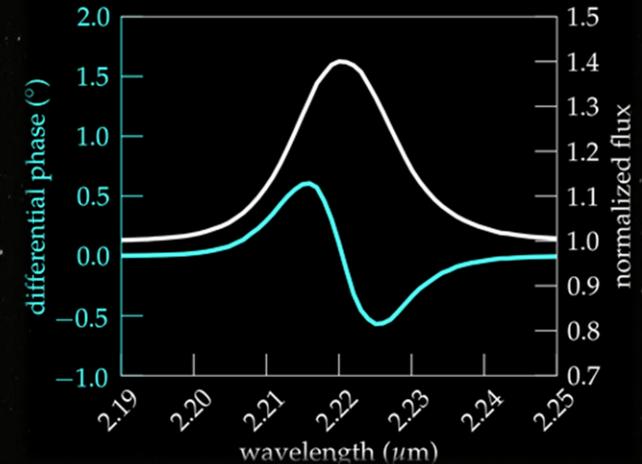


Spectro-Interferometric Signatures of the Broad Line Regions in Active Galactic Nuclei

Matthias Raphael Stock

Max Planck Institute for Extraterrestrial Physics

TORUS Conference, 10 December 2018

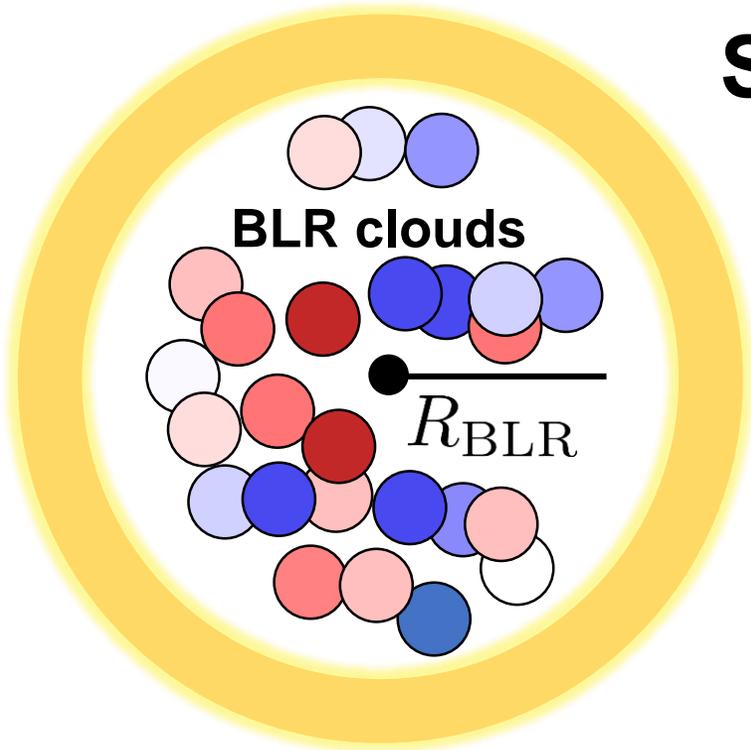


$$R_{\text{BLR}} \lesssim 0.1 \text{ mas}$$

GRAVITY at the VLTI

$$\Delta\theta = \frac{\lambda}{B} = \frac{2.2 \text{ } \mu\text{m}}{130 \text{ m}} \sim 3.5 \text{ mas}$$

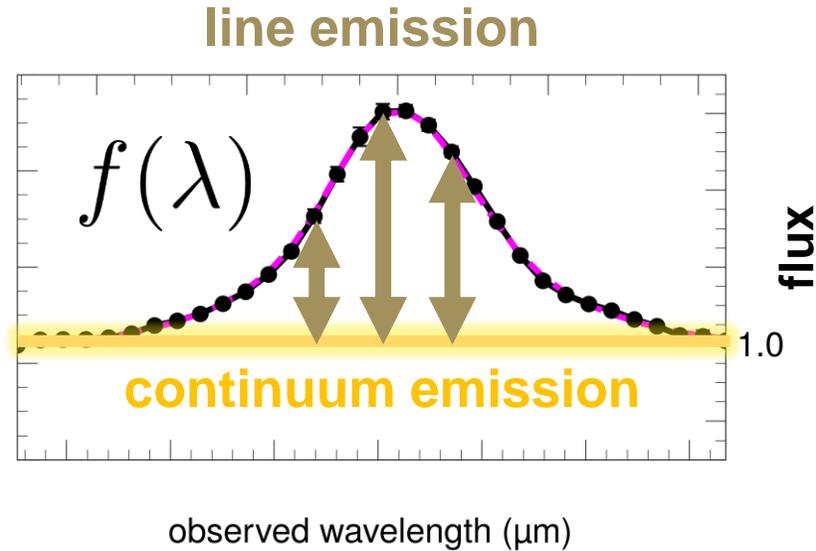
Spectro-Interferometry: Differential Phase



symmetric continuum

$$\Delta\phi(\lambda) = -2\pi f_{\text{line}}$$

line/continuum emission

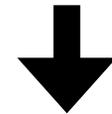


$$f_{\text{line}} = \frac{f(\lambda)}{1+f(\lambda)}$$

Quasars

$$f(\lambda) \lesssim 0.5 \Rightarrow |\phi| \lesssim 2.7^\circ$$

sensitivity of GRAVITY $\lesssim 0.5^\circ$
per baseline per hour



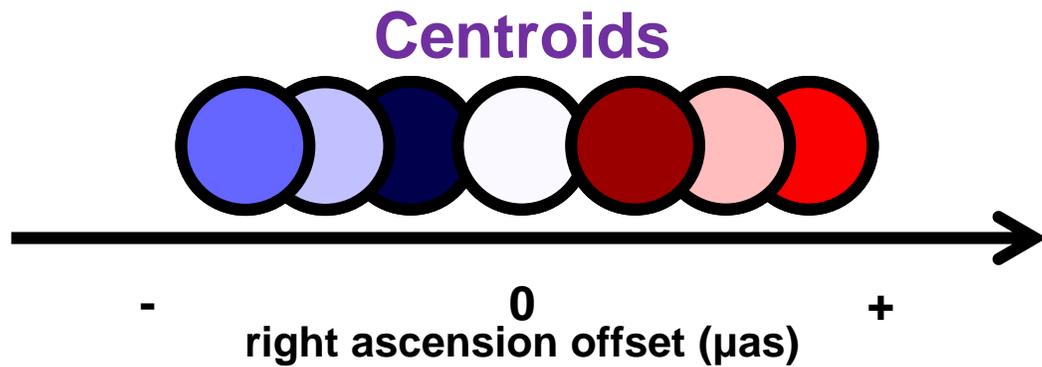
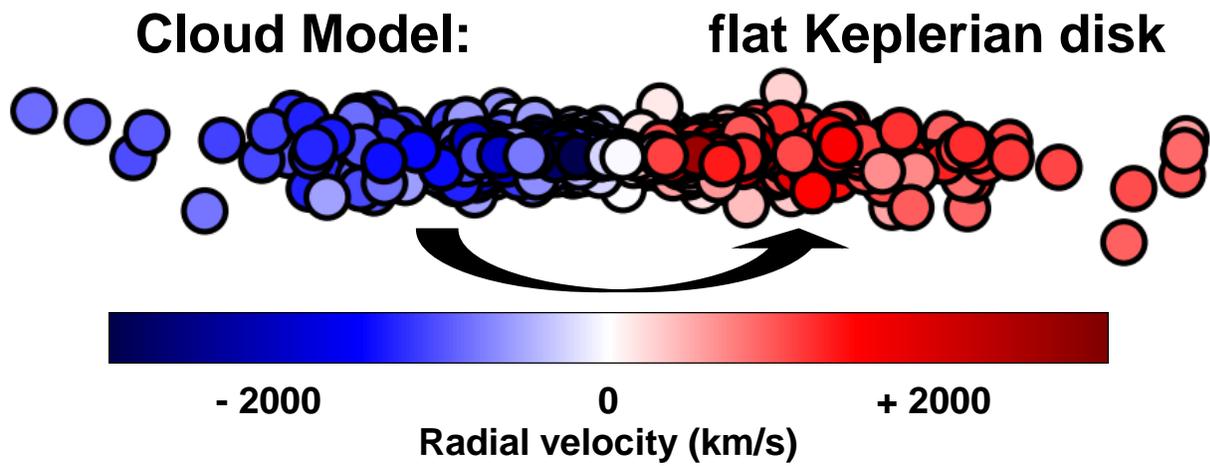
tens of micro-arcseconds
astrometric precision

baseline

centroid

$$\left[\frac{\vec{B}(\lambda)}{\lambda} \cdot \Delta\vec{x}(\lambda) \right]$$

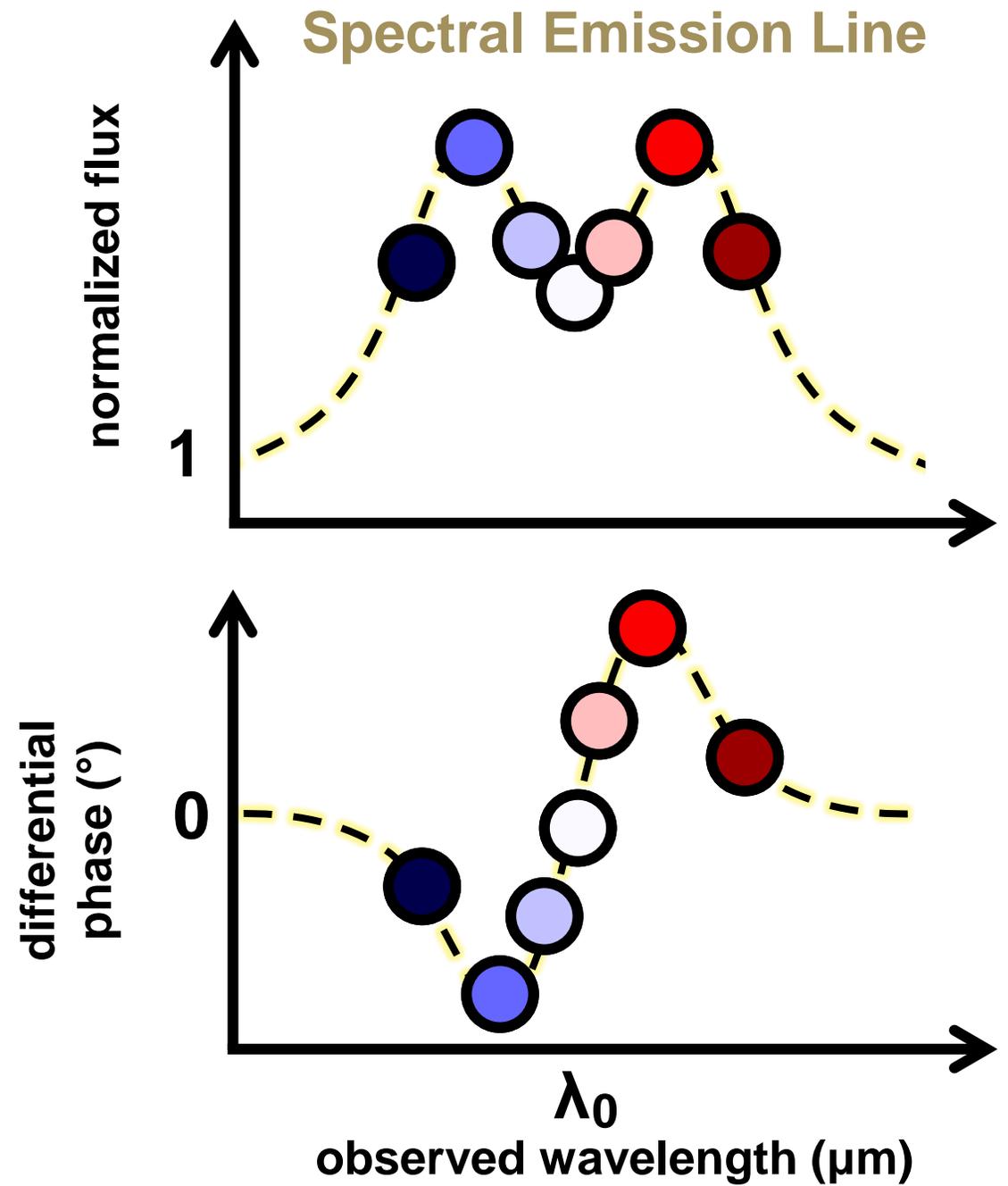
$\frac{100 \text{ m}}{2.2 \mu\text{m}} \cdot 100 \mu\text{as}$



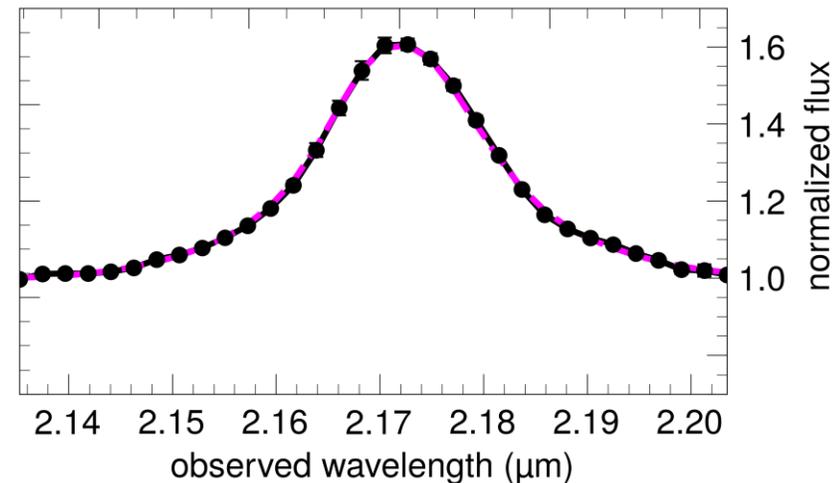
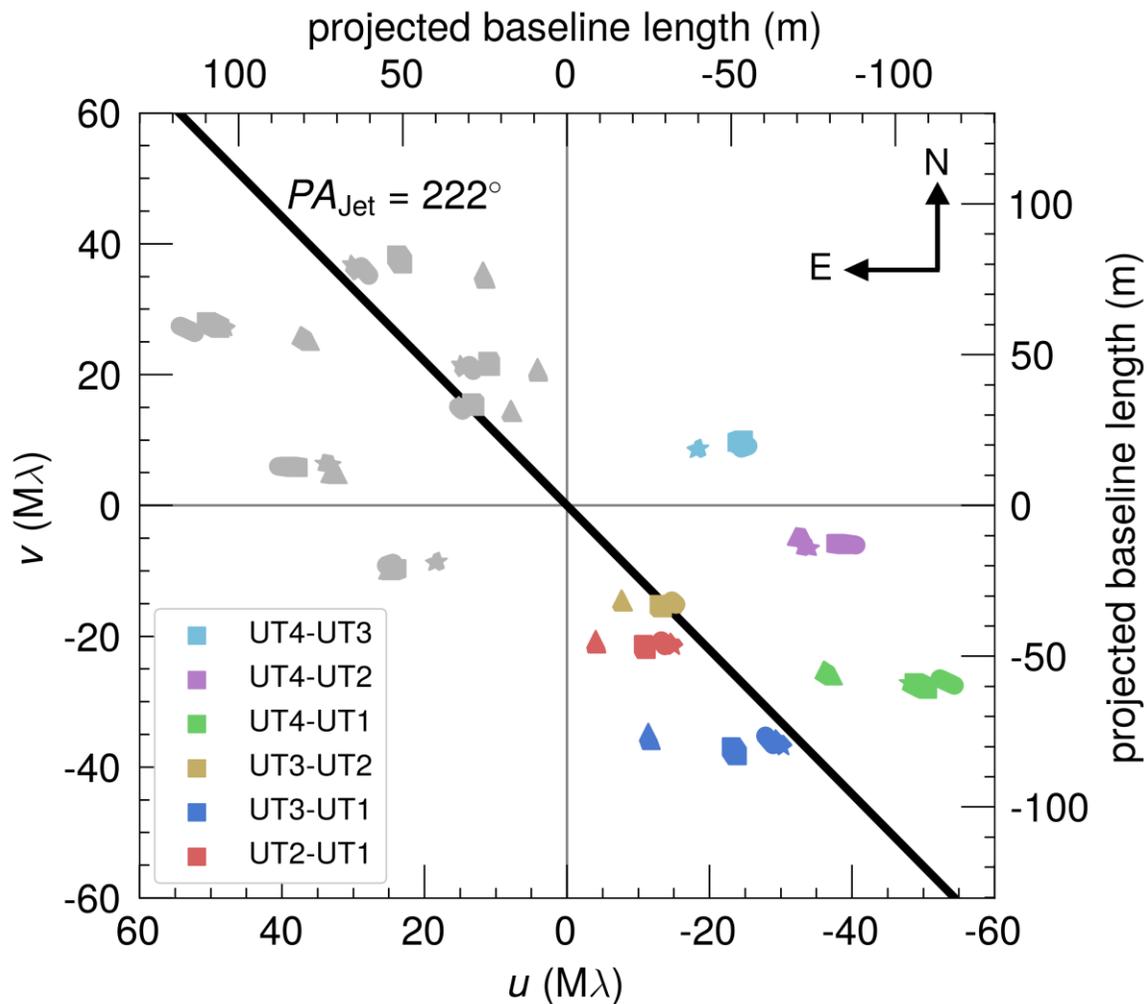
Line Intensity Map



S. Rakshit et al. 2015



Detection of the BLR of the Quasar 3C 273 with VLT/GRAVITY



line/continuum

baseline

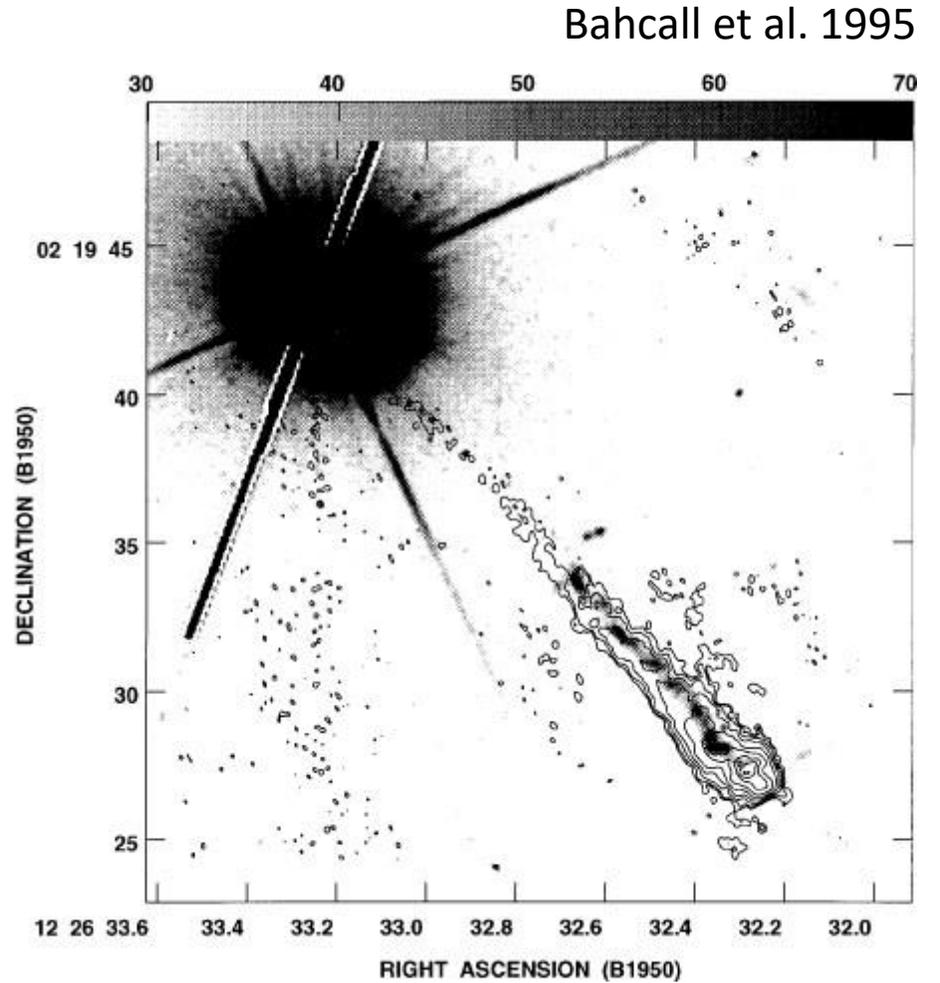
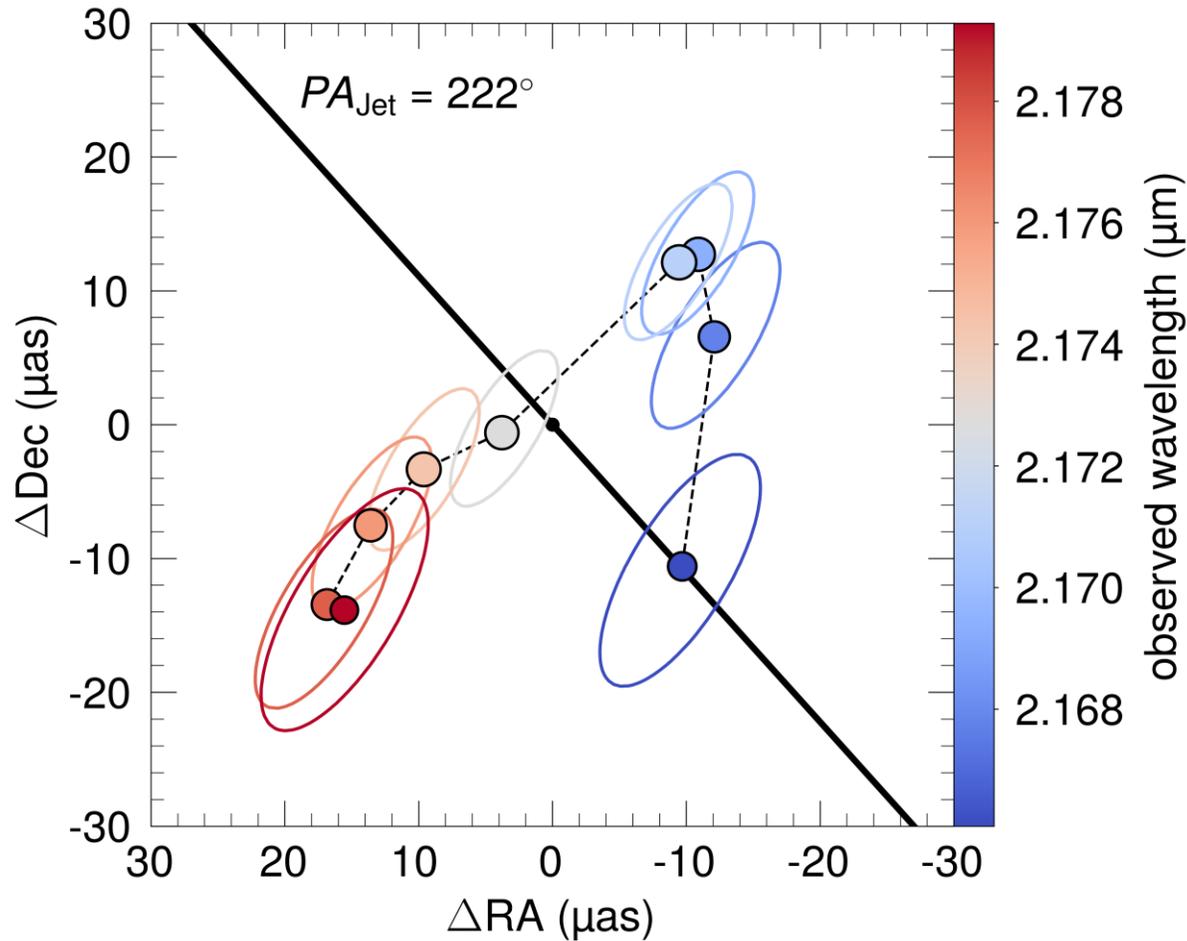
$$\phi_{\text{MC}}(\lambda, \vec{x}) = -2\pi \left(\frac{r(\lambda)}{1+r(\lambda)} \right) (\vec{u} \cdot \vec{x})$$

centroid

sum over all baselines

$$\chi^2(\lambda, \vec{x}) = \sum_{i=1}^{\text{Bl}} \left(\frac{\phi_{\text{GR}}(\lambda) - \phi_{\text{MC}}(\lambda, \vec{x})}{\Delta\phi_{\text{GR}}(\lambda)} \right)^2$$

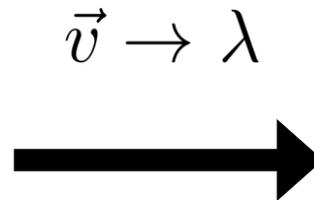
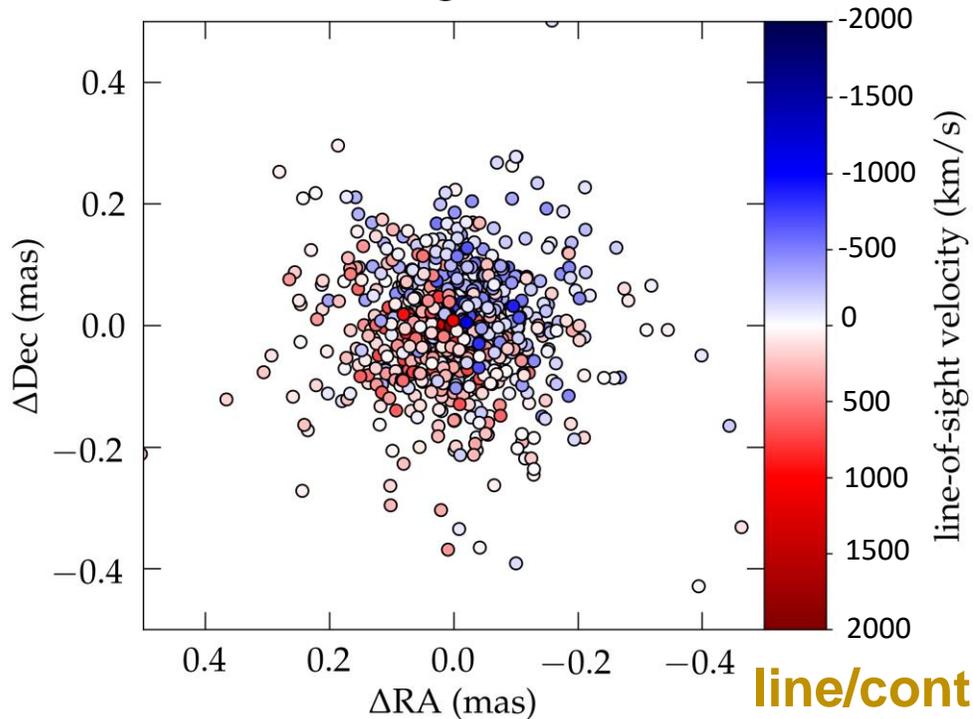
Ordered Rotation in the BLR of the Quasar 3C 273



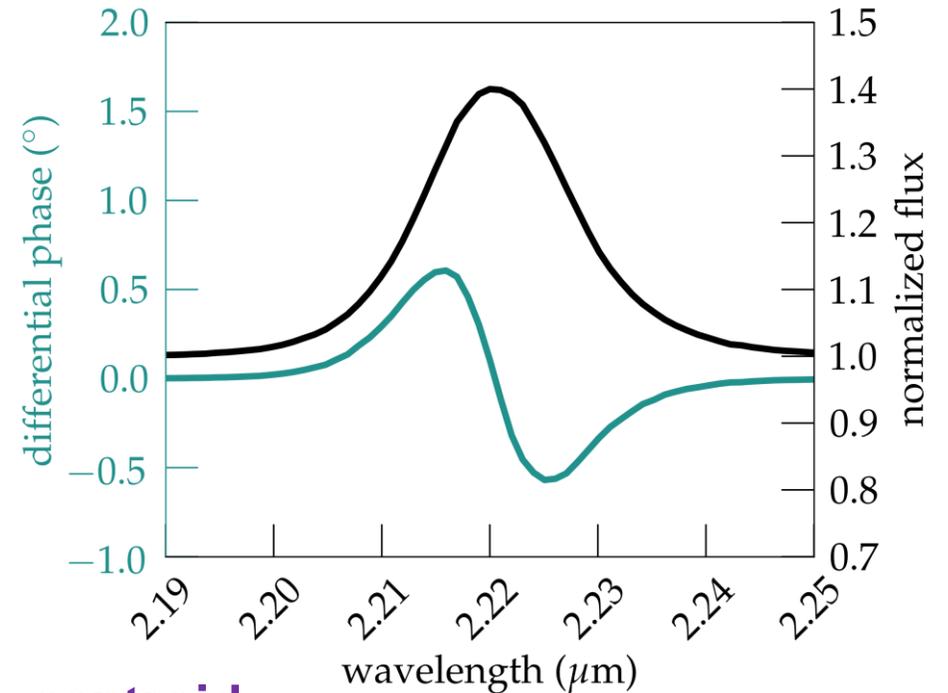
Phenomenological Model of the BLR

- following:
- A. Pancoast et al. 2014: reverberation mapping
 - S. Rakshit et al. 2015: interferometric observables

optically thin emission
from orbiting BLR clouds



wavelength bins



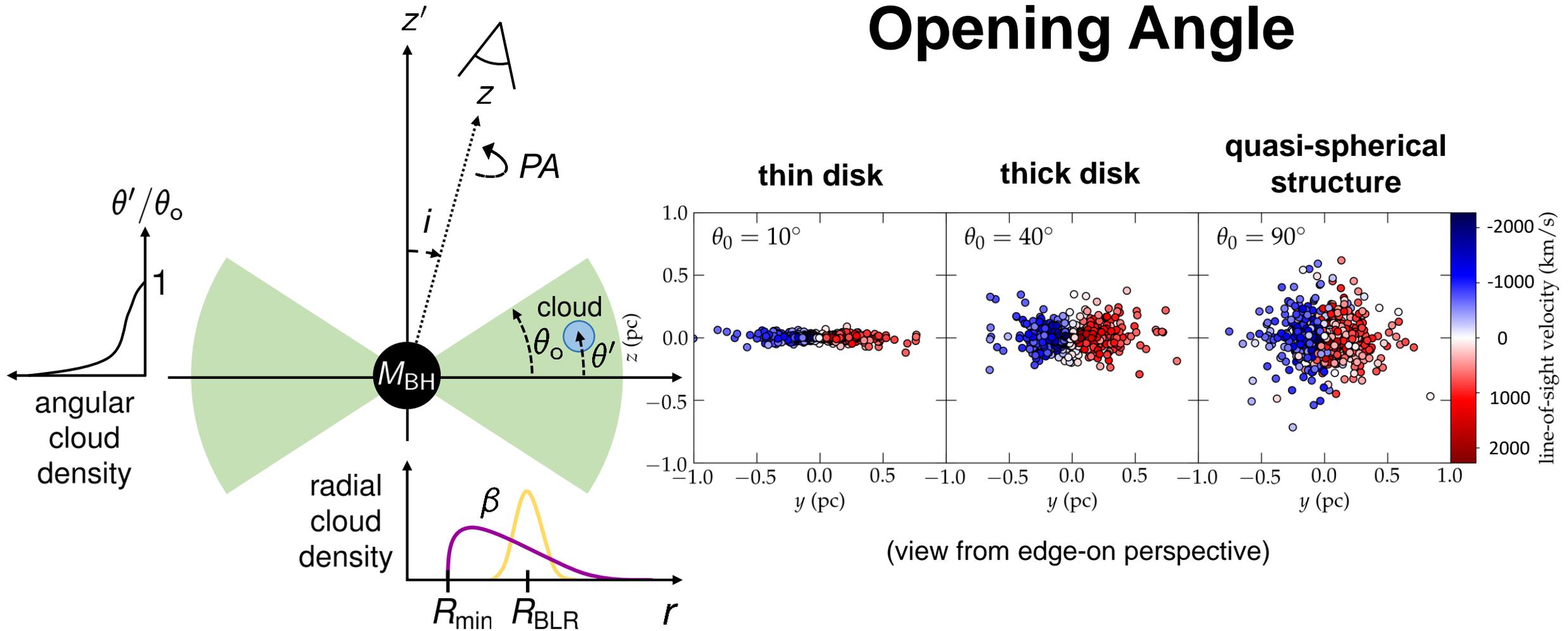
line/continuum
emission

baseline

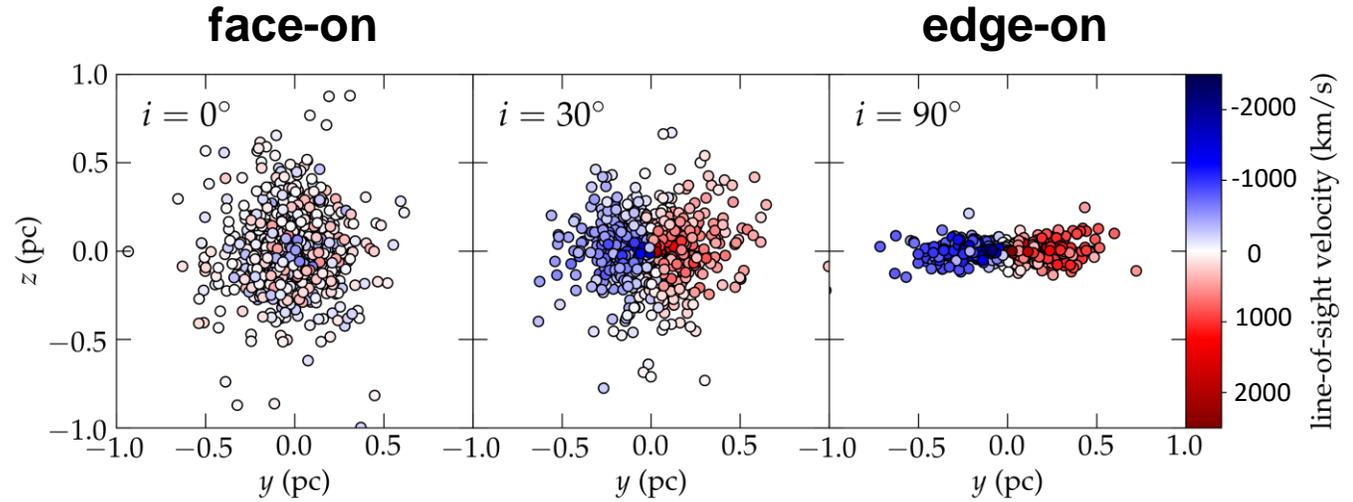
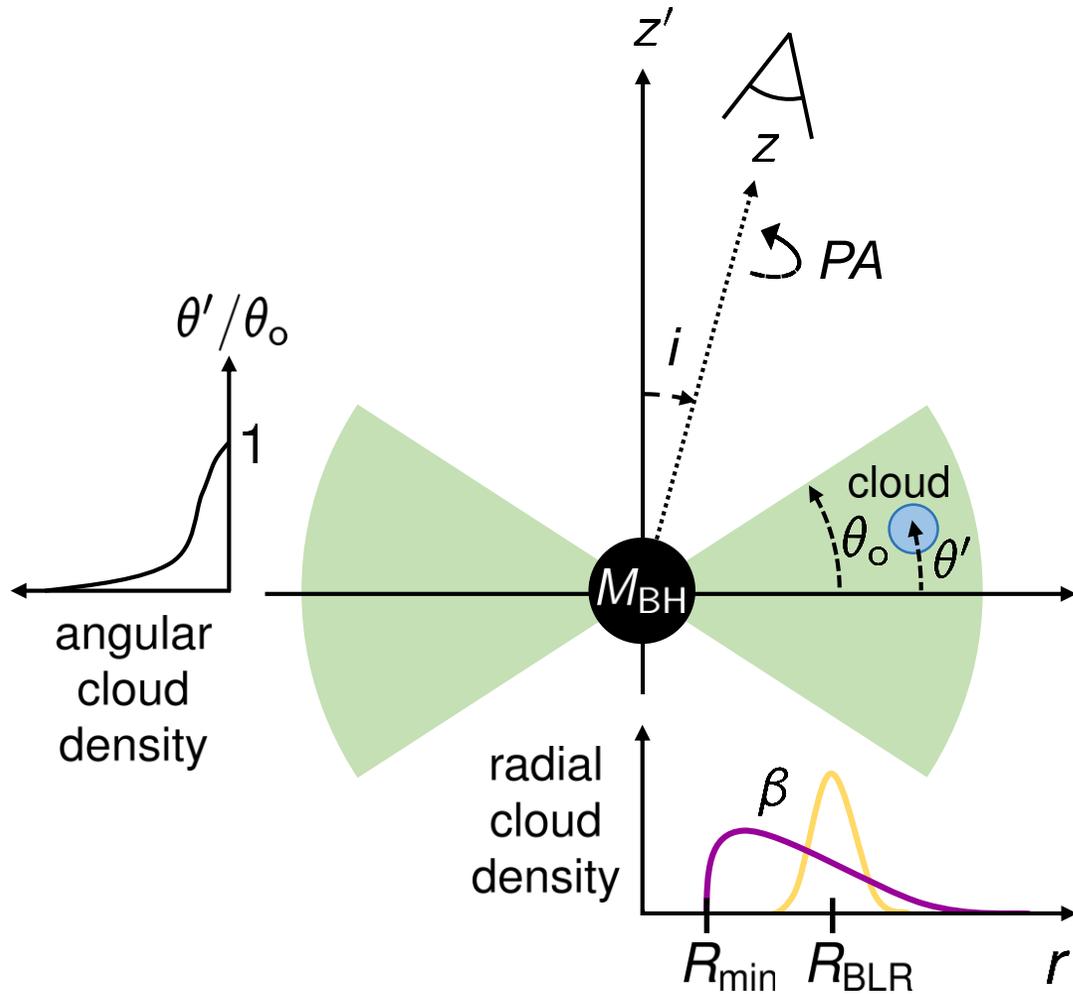
centroid

$$\Delta\phi(\lambda) = -2\pi \frac{f(\lambda)}{1+f(\lambda)} \left[\vec{u} \cdot \frac{\sum W_j \vec{x}_j}{\sum W_j} \right]$$

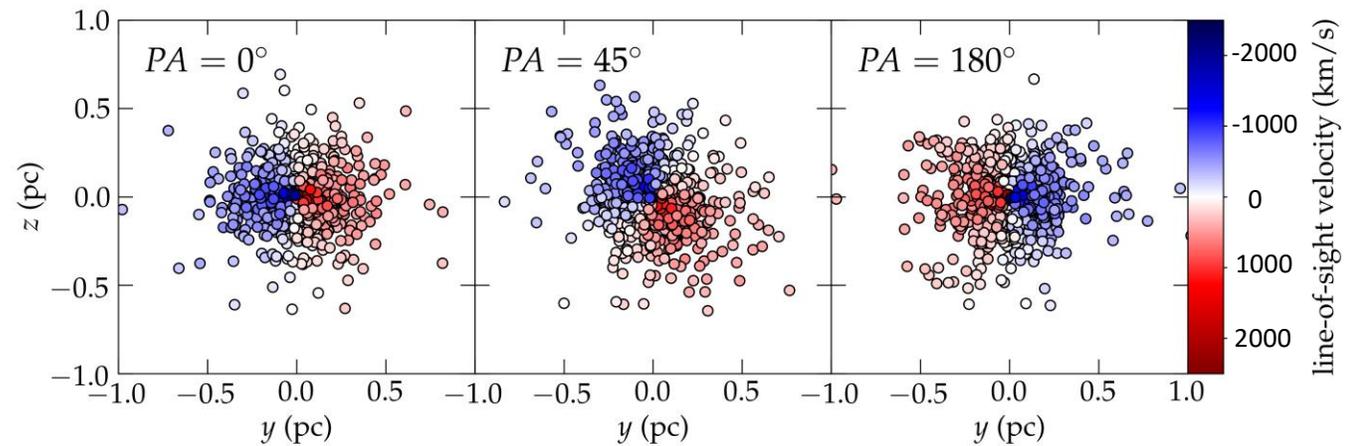
Opening Angle



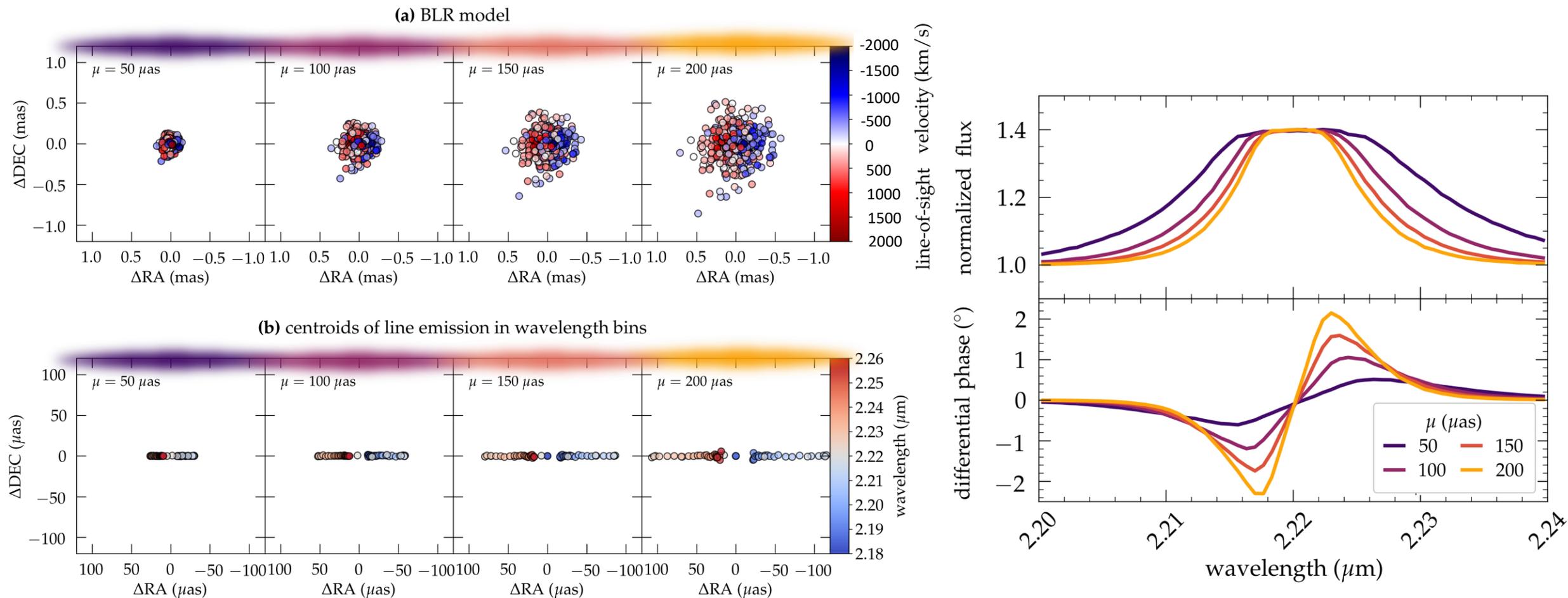
Inclination Angle



Position Angle

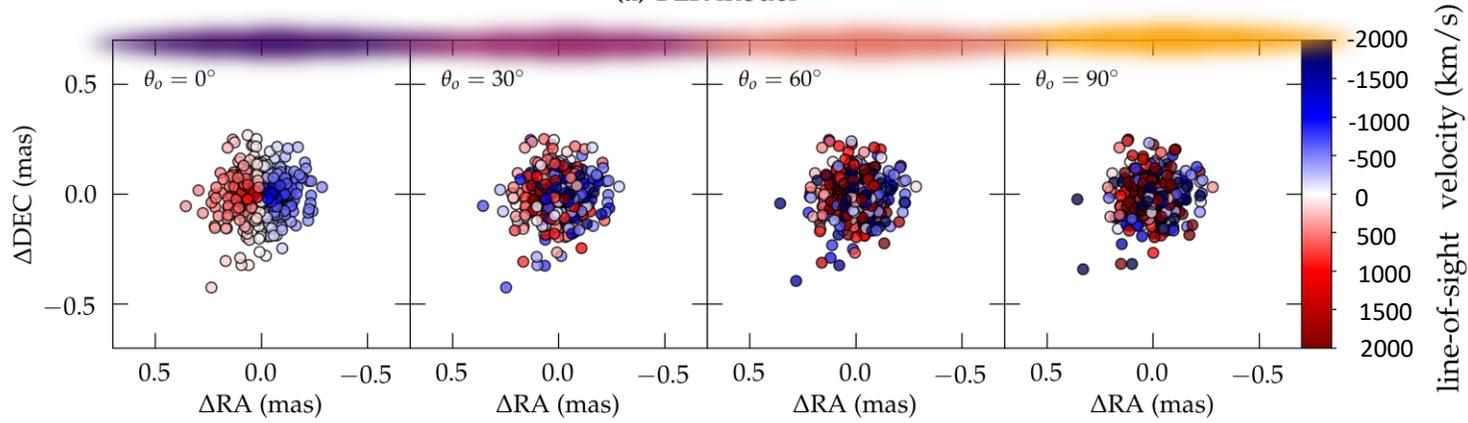


Mean Radius of the BLR

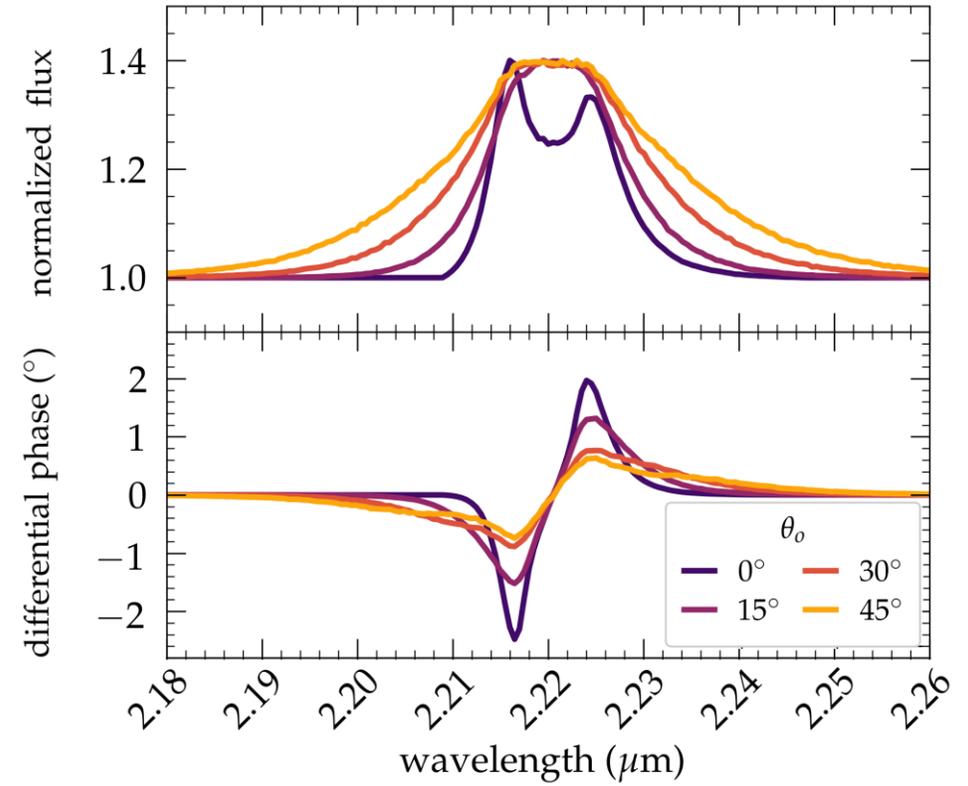
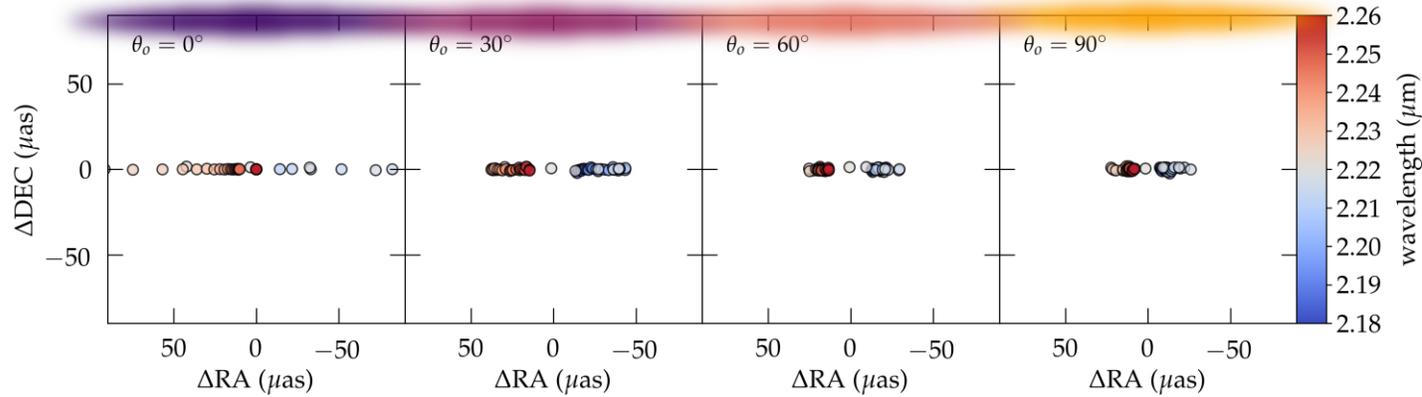


Opening Angle

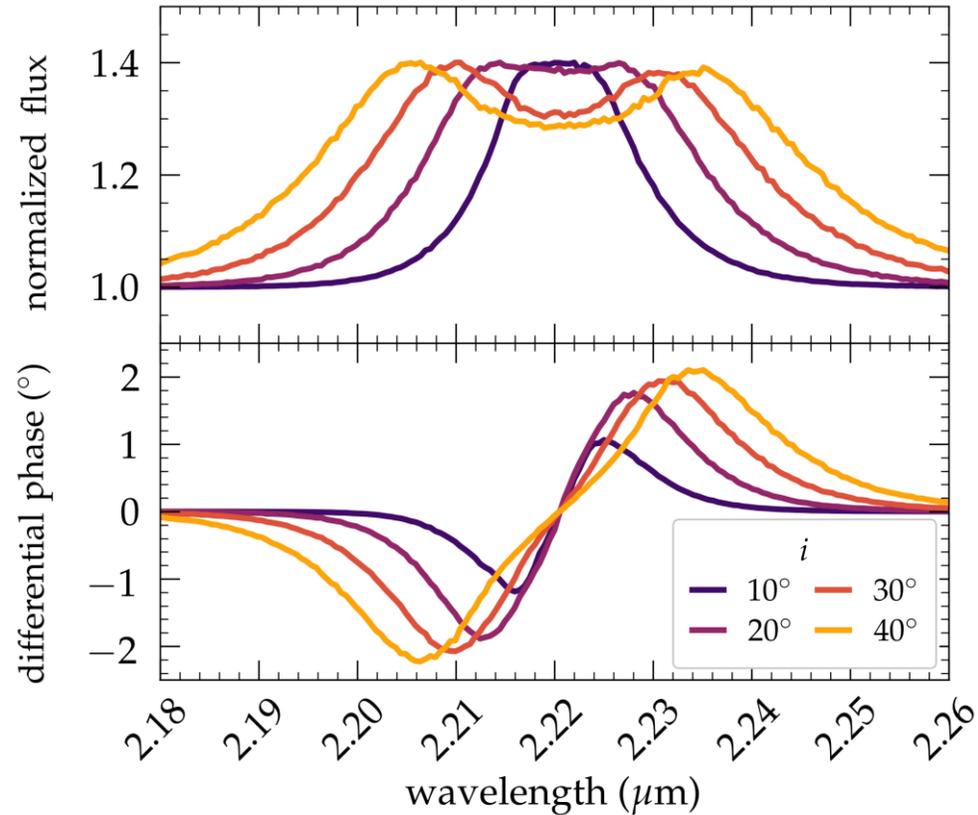
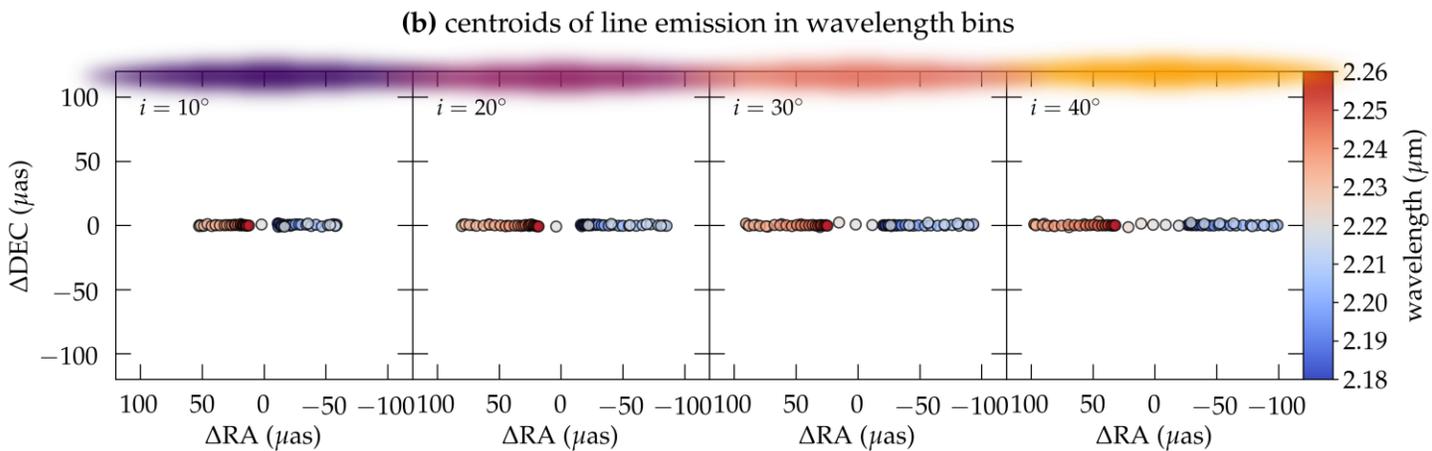
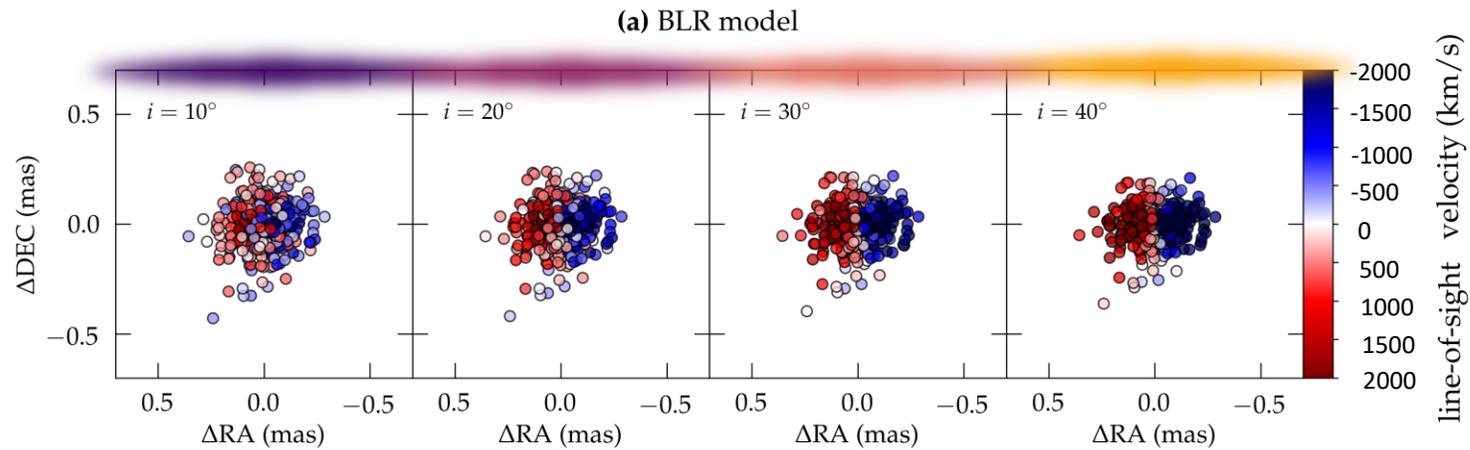
(a) BLR model



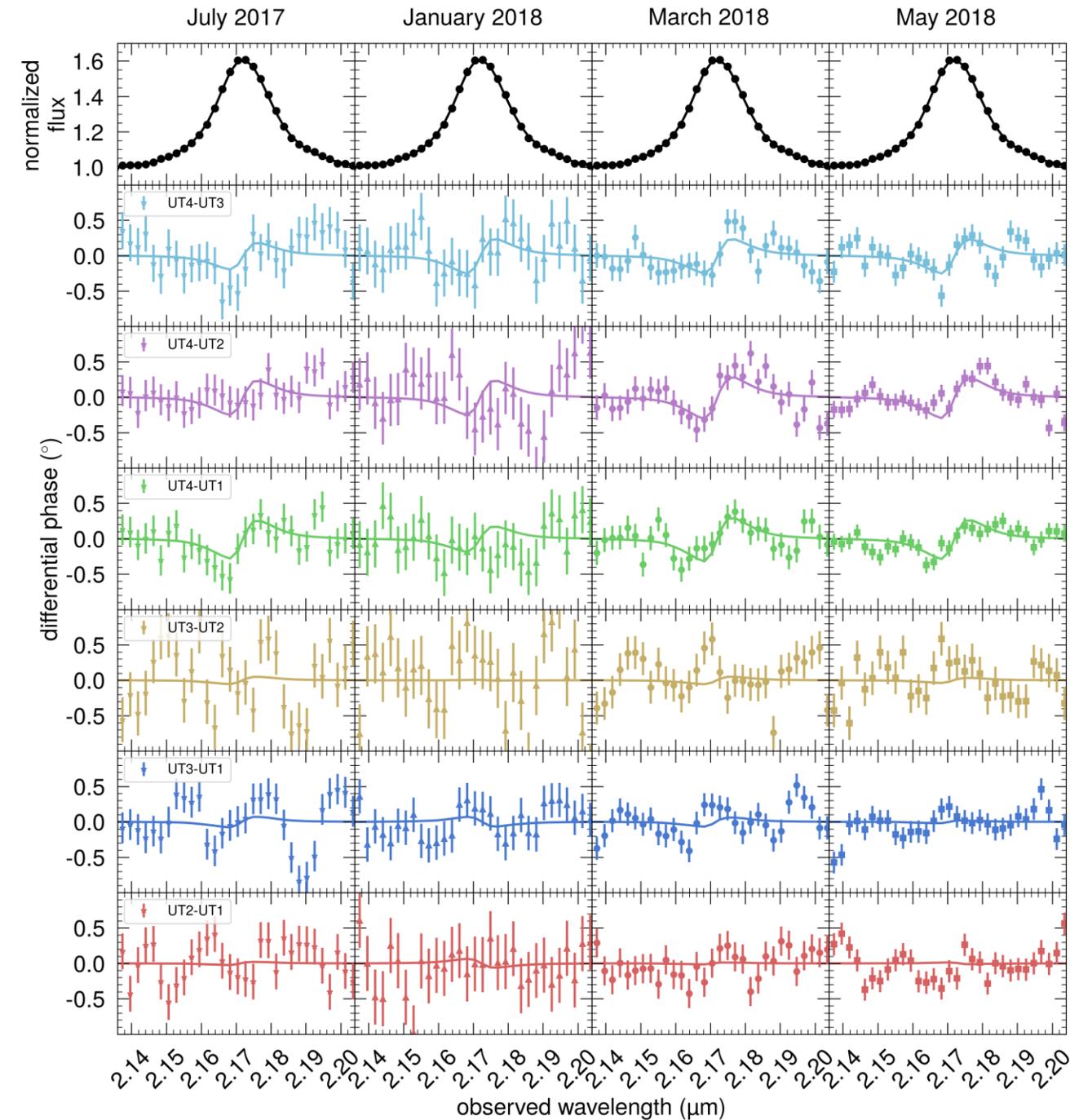
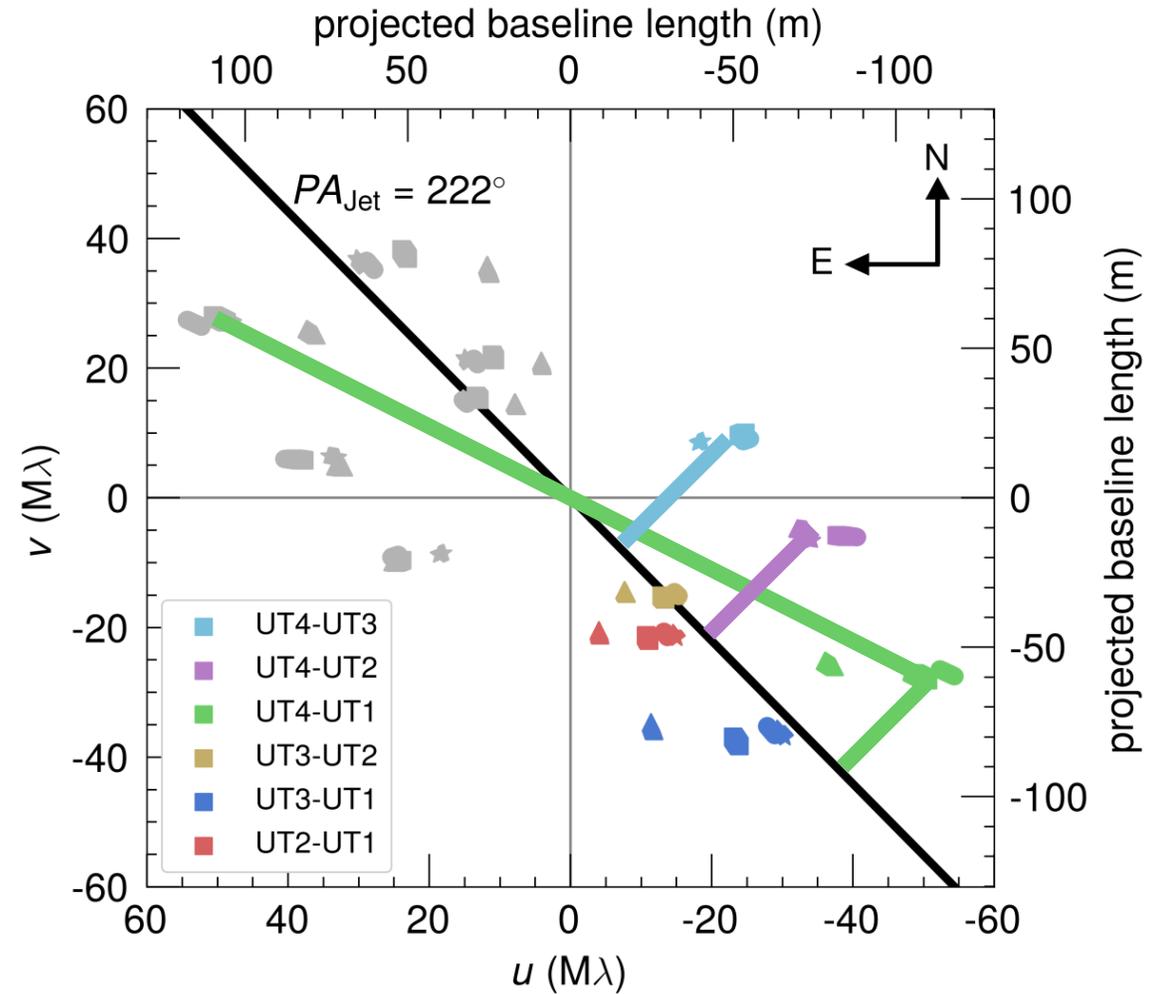
(b) centroids of line emission in wavelength bins



Inclination Angle

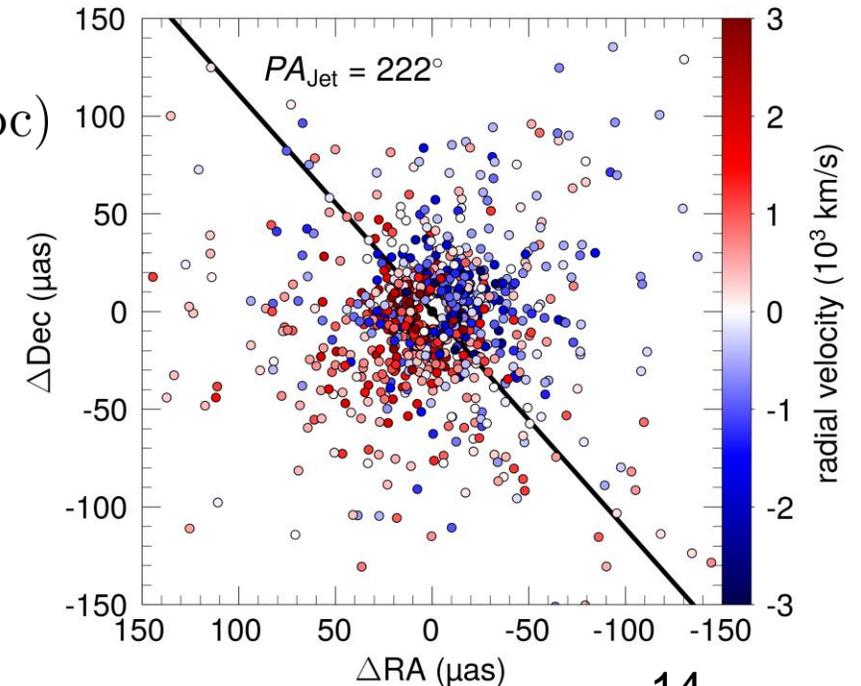
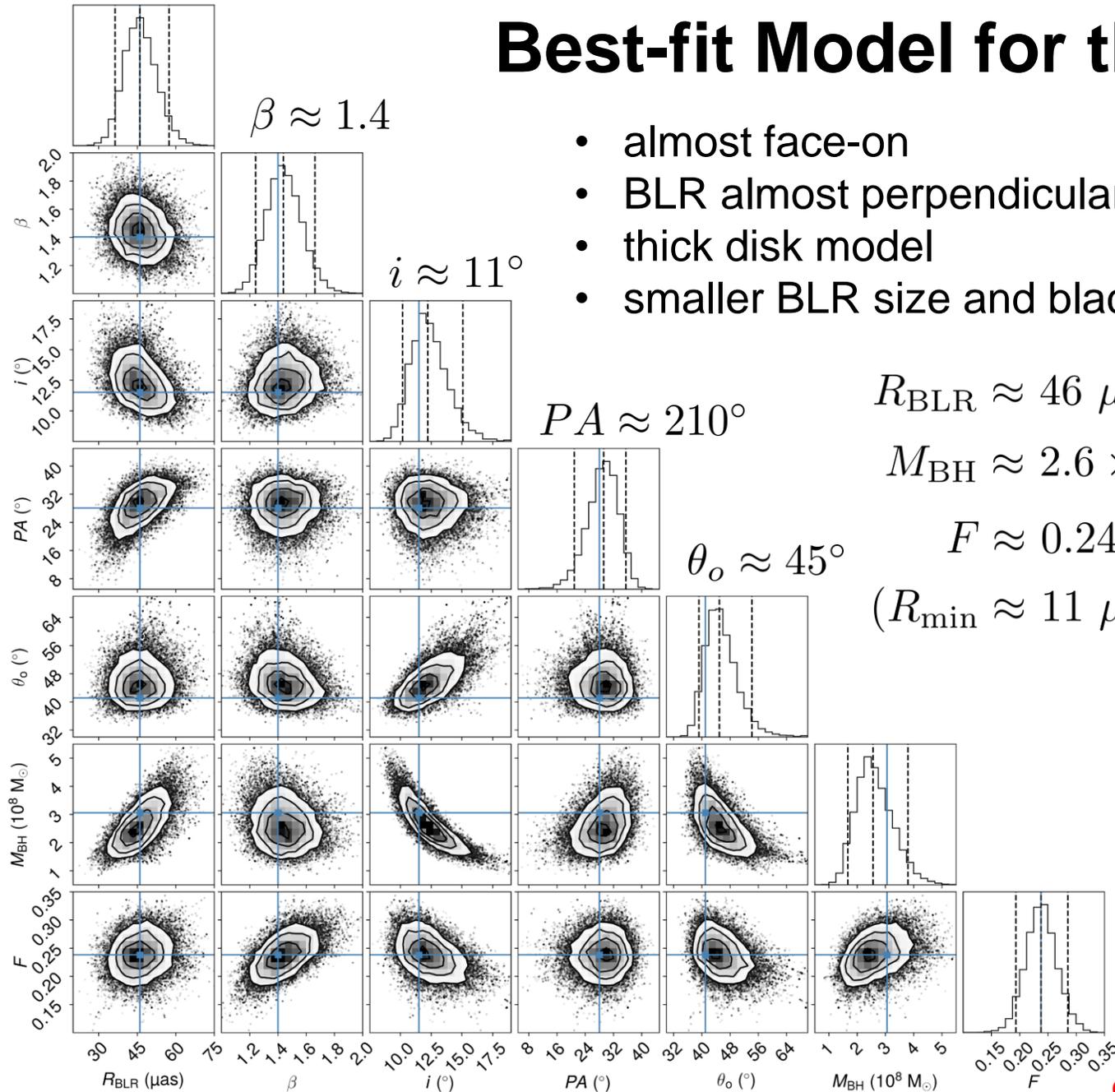


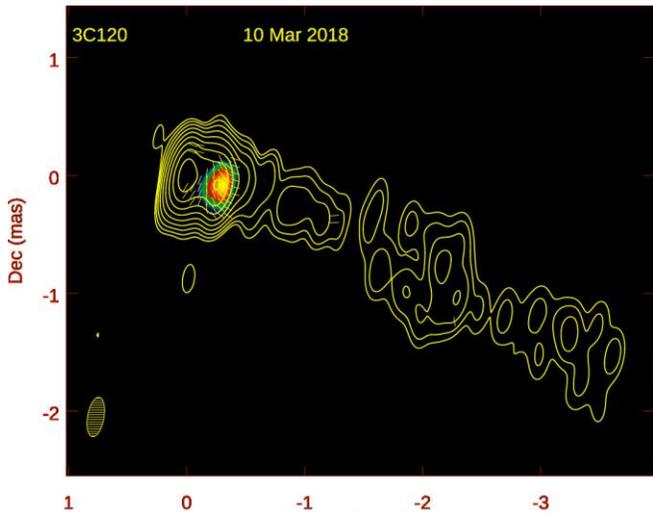
Best-fit Model for the BLR of Quasar 3C 273



Best-fit Model for the BLR of Quasar 3C 273

- almost face-on
- BLR almost perpendicular aligned to radio jet
- thick disk model
- smaller BLR size and black hole mass than from reverberation mapping





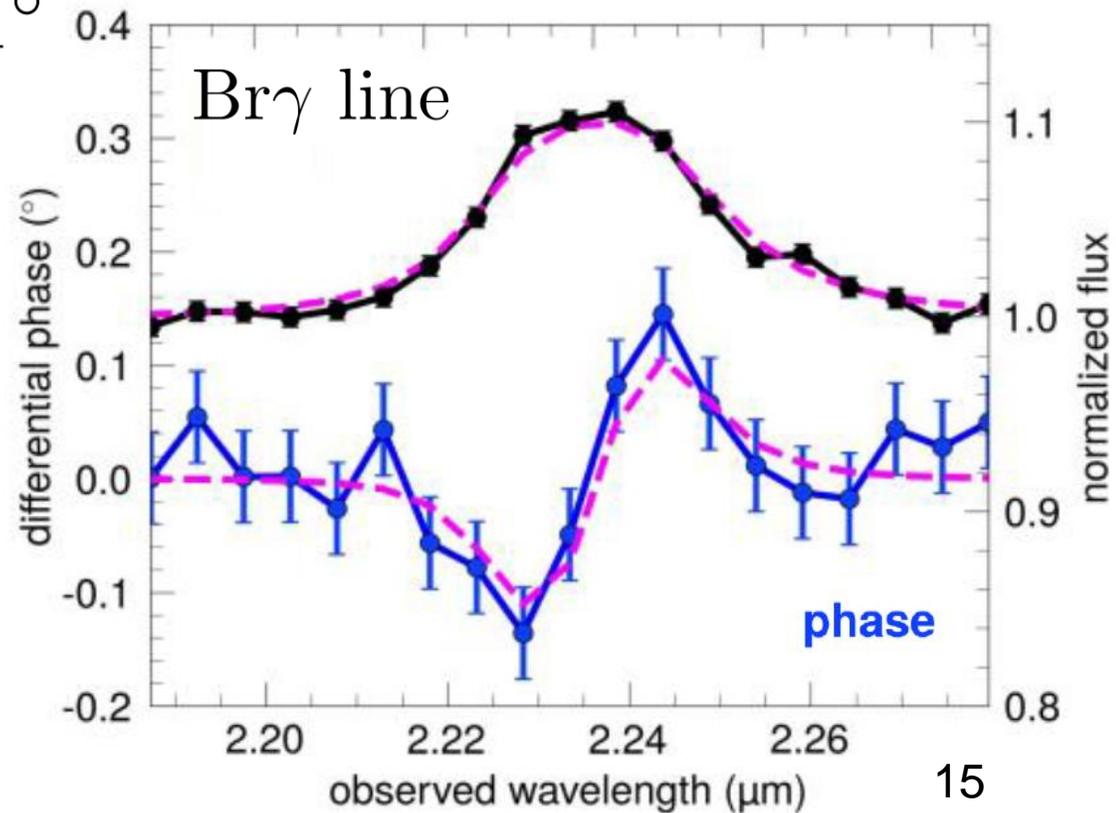
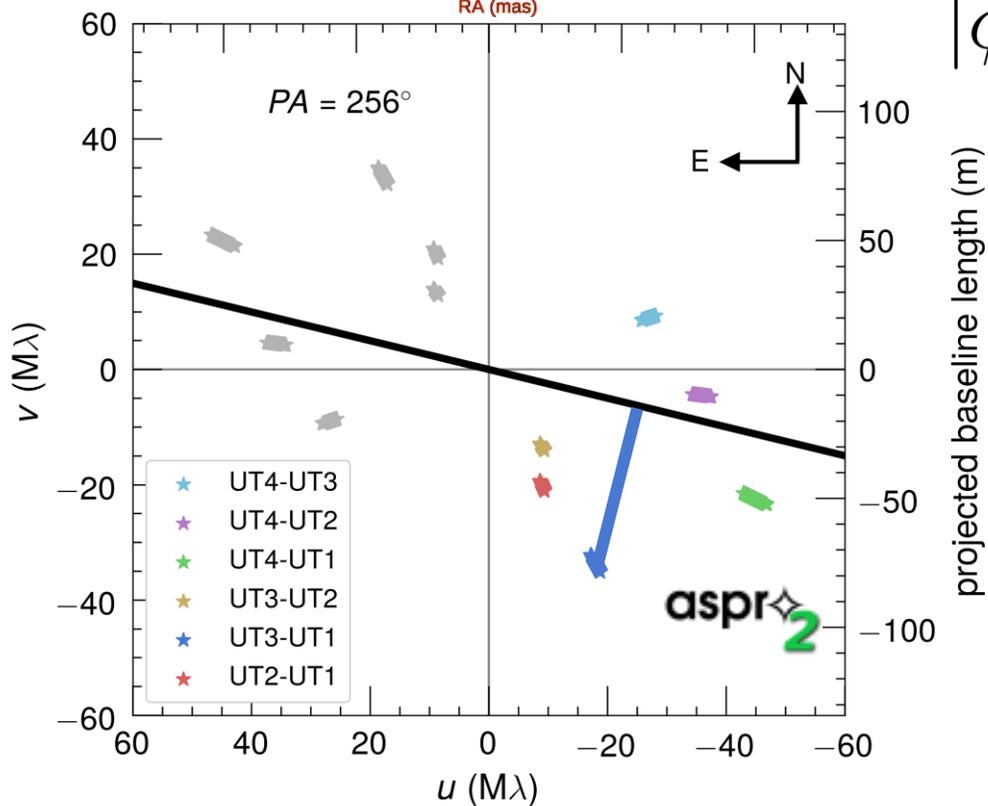
Predicting Phase Signal for 3C 120

- Interesting target to compare results from reverberation mapping

$$PA_{BLR} = PA_{Jet}$$

- Using similar model from Pancoast+ 2014 and reverberation size

$$|\phi| \lesssim 0.1^\circ$$



Summary

- GRAVITY detected the BLR of a quasar
- Flexible BLR model to predict emission lines & differential phases for observations

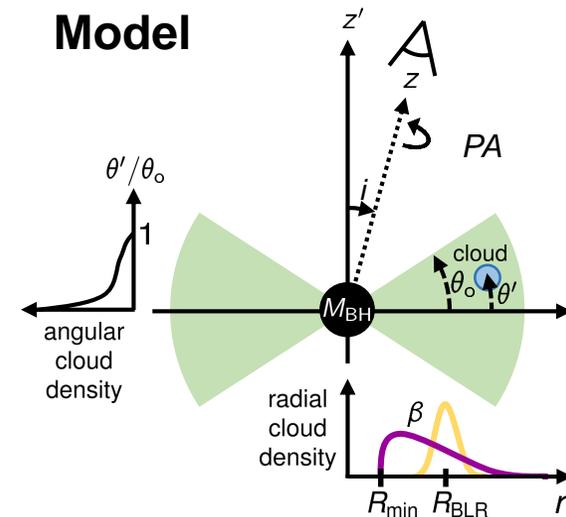
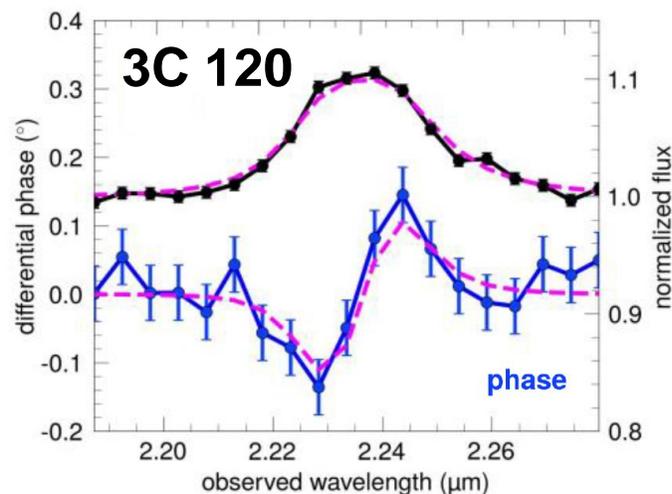
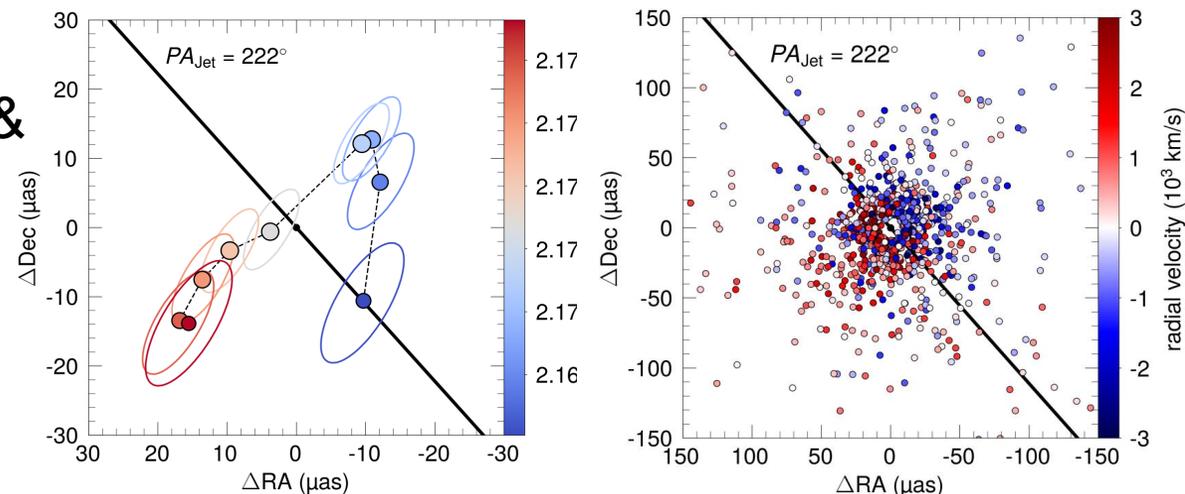
Direct Modeling of the BLR of 3C 273:

- Small differential phases due to: almost face-on inclined, thick disk structure perpendicular aligned to the jet/ unfortunate uv-coverage
- BLR size and black hole mass are smaller than from reverberation estimates

Outlook:

- Test for anisotropic emission etc...
- Developing disk-wind models

3C 273



**Thanks for your
attention!**