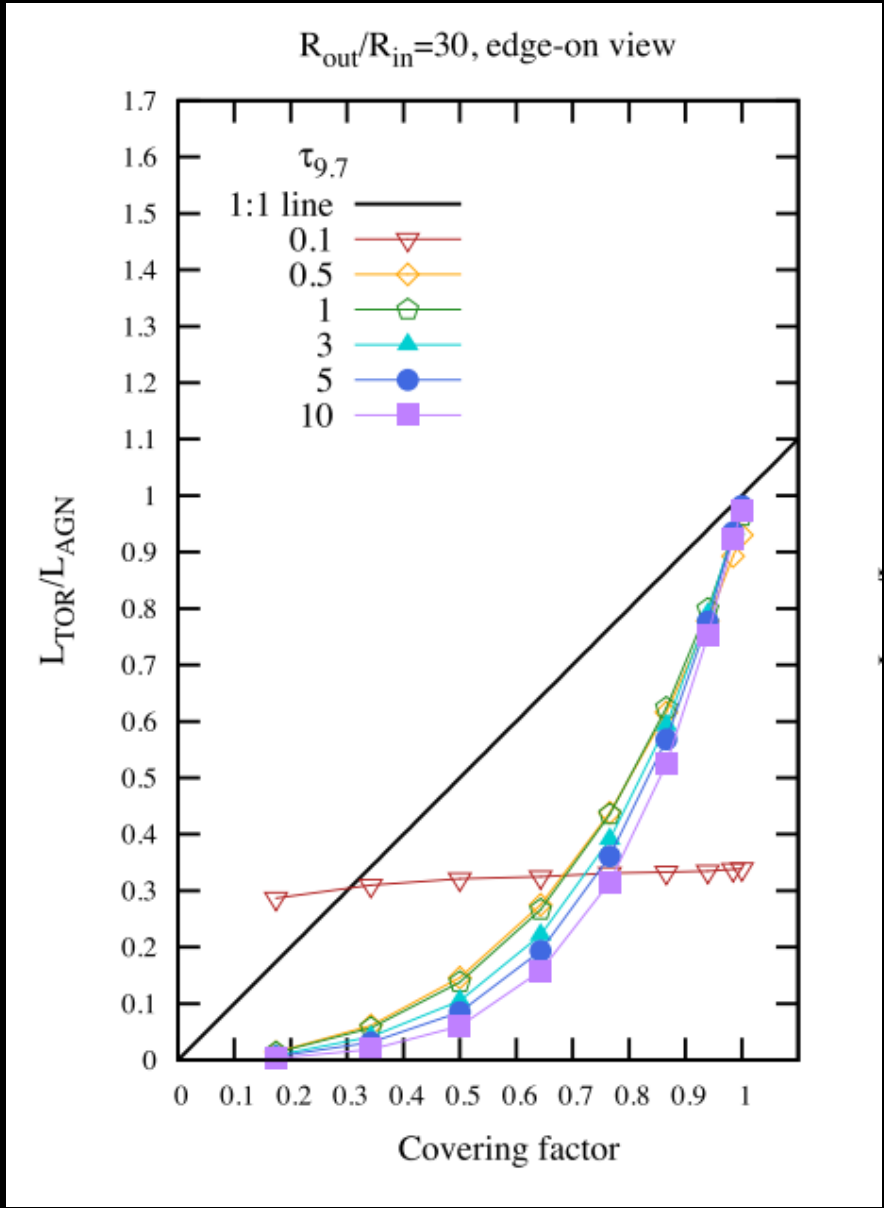
C_T: indicator of geometrical dust obscuration

$$L_{MIR} \propto L_{bol} C_T \Leftrightarrow C_T \propto L_{MIR}/L_{bol}$$

C_T(dust) vs. L_{bol} by Stalevski+16

C_T vs. L_{MIR}/L_{bol} (see Stalevski+16)

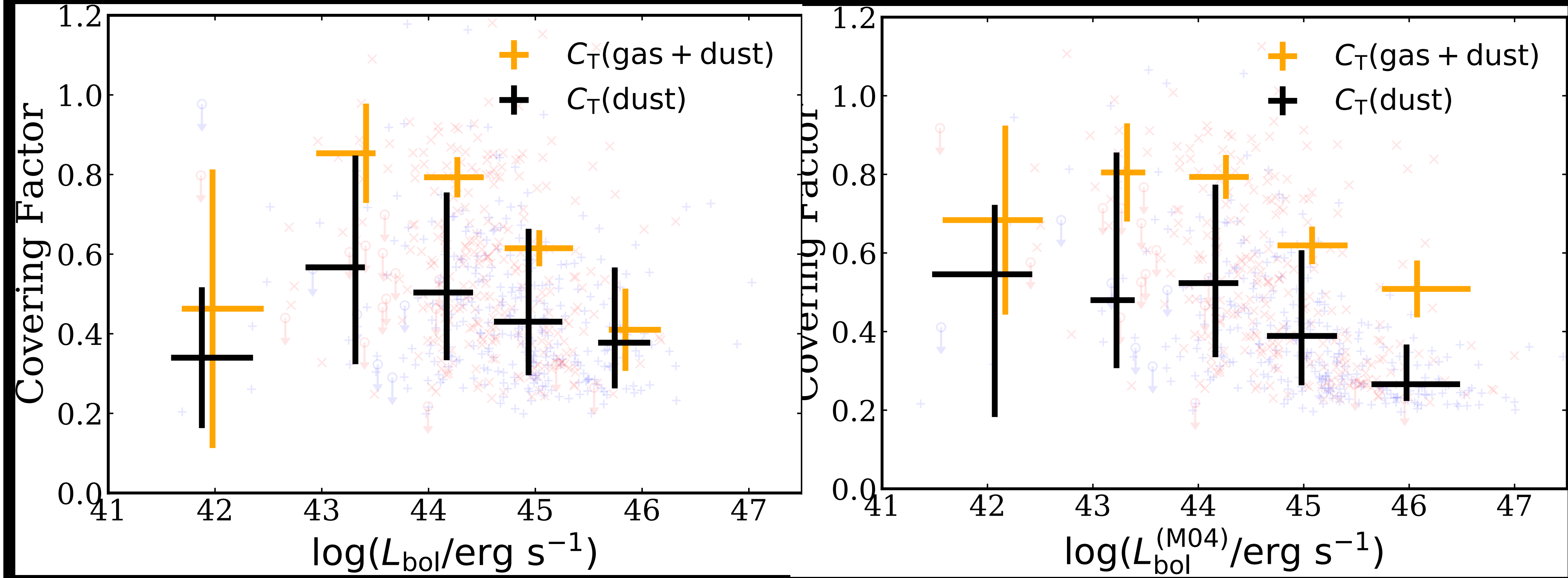




Dust Covering factor (C_T) vs. L_{bol}

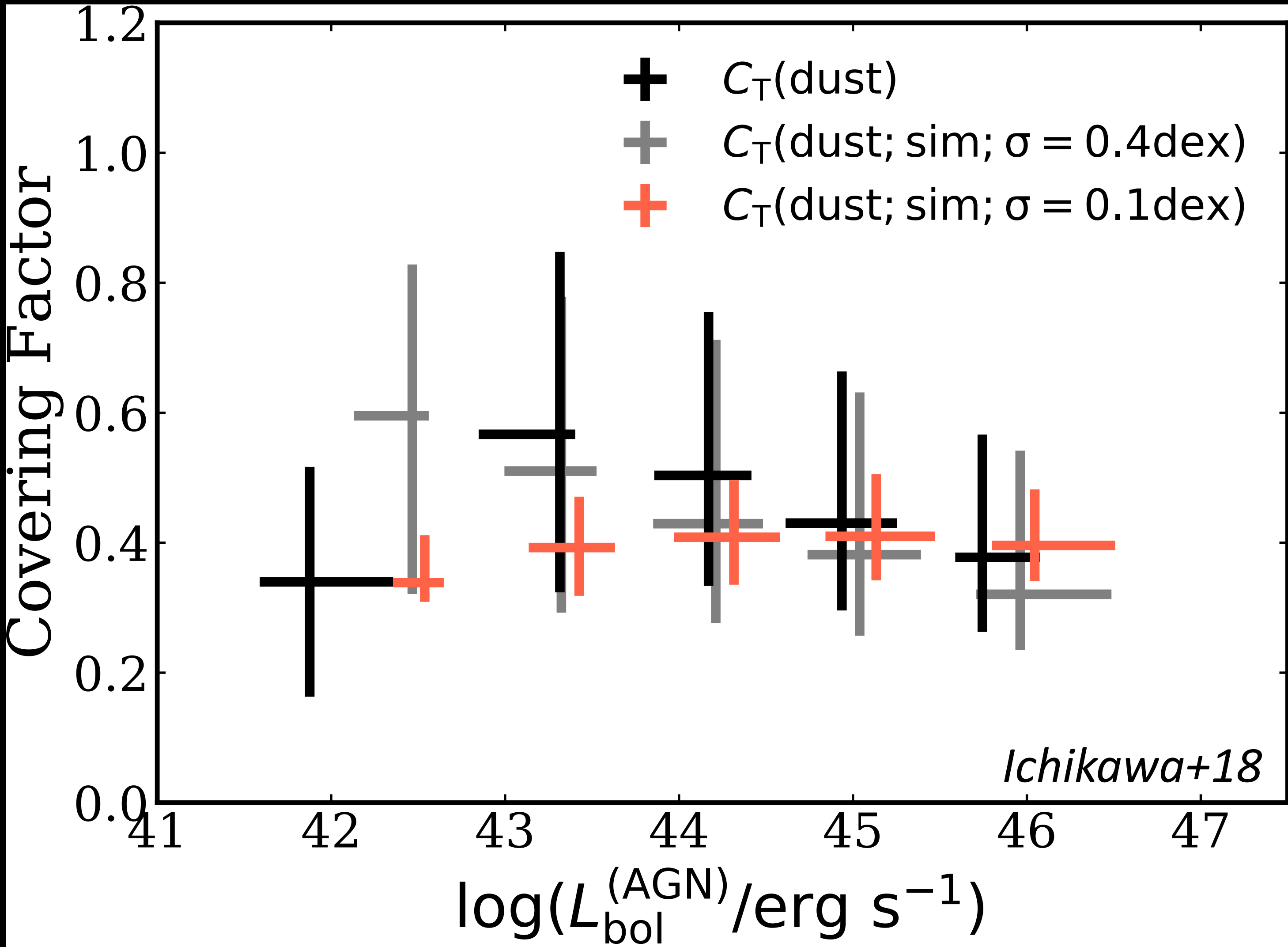
Lx => L_{bol} (Marconi+04) and **$C_T \propto L_{MIR}/L_{bol}$** (see Stalevski+16)

Ichikawa+18



☑ **Different bol-correction does not change the main result**

L_{bol} dependence of Dust Covering factor (C_T)

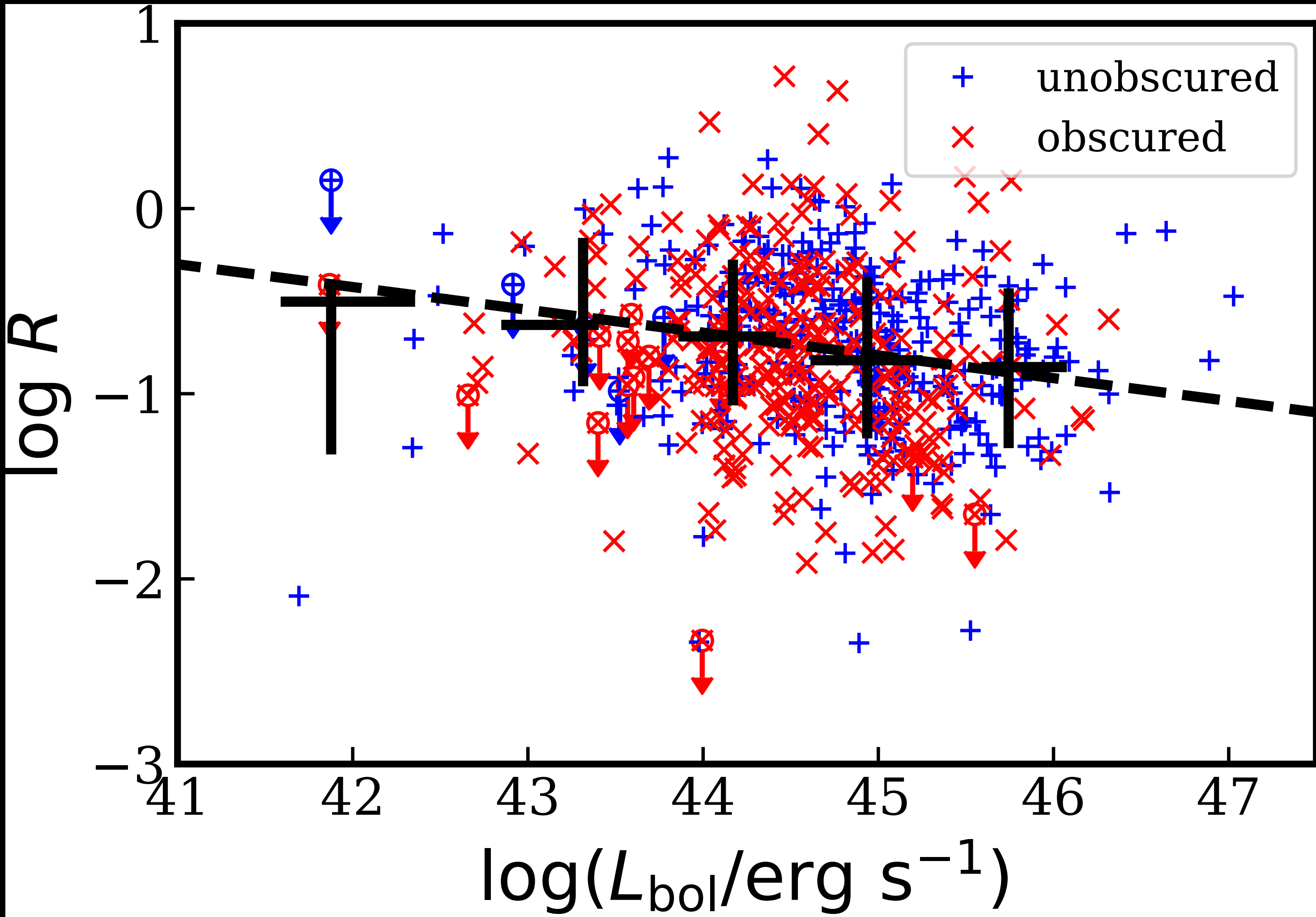


☑ Small scatter of L_X-L_{IR} relation gives a flatter L_{bol} dependence

☑ This is because $\log L_{MIR} \propto 1.06 \log L_X$

∴ slope b=1.06 (+/-0.03)

L_{bol} dependence of R = $L_{IR}(AGN)/L_{bol}$



☑ **Very shallow L_{bol} dependence w/ log R = 4.5 - 0.12 log L_{bol}**